

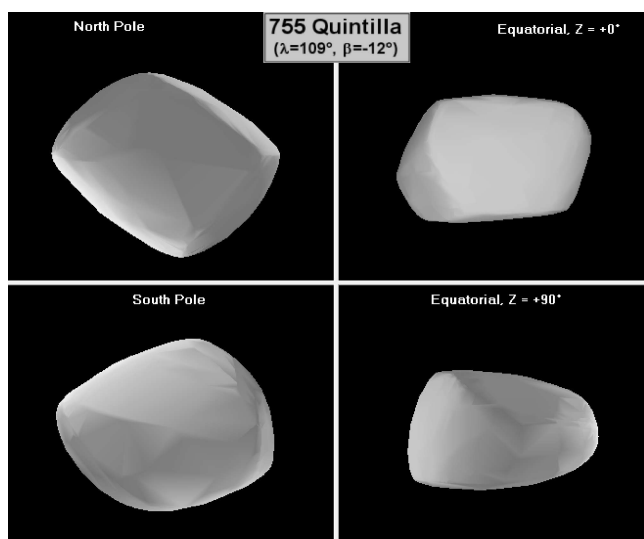


Figure 5: The shape model for 755 Quintilla ($\lambda = 109^\circ$, $\beta = -12^\circ$).

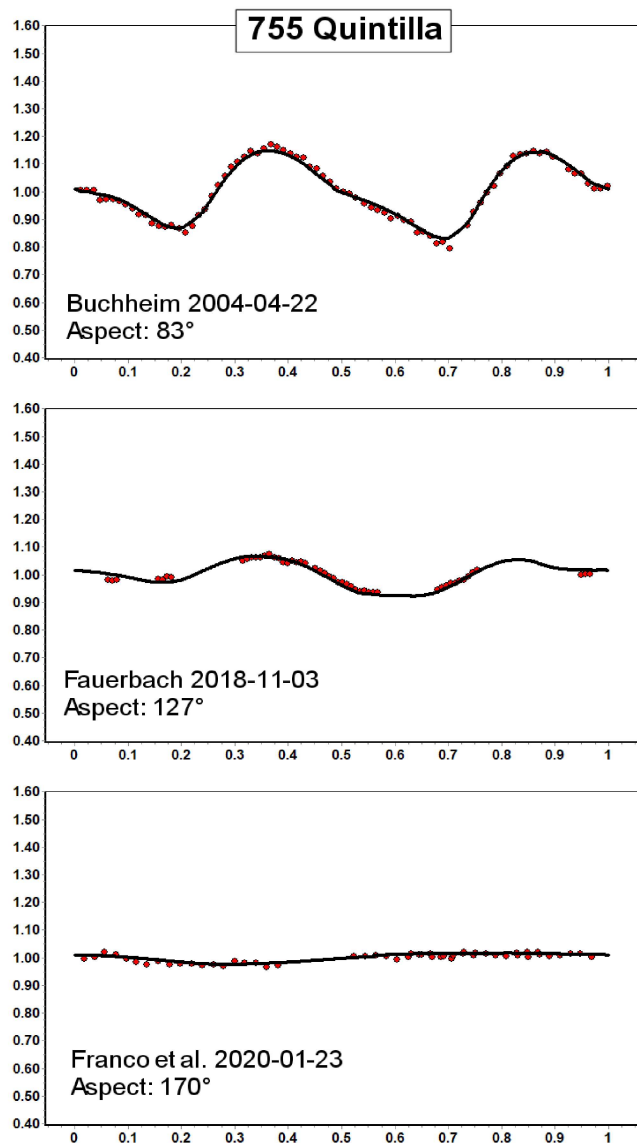


Figure 6: Model fit (black line) versus observed lightcurves (red points) for ($\lambda = 109^\circ$, $\beta = -12^\circ$) solution.

The analysis did not identify a unique solution (Durech et al., 2009), so we consider this to be a preliminary solution. Indeed, the pole search distribution is poorly constrained, especially along the ecliptic latitude. However, a check of the other four probable sidereal periods produced similar solutions with higher Chi-Sq and RMS values. We invite more observations of 755 Quintilla during the future oppositions, especially at large phase angles, in order to find a more robust solution.

References

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CORRIGENDUM

Franco, L.; Marchini, A.; Saya, L.-F.; Galli, G.; Baj, G.; Ruocco, N.; Mannucci, M.; Montigiani, N.; Tinelli, L.; Scarfi, G.; Aceti, P.; Banfi, M.; Bacci, P.; Maestripieri, M.; Papini, R.; Salvaggio F.; Mortari, F.; Bachini, M.; Casalnuovo, G.B.; Chinaglia, B. (2020). "Collaborative Asteroid Photometry from UAI: 2020 January - March." *Minor Planet Bulletin* **47**, 242-246.

Incorrect values were given for the lower limit of the secondary-to-primary mean diameter ratio D_s/D_p for the asteroids 1052 Belgica and 7132 Casulli. The correct values are: 1052 Belgica, 0.39 ± 0.02 ; 7132 Casulli, 0.33 ± 0.02 .