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Intensive photometry of southern Be variables. I. Winter objects (*)

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Summary. — We present results of an intensive photometric campaign on some bright southern Be stars to search for periodic light variations. In order to obtain good phase coverage, observations were conducted from two sites with different longitude : ESO and SAAO. Most of the stars observed are indeed variable with periods close to one day (the expected rotational period for these stars). We present our results for winter objects.

Key words : Be stars — variable stars — photometry — stellar pulsations.

1. Introduction.

Periodic variability in Be stars has attracted great interest in recent years. The discovery of unexpected short-period variations in these stars has renewed the search for a viable mechanism of mass ejection. The current consensus of opinion is that non-radial oscillations, driven by an unknown mechanism, is the likely cause of the mass loss leading to the Be phenomenon (see Percy, 1987, for a recent review). This conclusion is mostly based on line-profile variations seen in some Be stars.

The small-scale “moving bumps” seen in the line profiles of some Be stars is attributed to high-order non-radial pulsations (NRP). But moving bumps are also seen in B stars without emission lines and even in δ Scuti stars (Walker and Yang, 1987). It seems that NRP is not solely responsible for the mass ejection. It is also known that one of the distinguishing features of line-profile variations in Be stars is the presence of low-order line-profile variations which some attribute to low-order NRP (Penrod, 1987). But again, low-order NRP is present in non-Be stars as well (β Cep and δ Sct). Many β Cep variables are rotating more rapidly than some Be stars. High rotation coupled with NRP does not seem to be a sufficient condition for the Be phenomenon. The information from photometric observations has not been adequately considered and it is possible that the conclusion that NRP is responsible for the mass loss could be premature.

A viable hypothesis should be able to explain the distribution of observed periods, the fraction of Be stars which are periodic, their photometric amplitudes and evolution of the shapes of their light curves. Such information can only be obtained by intensive photometric monitoring of a large sample of Be stars. Owing to the fact that periods are close to one day, it is very important to obtain observations from different longitudes. These observations should be obtained over a short time as the light curve can evolve quite rapidly. These demands can only be met by an international campaign.

In this paper we present results of such a campaign conducted from ESO and SAAO as well as some observations obtained at SAAO only. We observed 17 bright southern Be stars ; most of them turned out to be short-period variables. These stars were chosen from the *Bright Star* catalogue on the sole criterion that they should be well placed for observing. A list of candidate and comparison stars is given in table I. Results for variable program stars are given in tables IIa and IIb.

2. Observations and reductions.

The observations at SAAO were obtained with the Volks photometer attached to the 0.5-m reflector at Sutherland. These were made mostly through a single filter, Strömgren *b*, which was chosen as a compromise between the expected larger amplitude towards shorter wavelength and the

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decreased sensitivity to atmospheric extinction at longer wavelengths. Some four-colour observations were also made. For the brightest stars, a neutral density filter had to be used to prevent over-illumination of the photocathode. In these circumstances we calibrated the ND filter by observing the comparison stars with and without the filter. The rms deviation suggests that this calibration is good to within 3 millimag (the expected observational error of one measurement). Systematic errors due to the use of the neutral density filters are probably negligible owing to the narrow bandpass of Strömberg *b*.

In the reduction procedure, mean extinction coefficients were obtained for each night using the comparison stars. Corrections for transparency changes during the night were made by calculating the deviation for each comparison star. Stars in close proximity were treated as a group and mean transparency corrections were applied to the group, which included the Be stars. Comparison stars which were variable were detected at this stage and could be eliminated. We searched for micro-variability in the comparison stars by calculating their periodograms; again such stars could be identified and eliminated. Finally, the mean magnitude of all comparison stars was adjusted to be in exact agreement with the mean *b*-magnitude for the same stars observed at ESO. This zero-point correction ensured that the Be stars observed from the two sites were on the same system.

The observations at ESO were made between 13 June and 6 July 1987. They were obtained with the six-channel *uvbyH β* photometer on the Danish 0.5-m telescope. This instrument uses a grating and slots to define the passbands; in consequence some systematic differences could be expected between the *b* magnitudes of the two observatories. In practice, these differences were very small (a few millimag) and posed no problem when combining the data sets. Because a new telescope control system with auto-centering and automated observations had just been installed and still had to be checked during the observing run, the quality and quantity of the observations was not as high as could be expected at the start of the run. Later on, most of the problems were solved and nowadays the new, automated system has proved to be very reliable.

For the brightest stars, the same procedure was used as at SAAO. The reduction procedure was similar except that a polynomial model for transparency changes and zero-point drift was calculated for each night. Furthermore, a general least-squares solution which included all observations for all comparison stars provided best estimates for their mean magnitudes. Variable stars among them were easily identified by the large standard deviations. They were not included in obtaining the final magnitudes for the comparison stars. A small zero-point difference between these results and the standard *uvby* system is still possible, but this is unimportant for the study of their light variations. The rms scatter for the comparison stars is in all cases less than 5 millimag.

3. Period finding.

We employed two methods of searching for periodicities. One is the standard Fourier periodogram analysis technique for unequally spaced data; the other is a modification of Stellingwerf's (1978) phase-dispersion minimization (PDM) technique. The first method is suited to the case where the underlying variation is approximately sinusoidal, but it can give misleading results when the variations are highly non-sinusoidal or when a multiple-wave light curve is present. It is known that many Be stars have a double-wave light curve which will not be detected by this method. Clearly it will be impossible to discriminate between a single-wave light curve and a double-wave light curve with nearly equal amplitudes.

The PDM technique will identify the period for all kinds of light curves, but since the general noise level is rather high in this method (Swingler, 1989), one has to be careful of spurious periodicities. Our final choice was based on a careful examination of the phase diagram resulting from the periods produced by the Fourier and PDM techniques. Our criterion was by nature rather subjective, but we attempted to choose the period which resulted in the curve with the least scatter. Except for one or two cases, there is little doubt that our choice is the correct one.

Besides the complication discussed above, a serious difficulty arises when the irregular light variations which are sometimes present in Be stars (attributed to effects of the circumstellar material) has a time scale comparable to the underlying periodic variation. In these cases it is not possible to obtain the correct period with any certainty. As some long-term activity is nearly always present, we normally removed the long-term trend by fitting a low-order polynomial to the data. This of course has the effect of reducing the scatter at the expense of some uncertainty in the shape and amplitude of the light curve. The derived period will in general be unaffected by this procedure.

It is difficult to derive reliable standard errors for the periods, but an estimate can be obtained from the half-width of the Fourier periodogram peaks. Typically, we observed for about four weeks in each season which leads to a standard deviation of about 0.02 d^{-1} in frequency which can be taken as a reliable guide to the error in the quoted frequencies. Sometimes we have combined many seasons in order to obtain a more accurate period. This is always possible to do, but invariably there will be many possible choices of frequency (each separated by approximately one cycle per year) which will produce as good a fit. In this case the real error is not reduced by combining the data.

4. The results.

ν Cen

This star is a single-line spectroscopic binary (Palmer, 1906). Wilson (1914) determined the orbital elements and obtained a circular orbit with $P = 2.62516 \text{ d}$. Hendry and Bahng (1981) found $H\alpha$ to have double emission

surrounding an absorption core. This observation indicates that ν Cen could be classified as a Be star, but it is probable that the emission originates as a result of binary interaction rather than an intrinsic property of the star itself. Spectroscopic observations suggest that this star may be a β Cep variable with a period of about 0.17 d and an amplitude of 5 km s^{-1} (Rajamohan, 1977 ; Kubiak and Seggewiss, 1982 ; Ashoka *et al.*, 1985). Waelkens and Rufener (1983) do not find any evidence for short period light variability, except for a variation with the same period as the binary period.

Our photometric observations (SAAO only) extend over two seasons, 1987 and 1988. The 1988 observations are particularly numerous and clearly show a variation with the same period as the orbital period (Fig. 1). Minimum light corresponds to maximum positive radial velocity. As discussed by Waelkens and Rufener (1983), the resulting light curve is most probably a reflection effect. A careful investigation did not reveal any signs of a variation at any other frequency with an amplitude exceeding 2 millimag.

μ Cen

This star is known as a “pole-on” Be star (v in $i = 155 \text{ km s}^{-1}$). Ghosh *et al.* (1987) have summarized the history of $H\alpha$ emission ; they reported an outburst in March/April 1987, about three months before the first set of our observations. A full discussion of an earlier outburst (February 1987) is given in Baade *et al.* (1988). Thimm (private communication) reported from spectra taken on May 29, 30 and on June 1, 1987 that $H\alpha$ changed from pure absorption to emission in 3 days. This was only a few days before the start of our observations. All indications are that this Be Star was in a very active phase during this time.

μ Cen was one of the first Be stars in which periodic line-profile variability was found (Baade, 1984). Baade observed three or four nearly equally-spaced moving bumps. There were also low-order changes in the line profile in which the symmetry of the whole line varied, but with a longer time scale. The period of the low-order variations was estimated to be 0.505 d, five times the period of the moving bumps. Baade interpreted these results in terms of NRP with $\ell = m = 2$ and $\ell = m = 10$. In our opinion the poor time coverage places considerable doubt on the accuracy of this period which is practically equal to half a day, since the aliasing problem must be quite considerable. More recently, Baade (1987a) obtained 9 nights of line profile observations in which this period is apparently confirmed.

Since two kinds of line-profile variations are seen in this star, Baade (1987b) claims that it is multiperiodic and hence proof that the NRP interpretation is the only feasible one. Harmanec (1987a) has suggested that the short period structure in Be stars can be represented by integer ratios, in which case we are dealing not with incompatible periods but harmonics of a longer period (which he identifies as the period of rotation). Baade (1987b) has criticized this suggestion. Since the period ratio in μ Cen is 5 : 1 according

to Baade, this star is perhaps not a convincing case of multiperiodicity.

We observed μ Cen for 29 nights in June/July 1987, 25 nights in March/April 1988 and 7 nights in June 1988. The latter two data sets are SAAO observations only. The 1987 data show that the star was photometrically very active at this time. There are two major excursions in brightness, each with a light range of 0.07 mag, separated by five nights.

A Fourier periodogram analysis shows no obvious trace of coherent periodicity : the three largest peaks being at frequencies of 0.05, 0.95 and 0.16 d^{-1} . These first frequencies appear to be artifacts caused by the large brightness excursions. The PDM periodogram shows essentially the same structure except for a large peak at 0.48 d^{-1} . This does not seem to be associated with the random brightness fluctuation and is the only coherent period we can extract from our 1987 data. The resulting light curve is a double-wave with a period of 2.10 d (Fig. 2).

Since our results from 1987 did not lead to a convincing period determination, we re-observed the star from SAAO in 1988. This time the star was less active, but surprisingly the periodograms look very similar to those of 1987. In particular the frequency at 0.48 d^{-1} was very strongly present. This seems to indicate that the 2.10 d period is physically real. As a final check we combined the data from both seasons which allowed us to refine the period to 2.1017 d. A plot of the light curve is shown in figure 2. The double-wave nature is very clear, but the large scatter indicates that there is an additional source of variation. A large increase in amplitude followed by a swift decrease can explain some of the apparent scatter in 1987. The peak-to-peak amplitude decreased from 0.08 mag in 1987 to 0.04 mag in 1988, though the large amplitude in 1987 could be mostly a result of a non-periodic brightness excursion during a few nights. But this cannot be invoked as the cause of the large scatter in 1988 ; this phenomenon is very common in the light curves of Be stars and can be termed “flickering”.

In the light of these results, we cannot confirm Baade’s 0.5 d period (frequency 2.0 d^{-1}) in μ Cen. We can possibly reconcile our results with this period if we note that the single-wave frequency (0.96 d^{-1}) is the one-day alias of Baade’s period. We certainly did not find any indication of a 0.5 d period in our data. As a test, we analyzed Baade’s (1984) radial velocity observations for the HeI line and constructed the phase diagram using a period of 2.1017 d. The resulting velocity curve has considerably greater scatter than the one constructed with 0.505 d, but is still quite reasonable and gives a double-wave curve. However, there are only 19 observations – far too few for a reliable period estimation.

η Cen

A description of recent emission-line activity in the shell star η Cen is given by Dachs *et al.* (1986). They conclude that at times the internal self-absorption by the shell is so

severe that it may obscure a certain fraction of the light from the photosphere. Mennickent and Vogt (1988) observed the star during 1987. The lower order Balmer lines showed weak variable emission wings and variable shell absorption cores.

Baade (1983) discovered significant line-profile variations with a time scale of hours. During a search for rapid spectroscopic variations, Ghosh *et al.*, (1988) discovered a continuum flux variability with a time scale of hours.

Our results clearly show large short-period variability. A nightly range in brightness of nearly 0.1 mag is not uncommon. The Fourier periodogram shows a very strong peak at 1.56 d^{-1} , but the resulting light curve has a large scatter. The PDM periodogram shows, in addition, large peaks at one-half and one-third this frequency. An examination of the phase plots show that a better fit is obtained by assuming a double-wave light curve with $f = 0.78 \text{ d}^{-1}$, but even then the scatter is large. By far the best phase plot is obtained with $f = 0.52 \text{ d}^{-1}$ which results in a triple-wave light curve. While double-wave light curves are common, only one case of a triple-wave light curve (α Eri, Balona *et al.*, 1987) is known.

This surprising result prompted us to re-observe η Cen for a further season in 1988 (SAAO observations only). Our analysis of these data showed quite convincingly that $f = 0.52 \text{ d}^{-1}$ is indeed present once again and gives a light curve with the least scatter. On this basis we feel that a triple-wave light curve with period of 1.927 d is likely to be the correct interpretation (Fig. 3). The peculiar light curve and the very large amplitude (0.14 mag in 1987, 0.10 mag in 1988) makes this star unique among the periodic Be variables. Line profile observations of η Cen are potentially of great importance in view of the large amplitude. If NRP is involved, the geometrical distortions and/or temperature variation must be exceedingly severe to give rise to such a large amplitude for an $\ell = |m| = 3$ mode.

HD 137518

This star was included in our program at the suggestion of C. Waelkens who found it to be a large-amplitude variable. Very little is known about this star. It has been classified as a luminous blue supergiant (B5Ia) by McConnell and Bidelman (1976). On the other hand, Garrison *et al.*, (1977) obtain a classification B1III_{ne}p with a comment that the HeI lines are strongly veiled and the suspicion that it is a double-line spectroscopic binary. In the *Michigan Spectral Catalogue* (Houk, 1978), it is given the classification B1/2(I/III_n), while Mermilliod (1987) tabulates it as O9Vn.

The photometric variability was discovered by Strohmeier *et al.* (1964) who found a range of 0.5 mag from sky patrol plates of the Bamberg Southern Station. The variability was later confirmed by Kozok (1985).

Our 1987 observations show a light range of 0.5 mag, but we could not find a definite periodicity. The Fourier periodogram shows a strong peak at $f = 0.11 \text{ d}^{-1}$ or its

one-day alias at 0.89 d^{-1} . Smaller peaks are present at $f = 0.36$ and 1.29 d^{-1} . The PDM periodogram is similar except that $f = 0.12 \text{ d}^{-1}$ is by far the strongest. We examined the light curve at many different frequencies given by the Fourier and PDM periodograms, but no single choice produced a reasonable light curve. We could assume multiperiodicity of course. If we prewhiten the data by removing a sinusoid with $f_1 = 0.115 \text{ d}^{-1}$, we obtain a noisy Fourier periodogram in which the highest peak is at $f_2 = 0.55 \text{ d}^{-1}$. A multiperiodic Fourier fit with f_1 and f_2 leads to a rms scatter of 0.08 mag which is clearly unacceptable. Removing these two frequencies leads to a strong peak at $f_3 = 0.35 \text{ d}^{-1}$, which is the third harmonic of f_1 . A Fourier fit with f_1 , f_2 and f_3 gives an rms scatter of 0.044 which is still very much higher than the expected observational error. Further prewhitening just gives a periodogram in which the noise level steadily increases towards low frequencies.

It is clear that no direct evidence for coherent multiperiodicity exists as even a solution with three frequencies is quite insufficient to describe the data. A solution can always be found by including more and more frequencies until the resulting rms scatter is sufficiently small, but there is no guarantee that the resulting solution will bear any resemblance to the true physics of the problem. We prefer to believe that the only likely periodicity in HD137518 is $P = 8.70 \text{ d}$ ($f = 0.115 \text{ d}^{-1}$) and its harmonics (Fig. 4). Prewhitening by a Fourier curve leads to a further possible periodicity of 0.685 d, but the reality of this period is far from convincing. The residual variability can be attributed to random fluctuations.

This star is a very puzzling object. Without detailed spectroscopic observations it seems impossible to determine its nature. The large light amplitude suggests some type of eclipsing phenomenon.

κ^1 Aps

Slettebak (1982) describes how two spectra, taken one day apart, show striking differences in the widths of the absorption lines (from $v \sin i = 250 \text{ km s}^{-1}$ to 350 km s^{-1}), suggesting that this star may be a double-line spectroscopic binary. Mennickent and Vogt (1988) observed this star during 1988 and found the spectrum characterized by narrow absorption cores flanked by weak emission at $H\beta$.

The Fourier periodogram of this star shows it to be another periodic variable with a frequency of 1.61 d^{-1} . The PDM periodogram also gives a strong signal at half this frequency, indicating a double-wave light curve. However, the phase plot shows that this effect arises from only five data points which deepens the second minimum; apart from this there is no strong reason to believe that the light curve is double-wave.

Observations at SAAO during 1988 give much the same results, but this time the double-wave nature is more convincing. This time the amplitude difference in the two waves is based on more observations and there are also distinct differences in their shapes. By combining the data

from the two seasons we find the best double-wave period to be 1.238 d (Fig. 5). There has been relatively little change in overall amplitude between the two seasons.

η^1 TrA

This star turned out to be one of the few constant Be stars in our sample. No sign of periodicity is evident in either the Fourier or PDM periodograms. The rms scatter of all observations is 3 millimag – the value expected for a constant star. It was not observed in 1988.

48 Lib

This star has a long history of documented spectroscopic observations. It is well known for the long-lasting presence of numerous sharp shell absorption lines and for strong regular variations of Balmer emission line profiles and radial velocities of shell absorption lines with a quasi-period of about ten years (cf., e.g. Aydin and Faraggiana, 1978). Dachs *et al.* (1986) has described the emission-line history in recent years. Mennickent and Vogt (1988) found the spectrum of 48 Lib in 1987 to be characterized by many faint FeII shell lines, some of them surrounded by emission. H β and H γ had deep central absorption cores with emission flanks ($V < R$). The V/R ratio was variable.

The 1987 data show a very strong peak at 2.49 d^{-1} in the Fourier periodogram. The PDM periodogram shows, in addition, smaller peaks at one-half and one-third of this frequency. Examination of the phase plots gives us no reason to suspect that any of the lower frequencies are physically real. There is no appreciable reduction in scatter or differences in the waves. The phase plot for the single-wave period of 0.40 d shows a rather asymmetrical light curve with a gradual decline and sharp rise (Fig. 6).

To confirm the period, we re-observed the star during 1988. We obtained the same results. The best period from the combined data is 0.4017 d. This is the shortest period we have found among the periodic Be stars.

Ringuelet-Kaswalder (1963) reported periodic radial velocity variations with $P = 0.1154295 \text{ d}$ suggesting a pulsation similar to that of the β Cep variables. Such a variation has never been confirmed, though rapid changes in spectral appearance during one night is not unknown (Aydin and Faraggiana, 1978). We phased Ringuelet-Kaswalder's data on our photometric period, but were unable to produce a satisfactory velocity curve.

χ Oph

Recent variations in the emission lines in this well-observed Be star are discussed by Dachs *et al.* (1986). The considerable number of radial velocities in the literature were analyzed by Harmanec (1987b) who tentatively proposed that χ Oph is a spectroscopic binary with a period of 34.12 d in a highly eccentric orbit ($e = 0.699$).

Balona and Engelbrecht (1987) observed this star photometrically during 1985 and suggested a possible period of 0.935 d. The available data were insufficient to discriminate between this period and its one-day alias at 14.3 d.

The shorter period was considered the more likely owing to the considerable brightness variations during the course of a night. In neither case is the resulting light curve particularly convincing.

Our 1987 data show much the same behaviour. Again, the strongest frequencies in both the Fourier and PDM periodograms occur at either 0.935 d or 14.3 d, but this time the longer period has a much higher peak. The resulting phase plot shows much less scatter for 14.3 d than for 0.935 d. Thus, by combining data from two separate sites we eliminate the alias problem and conclude that the true period is 14.3 d. Combining the 1985 and 1987 data yields a best period of 13.774 d (Fig. 7). A small number of observations were obtained in 1988.

This is the longest period so far detected for a periodic Be star. The question arises as to whether it represents an orbital variation. We looked for, but could not find, evidence for a radial velocity variation with this period from the data collected by Harmanec (1987b).

In spite of the fact that the 13.774 d period fits the photometric data, it is clear that there are variations on a much shorter time scale. However, these cannot be periodic as they would leave a signature in the periodogram. We are forced to conclude that this is one more example of the “flickering” phenomenon so common amongst these stars.

ζ Oph

ζ Oph (O9.5Ve) is one of the most rapidly rotating stars known. Emission lines at H α and HeI were observed between July 1973 and April 1974 (Irvine, 1974 ; Niemela and Mendez, 1974 ; Barker and Brown, 1974). During 1979 there was no hint of emission at H α , but by March 1980 H α appeared as a well-developed double-emission feature quite similar to that observed during 1974 (Ebbets 1981).

Walker *et al.* (1979) were the first to discover the by now well-known phenomenon of “moving bumps” in the line profile of B stars. It was in this star, ζ Oph, that they were first detected. At that time they were interpreted as nonuniformities in the stellar photosphere carried across the line of sight by rotation. In this way a rotational period of 21.7 hours was deduced, leading to an estimated $v \sin i = 560 \text{ km s}^{-1}$. This value is considerably higher than observed ($v \sin i = 370 \text{ km s}^{-1}$), but it is sensitive to the adopted stellar radius. Harmanec (1989) has shown that a somewhat smaller radius, which is still consistent with the physical parameters for this star, will bring the expected projected rotational velocity into good agreement with the observed value.

Walker *et al.* (1981) presented further observations of this new phenomenon and proposed an alternative model in which the moving bumps arise from obscurations in the circumstellar material. They also suggested NRP as a possible explanation.

Vogt and Penrod (1983) obtained extensive high-resolution spectra of ζ Oph and confirmed the moving bumps seen

by the earlier workers. They rejected the interpretation in terms of rotational modulation or obscuration by circumstellar material. While these models give rise to moving bumps in agreement with the observations, the expected light variations are not seen. They favour the NRP model. Harmanec (1989) has criticized this conclusion as it is based on only one night of photometry which does indeed show some variability. Nevertheless, the NRP model has been accepted by most groups as the most probable explanation in this and other Be stars.

It is clear that intensive photometric observations should shed considerable light on the nature of moving bumps. According to Vogt and Penrod (1983) the NRP interpretation should give rise to periodic light variations of amplitude 0.02 mag with a period of a few hours. Such periodic variations are easily detectable.

We started our photometric campaign on this star in 1985. The results of this first season were briefly reported in Balona and Engelbrecht (1987). During 1985 ζ Oph showed a clear indication of pulsation with a period of 0.193 d and amplitude of 0.02 mag. There was also a longer period of 1.075 d with amplitude 0.03 mag (Fig. 8a). The short-period pulsations were interpreted as NRP, but not as a new phenomenon. Since ζ Oph lies on the hot end of the β Cep instability strip, it seemed reasonable to classify the star as a β Cep variable. However, the presence of the longer period in a β Cep star is most unusual; rather it is indicative of a periodic Be star. It is also of particular interest to note that if the light curve is phased on the long period of 1.075 d the short-period variation is easily visible as a sinusoidal modulation of the light curve (Fig. 8b). It looks very much as if the short period is practically one-sixth of the long period to within observational errors. This seems to confirm Harmanec's (1987a) suspicion that whenever a short-period is present it turns out to be a sub-multiple of the rotation period.

We observed ζ Oph during 1987 expecting to confirm the provisional results announced by Balona and Engelbrecht (1987). The Fourier and PDM periodograms both show some power at frequencies of 1.00 or 2.00 d^{-1} with a peak-to-peak amplitude of 0.01 mag. There is no sign of the short period variation seen in 1985. While the long period is close to the value seen in 1985, it is practically the same as one cycle per day and, as such, must give rise to a grave suspicion that this is an artifact. Indeed, since we had to observe ζ Oph at SAAO using a neutral density filter it would not be entirely surprising if some slight systematic error was present in the data. Under these conditions we can understand the results of the periodogram and we do not consider the one day period to be real. Indeed, the rms error of one observation from both sites is only 4 millimags., i.e. close to the expected rms error of a constant star. We conclude that ζ Oph was constant in light during June/July 1987.

Some observations at SAAO during early 1988 again

indicated a constant light. But in June of that year a dramatic change took place: our results of four photometric nights show marked photometric activity with night-to-night variations exceeding 0.01 mag. It would be very interesting to discover if this is associated with the development of emission at $\text{H}\alpha$. Further photometric monitoring is planned for 1989.

ι Ara

Our photometric observations of this star now cover four seasons (1985 - 1988). The 1985 light variations were particularly difficult to interpret in spite of the fact that there were no long-term trends or other difficulties. The problem is the very large flickering associated with this star which tends to mask the underlying variation almost completely. Balona and Engelbrecht (1987) obtained a double-wave light curve with period 0.515 d as the most likely solution for this season. During 1986 the amplitude appeared to increase steadily during the observing period of one month, but a period of 0.56 d could be extracted. This is the same as the period obtained in 1985 within the observational error and results in a single-wave light curve.

Our 1987 data showed that once more the star was active with a short time scale. The strongest peak in the Fourier periodogram occurs at 0.265 d, but the PDM periodogram indicates that a double-wave solution is better as the strongest peak occurs at 0.53 d. The 1988 observations are too few to derive a period, but a solution with $P = 0.56$ d fits the data well.

The conclusion is that in spite of the low signal-to-noise ratio, a period close to 0.55 d is derived independently from each observing season. On this account we feel confident that this is indeed the correct period. Combining all the data gives a best solution of $P = 0.5565$ d (Fig. 9). The flickering in this star is one of the largest we have encountered. As a consequence the periodicity in this star is somewhat blurred, but periodicity implies coherence over many cycles and this period was the only one where such a coherence was observed. The rms deviation from the light curve is as large as 0.02 mag though the amplitude is as large as 0.10 mag.

α Ara

Mennickent and Vogt (1988) found the spectrum in 1986-7 to show double emission in $\text{H}\beta$ and $\text{H}\gamma$ ($V \simeq R$) with small variations in the V/R ratio on a short time scale.

This star was immediately recognized as a large-amplitude short-period variable during the first observing run in 1985. The Fourier periodogram shows strong peaks at frequencies of 2.04 or 3.04 d^{-1} and its one day aliases. These two frequencies are also the strongest in the PDM periodogram, but the double-wave solution at 1.52 d^{-1} is nearly as strong. An examination of the light curve does indicate that there is a significant difference in amplitude between the two waves in the double-wave solution. Thus we adopted a period of 0.658 d ($f = 1.52 \text{ d}^{-1}$, Balona and Engelbrecht, 1987).

In 1986 the highest peak in the periodograms occurs at

2.04 d^{-1} ; the solution for 1985 thus appears to be a one-day alias of the true period. That a wrong choice was made is not surprising in view of the closeness of the frequency to a half a day and the severe aliasing problem.

Our 1987 data shows that the ambiguity is completely resolved thanks to the different longitudes of the observing sites. The true period is either a single-wave with $P = 0.49 \text{ d}$ or a double-wave with $P = 0.98 \text{ d}$. It is very difficult to be sure of the double-wave solution since the two waves are of nearly equal amplitude, but it is a distinct possibility. Unfortunately, the period is so close to one day that we cannot use the other observations (all made at SAAO) to settle this problem. By combining all the data we find a best period of 0.9807 d for the double-wave solution (Fig. 10). Results from 1988 indicate that this period fits the data well.

66 Oph

During 1987 this star showed double emission at $H\beta$ and $H\gamma$ (Mennickent and Vogt, 1988).

We observed this star for one season only (1987). Apart from a gradual increase in brightness, there is no indication of periodicity. Nevertheless, the star is variable as the scatter (0.01 mag) is much larger than expected. It appears that we are seeing only the flickering component.

V986 Oph

This star is not known as a Be star, but we have included it here owing to its great interest in connection with the role of NRP in the Be phenomenon. V986 Oph has a long history of photometric variability. Lynds (1959) determined a period of 0.289 d and classified it as a β Cep variable. Jerzykiewicz (1975) confirmed this period from seven nights of photometry and deduced a peak-to-peak amplitude of 0.014 mag . However, not all the data could be described by a simple sinusoidal variation with this period. Pike and Lloyd (1979) failed to detect any radial velocity variations with the photometric period, though a systematic decrease in nightly mean velocities was observed.

To clarify the nature of this star, C. D. Pike organized an international campaign for July 1980. Fullerton *et al.* (1985) analyzed both the photometry and radial velocities obtained during the campaign and found a period of 0.325 d with a light amplitude of 0.01 mag in B . No significant radial velocity variations were found, but the systematic variation in mean nightly velocity was confirmed and orbital elements derived ($P = 25.56 \text{ d}$, $2K = 35 \text{ km s}^{-1}$, $e = 0.23$). They also obtained an extensive set of high dispersion line-profile observations and discovered high-order profile variations progressing from blue to red with a period of about 0.3 d . They deduced $\ell = 6$, but at times $\ell = 4$ and $\ell = 8$ are excited.

This confusing picture of V986 Oph prompted us to include it in our observing project. Our photometry for 1985 was obtained over 10 nights and showed light variations at a low level. Fourier periodogram analysis showed peaks at 1.74 and 1.16 d^{-1} and their one-day aliases (Fig. 11a - top panel). An unique period or periods could not be

determined from these data alone. Combining these data with all available photometry showed the most likely period to be 0.30 d or 0.23 d , its alias.

Our 1987 data again shows micro-variability, but owing to the spacing in longitude the aliasing problem largely disappears. The Fourier periodogram shows three strong peaks at $f_1 = 3.30$, $f_2 = 1.41$ and $f_3 = 0.77 \text{ d}^{-1}$ with semi-amplitudes of 4.7, 4.6 and 3.8 millimag respectively (Fig. 11a - middle panel). Prewhitening by f_1 leaves f_2 and f_3 with f_2 slightly stronger (Fig. 11a - bottom panel). Further prewhitening by either f_2 or f_3 leads to a pure noise spectrum. Our conclusion is that there are two periods present: the short period is well determined to be 0.303 d (Fig. 11b), but the long period is uncertain, it is either 1.299 d or 0.709 d .

The short period is the same as the one originally suggested by Lynds and Jerzykiewicz and also found in our 1985 data. We cannot confirm the period found by Fullerton *et al.* (1985); indeed there is no indication of a 0.325 d period in our data at all. We note, however, that the 0.303 d period is the same as their estimate of the period of the high-order line-profile variations. The low photometric amplitude is consistent with a high-order NRP mode. The nature of the long period is uncertain. It could be another NRP mode (either rotationally split p -mode or a g -mode), but the period is close to the expected period of rotation so that some kind of rotational modulation is not ruled out.

There appears to be no compelling reason to classify this star as anything other than a β Cep variable as originally suggested by Lynds (1959). Admittedly, it has the longest known period of a star in this class, but this is not surprising in view of its early spectral type and luminosity. It does apparently oscillate with somewhat higher order modes than the majority of other β Cep variables, but this is not a criterion for reclassification. The only reasonable criterion for classification is a definite indication of a different pulsation mechanism. Since there is no reason to believe that the pulsation mechanism in this star is any different from other β Cep variables, we feel that reclassification serves merely to confuse the issue.

ϵ Cap

Mennickent and Vogt (1988) observed this star in 1986 and found deep central absorption with no traces of emission at $H\beta$. Dachs *et al.* (1986) found $H\alpha$ to have a central absorption and emission wings in 1982.

Our 1985 photometry showed that ϵ Cap underwent a sharp drop in brightness of nearly 0.2 mag in just three days followed by a rapid recovery. Excluding this period, the periodogram shows a set of peaks with $f = 2.60 \text{ d}^{-1}$ and its one-day aliases, but a double-wave light curve with half this frequency is possible as the waves have distinctly different amplitudes. This solution ($P = 0.769 \text{ d}$) was proposed by Balona and Engelbrecht (1987).

The 1986 photometry again showed indications of peri-

odicity, but we could not confirm the solution derived for 1985. Instead a frequency close to two cycles per day was indicated. The severe aliasing problem prevented a meaningful solution.

The combined ESO and SAAO data of 1987 showed a slow brightening over the observing run. Removing this trend showed a possible period very close to one day. By combining all available data we determined the best value to be 1.030 d (Fig. 12). Because it is so close to one day and because of the discrepancy obtained in 1985, this period must be regarded as very uncertain.

η PsA

Mennickent and Vogt (1988) find H β to be very diffuse with weak double emission ($V \simeq R$) around a broad absorption core.

We were unable to detect any clear periodicity in η PsA during 1985, but the star was clearly a low-amplitude variable. A Fourier periodogram suggested a possible period of 0.774 d, which is the value quoted by Balona and Engelbrecht (1987). During 1986, the low-level variability (flickering) was again evident. No reliable period could be extracted. The same is true for the combined ESO and SAAO data of 1987.

HR 8408

Slettebak (1982) discovered HR 8408 to be a Be star showing faint double emission ($V \simeq R$) at H β surrounding an absorption core. Mennickent and Vogt (1988) did not see emission during 1986, but suspect incipient emission at H β . Buscombe and Morris (1961) found a radial velocity range of 40 km s $^{-1}$.

This star was originally used as a photometric standard before we discovered its variability. The 1985 SAAO photometry shows clear indications of periodicity with a frequency $f = 2.53$ d $^{-1}$, though a double-wave solution with half this frequency shows the two waves to have significantly different amplitudes.

The 1986 data produced the highest peak in the Fourier periodogram at $f = 1.53$ d $^{-1}$, i.e. the one-day alias of the solution adopted for 1985. Again, a double-wave solution appeared to be significant.

The combined ESO and SAAO data of 1987 removed the aliasing problem. The correct frequency was found to be $f = 1.53$ d $^{-1}$ or the double-wave solution at $f = 0.765$ d $^{-1}$. By combining all available data we refined the period to 1.3106 if we adopt the double-wave solution (Fig. 13). This value (or possibly the single-wave period of 0.6553 d) seems to be well established.

\omicron Aqr

The emission features in this star have been relatively static since 1975. The spectrum shows double emission in the Balmer lines ($V \simeq R$) and a central shell absorption feature (Mennickent and Vogt 1988).

Photometry during 1985 showed a steady brightening of the star over a period of three months with a large excursion

in magnitude on one occasion. These complications made period extraction very difficult. A subset of the data indicated a possible double-wave light curve with $P = 1.449$ d (Balona and Engelbrecht 1987). During 1986 a much larger amount of data was obtained and the star was more quiescent. Both the Fourier and PDM periodograms show a strong signal at $f = 1.40$ d $^{-1}$ ($P = 0.71$ d) - half the period found in 1985. However, a double-wave light curve with the period found in 1985 also described the 1986 data well.

The combined ESO and SAAO data for 1987 again shows a strong signal with $f = 1.39$ d $^{-1}$ or the double-wave equivalent at $f = 0.70$ d $^{-1}$. This time the amplitude difference in the waves did not appear to be significant. Considering the good agreement of the period determinations of three seasons, we can state with little doubt that the period of this star is $P = 1.4325$ d from all available data (Fig. 14). This gives a double-wave light curve with distinctly different amplitudes for the two waves in 1985 and 1986. An increase in amplitude from 0.01 mag in 1986 to 0.02 mag in 1987 seems to have occurred.

One of the aims of this project was to determine the temperature range at which periodic Be stars could be found. \omicron Aqr is one of the coolest stars (B7IVe) which has been found to be periodic. Our impression is that periodicity is most likely to be found amongst earlier spectral types.

ψ^2 Aqr

This star is not known as an emission-line star. We decided to include it in our observing program owing to its very rapid rotation ($v \sin i = 332$ km s $^{-1}$) and to test the hypothesis that rapidly-rotating non-Be stars do not show low-order profile variations (which are expected to give rise to periodic light variations). Abt and Levy (1978) regard the star as a probable spectroscopic binary.

The SAAO data for 1986 showed a strong peak at $f = 1.87$ d $^{-1}$ in the Fourier periodogram. The PDM periodogram suggested that a double-wave solution with half this frequency is more appropriate. The ESO and SAAO data of 1987 confirmed these conclusions. We have adopted a double-wave solution with $P = 1.073$ d obtained from an analysis of all existing photometry. A large decrease in amplitude seems to have occurred between the two seasons (Fig. 15).

The light variation in this star is in every respect similar to those in Be stars. It would be of great value to look for possible emission at H α in this star. If it can be shown that emission is definitely not present but that short-period light variations are present, it could offer an important clue to the relationship between the short-period variations and the Be phenomenon.

5. Colour variations.

There are very few observations of periodic Be stars observed in more than one colour. Van Vuuren *et al.* (1988) found that the colour variations in periodic Be stars in the

open cluster NGC 3766 were very small. The evidence indicates that the star is bluest (hottest) when brightest. The ESO observations produced simultaneous *uvby* colours. In figures 16 - 20 we show light and colour curves for some of the periodic Be stars observed at ESO. In general, the conclusions of van Vuuren *et al.*, (1988) are supported. The best evidence comes from observations of 48 Lib (Fig. 17). There is a small phase shift between *b* and *u-b*, but the star is brightest when bluest.

6. Conclusions.

A detailed interpretation of the results found in this paper and others in the series will be presented elsewhere. Here we summarize the main findings.

We found periodic light variations in almost all Be stars. Of the 17 candidates, only three stars were constant in light, three others had uncertain periods. The Be star HD137518 has a very unusual light curve which may not be strictly periodic or else could be another example of a triple-wave light curve as found in η Cen and α Eri.

In most cases the light curves are double waves with unequal maxima and minima. Sometimes the waves are of nearly equal amplitude, so it is quite possible that the periods of some apparently single-wave variables may be in error by a factor of two. Such a situation was found for two Be stars in the cluster NGC3766 (Ahmed 1 and Ahmed 15) where an apparently single-wave light curve became a double-wave light curve with twice the period in the course of a year (van Vuuren *et al.* 1988). The triple-wave light curve of η Cen is remarkable for its large amplitude, but has a counterpart in the bright star α Eri (Balona *et al.*, 1987).

It has become increasingly clear, both in this work and in previous studies of the light variability of periodic Be stars, that the light variations are never adequately described

by purely periodic variations. There is always a residual scatter which is often many times larger than the expected observational error. We have called these random, short-term variations *flickering*. It seems that flickering is closely associated with periodic variability. The constant Be stars are really constant and do not show this flickering (or at least not to the same extent).

In all cases we found the period of light variation to be consistent with the expected period of rotation of the star. There is no clear evidence for multiperiodicity in any of the stars. In the case of ζ Oph we found a β Cep - like variation with a period of 0.193 d superimposed on a longer period of 1.075 d which is practically six times the length of the shorter period and could be the rotation period. This behaviour was only seen for one season. This star is clearly an exception to the rule and deserves closer study. For χ Oph (and possibly HD137518) we found evidence for a period longer than the expected period of rotation. It is not very clear whether this variation is intrinsic to the star (as is the case in general) or a result of a close companion or circumstellar material.

The results of this work clearly show the importance of multi-site observations for these short-period variables. The aliasing problem that still existed in some stars was completely resolved by combining results from the two sites.

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TABLE I. — *A summary of stars observed. Notes are given for program stars ; the others are comparison stars. The abbreviation 2W and 3W stands for double-wave and triple -wave light curve respectively. The grouping includes stars which are in close proximity and for which the same transparency corrections were applied. The projected rotational velocity, the mean Strömgren b magnitude and the total number of observations is shown.*

HR	HD	Name	MK type	vsini		N	Notes
HR 5190	HD 120307	ν Cen	B2IV	91	3.33	279	P = 2.621 d.
HR 5193	HD 120324	μ Cen	B2IV-Ve	175	3.39	360	P = 2.1017 d, 2W.
HR 5439	HD 127971		B7V		5.872	202	
HR 5440	HD 127972	η Cen	B1.5Vne	333	2.32	352	P = 1.927 d, 3W.
HR 5668	HD 135348		B3IV		6.026	178	
	HD 137518		B1IIIInep		7.94	183	P = 8.696 d?
HR 5730	HD 137387	κ^1 Aps	B1pne		5.39	308	P = 1.238 d, 2W.
HR 5786	HD 138867		B9V		5.950	246	
HR 6172	HD 149671	η^1 TrA	B7IVe		5.89	202	Constant.
HR 6233	HD 151441		B8II-III		6.150	215	
HR 5902	HD 142096		B2.5V	197	5.065	230	
HR 5941	HD 142983	48 Lib	B5IIIpe	393	4.81	232	P = 0.4017 d.
HR 5993	HD 144470		B1V	142	3.986	401	
HR 6112	HD 147933		B2V	303	4.781	134	
HR 6118	HD 148184	χ Oph	B2IV:pe	134	4.61	346	P = 13.774 d.
HR 6175	HD 149757	ζ Oph	O9.5Vne	379	2.67	313	P ₁ = 1.075, P ₂ = 0.193 d.
HR 6224	HD 151133		B9.5III	100	6.050	131	
HR 6451	HD 157042	ι Ara	B2IIIIne	369	5.25	542	P = 0.556 d, 2W.
HR 6460	HD 157243		B7III	150	5.120	684	
HR 6475	HD 157599		B8-9V		6.190	405	
HR 6510	HD 158427	α Ara	B2IIIIne	298	2.81	545	P = 0.9807 d, 2W.
HR 6712	HD 164284	66 Oph	B2Ve	221	4.77	160	Constant.
HR 6719	HD 164432		B2IV		6.352	287	
HR 6732	HD 164716		B9V		6.878	66	
HR 6747	HD 165174	V986 Oph	B0IIIIn	434	6.20	274	P ₁ =0.303, P ₂ =1.3/0.7 d.
HR 8260	HD 205637	ϵ Cap	B2.5Vpe	293	4.36	567	P = 1.03 d?
HR 8293	HD 206546		A3m	47	6.336	581	
HR 8386	HD 209014	η PsA	B8Ve		5.39	598	Constant.
HR 8408	HD 209522		B4IVne	300	5.90	597	P = 1.3106 d, 2W.
HR 8446	HD 210300		A5V		6.500	482	
HR 8402	HD 209409	o Aqr	B7IVe	227	4.67	526	P = 1.433 d, 2W.
HR 8451	HD 210419		A1Vnn	254	6.270	455	
HR 8533	HD 212404		A0V	65	5.750	406	
HR 8840	HD 219402		A3V		5.561	389	
HR 8858	HD 219688	ψ^2 Aqr	B5V	332	4.32	388	P = 1.073 d, 2W.

TABLE IIa (continued)

Table with multiple columns for HR numbers (5941, 6118, 6175, 6451), HJD values, and component letters (a, b). The data is organized into four main sections, each representing a different set of stars, with each section containing two columns of star identifiers.

TABLE IIa (continued)

HR 0858	HJD	b	HJD	b	HJD	b	HR 0858	HJD	b
	6682.4665	4.346	6692.5017	4.339	6700.2841	4.344		6983.6001	4.328
	6682.4823	4.343	6692.5257	4.346	6700.3067	4.348		6983.6180	4.320
6676.3843	4.324	6682.4971	4.341	6693.2458	4.346	6700.3342	4.349	6967.6069	4.328
6676.4039	4.329	6682.5135	4.343	6693.2693	4.345	6702.4550	4.344	6967.6253	4.323
6676.4192	4.330	6682.5289	4.344	6693.2913	4.354	6702.4786	4.340	6967.6373	4.319
6676.4355	4.333	6682.5453	4.346	6693.3143	4.352	6702.5000	4.345	6967.6477	4.319
6676.4539	4.337	6682.5615	4.347	6693.3366	4.356	6702.5202	4.328	6967.6578	4.318
6676.4716	4.339	6682.5776	4.347	6693.3627	4.349	6702.5367	4.337	6970.5551	4.321
6676.4939	4.343	6682.5908	4.344	6693.3874	4.353	6703.2821	4.355	6970.5805	4.324
6676.5127	4.342	6682.6036	4.344	6693.4107	4.346	6703.3038	4.343	6970.6139	4.327
6676.5294	4.349	6684.3513	4.330	6693.4305	4.348	6703.3265	4.349	6970.6373	4.329
6676.5464	4.343	6684.3726	4.329	6693.4502	4.339	6703.3491	4.346	6970.6543	4.329
6676.5623	4.342	6684.3962	4.322	6693.4755	4.335	6703.3689	4.349	6972.5345	4.334
6676.5747	4.347	6684.4185	4.331	6693.4987	4.335	6703.3899	4.354	6972.5559	4.335
6676.5898	4.341	6684.4352	4.329	6693.5202	4.325	6703.4100	4.353	6972.5784	4.334
6676.6024	4.341	6684.4512	4.338	6693.5372	4.316	6703.4302	4.352	6972.5967	4.332
6677.3567	4.344	6684.4862	4.339	6693.5524	4.312	6703.4442	4.353	6972.6102	4.331
6677.3754	4.355	6686.2812	4.343	6693.5647	4.309	6703.4838	4.348	6972.6236	4.330
6677.3980	4.328	6686.2987	4.343	6694.2514	4.320	6703.4992	4.345	6972.6363	4.334
6677.4197	4.327	6686.3157	4.342	6694.2981	4.333	6703.5167	4.349	6972.6489	4.331
6677.4380	4.319	6686.3317	4.345	6694.3221	4.338	6703.5341	4.344	6972.6615	4.335
6677.4543	4.324	6686.3492	4.347	6694.3513	4.341	6959.5189	4.334	6972.6745	4.333
6677.4703	4.310	6686.3680	4.347	6694.3779	4.334	6959.5464	4.329	6972.6868	4.336
6677.4879	4.322	6686.3874	4.351	6694.3993	4.334	6959.5968	4.321	6975.5328	4.327
6677.5059	4.319	6686.4054	4.349	6694.4226	4.336	6959.6190	4.315	6975.5756	4.327
6677.5223	4.317	6686.4218	4.352	6694.4471	4.327	6959.6385	4.311	6975.5923	4.329
6677.5412	4.333	6686.4365	4.353	6695.2731	4.336	6959.6543	4.318	6975.6064	4.328
6677.5623	4.332	6686.4519	4.354	6695.3186	4.350	6959.6666	4.318	6975.6216	4.329
6677.5792	4.326	6686.4668	4.352	6695.3453	4.340	6959.6793	4.319	6975.6362	4.327
6680.2589	4.324	6686.4823	4.348	6695.3725	4.345	6960.5232	4.347	6975.6502	4.326
6680.2748	4.322	6686.4994	4.348	6695.3940	4.347	6960.5705	4.339	6975.6639	4.324
6680.2937	4.334	6686.5167	4.339	6695.4162	4.349	6960.5949	4.342	6978.4783	4.321
6680.3115	4.334	6686.5325	4.337	6695.4384	4.349	6960.6137	4.335	6978.5277	4.326
6680.3296	4.329	6686.5485	4.332	6695.5187	4.359	6960.6306	4.326	6978.5487	4.329
6680.3475	4.330	6686.5592	4.333	6695.5354	4.370	6960.6433	4.322	6978.5677	4.329
6680.3650	4.338	6686.5702	4.328	6696.2696	4.332	6960.6550	4.326	6978.5886	4.332
6680.3865	4.343	6686.5829	4.330	6696.2919	4.333	6960.6671	4.322	6978.6033	4.335
6680.4088	4.348	6686.5947	4.329	6696.3132	4.342	6961.4888	4.324	6980.4987	4.330
6680.4312	4.352	6688.2442	4.306	6696.3362	4.344	6962.5455	4.347	6980.5345	4.339
6680.4471	4.362	6688.2636	4.314	6696.3629	4.341	6962.5731	4.345	6980.5526	4.340
6680.4639	4.359	6688.2995	4.336	6696.3845	4.341	6962.5970	4.344	6980.5713	4.339
6680.4906	4.353	6688.3175	4.331	6696.4058	4.336	6965.5223	4.320	6980.5907	4.334
6680.5070	4.354	6688.3360	4.335	6696.4280	4.345	6965.5637	4.316	6980.6034	4.330
6680.5412	4.353	6690.3101	4.330	6696.4491	4.354	6965.6043	4.327	6980.6165	4.325
6680.5580	4.350	6690.3343	4.336	6697.2586	4.334	6965.6293	4.324	6980.6314	4.325
6680.5884	4.366	6690.3702	4.323	6697.2817	4.331	6965.6441	4.327	6980.6437	4.325
6682.2413	4.346	6690.3996	4.335	6697.3046	4.332	6965.6592	4.329	6980.6550	4.325
6682.2603	4.342	6690.4272	4.327	6697.3287	4.335	6966.5474	4.329	6980.6676	4.323
6682.2768	4.337	6690.4497	4.327	6697.3509	4.334	6966.5740	4.322	6981.4877	4.324
6682.2937	4.339	6690.4753	4.340	6697.3769	4.334	6966.5906	4.321	6981.5214	4.333
6682.3106	4.335	6690.4962	4.335	6697.3983	4.342	6966.6054	4.319	6981.5445	4.335
6682.3271	4.335	6690.5167	4.342	6697.4195	4.339	6966.6213	4.319	6981.5722	4.334
6682.3439	4.338	6690.5384	4.340	6698.2647	4.349	6966.6441	4.320	6981.5970	4.341
6682.3634	4.336	6690.5549	4.349	6698.2863	4.349	6966.6587	4.324	6981.6115	4.339
6682.4011	4.328	6690.5787	4.353	6698.3081	4.344	6967.5019	4.338	6981.6253	4.336
6682.4189	4.352	6692.4370	4.346	6698.3303	4.340	6967.5513	4.338	6981.6408	4.334
6682.4346	4.338	6692.4575	4.344	6698.3537	4.340	6967.5746	4.331	6981.6549	4.331
6682.4501	4.339	6692.4789	4.342	6700.2636	4.353	6967.5907	4.328	6981.6669	4.325

TABLE IIb. — Four-colour Strömgren photometry for periodic Be stars observed at ESO. Observations between JD2446991 and 2447000 were obtained at SAAO. The heliocentric Julian day is with respect to JD 2440000.000.

Table with columns for star name (HR 5193, HR 5440, HD 137518, HR 5730), HJD, and photometric bands (u, v, b, y). It contains multiple columns of data for each star, organized by observation date (HJD).

TABLE IIb (continued)

HR 5941

Table with 5 columns: HJD, u, v, b, y. Contains star data for HR 5941.

HR 6118

Table with 5 columns: HJD, u, v, b, y. Contains star data for HR 6118.

HJD

Table with 5 columns: u, v, b, y. Continuation of star data for HR 5941.

HR 5941

Table with 5 columns: HJD, u, v, b, y. Contains star data for HR 5941.

HR 6118

Table with 5 columns: HJD, u, v, b, y. Contains star data for HR 6118.

TABLE IIb (continued)

HR 8408	HJD	u	v	b	y	HR 8858	HJD	u	v	b	y				
	6973.7103	6.697	6.226	5.900	5.947	6961.7780	5.234	4.668	4.313	4.370	6977.8993	5.229	4.673	4.316	4.379
	6977.7457	6.687	6.220	5.886	5.943	6961.8076	5.235	4.668	4.315	4.373	6978.7713	5.271	4.694	4.335	4.396
6960.7486	6.688	6.225	5.891	5.942		6961.8295	5.235	4.666	4.313	4.372	6978.8033	5.260	4.689	4.329	4.390
6960.7697	6.689	6.220	5.888	5.939		6961.8431	5.241	4.671	4.318	4.375	6978.8292	5.258	4.691	4.331	4.394
6960.8037	6.699	6.232	5.900	5.951		6961.8581	5.242	4.675	4.319	4.378	6979.7727	5.276	4.695	4.341	4.398
6960.8410	6.699	6.231	5.900	5.952		6961.8695	5.244	4.676	4.321	4.378	6979.7958	5.267	4.686	4.331	4.391
6960.8676	6.695	6.230	5.900	5.951		6961.8883	5.253	4.671	4.324	4.381	6979.8163	5.263	4.689	4.330	4.394
6960.8785	6.697	6.231	5.901	5.953		6961.9007	5.252	4.672	4.325	4.380	6979.8317	5.266	4.685	4.331	4.391
6960.8869	6.692	6.229	5.899	5.951		6962.7792	5.250	4.676	4.319	4.373	6979.8446	5.266	4.687	4.330	4.393
6960.8946	6.695	6.226	5.899	5.949		6962.8157	5.231	4.666	4.308	4.371	6979.8735	5.256	4.680	4.324	4.387
6960.9011	6.693	6.229	5.898	5.949		6962.8336	5.225	4.661	4.303	4.367	6979.9086	5.250	4.674	4.318	4.380
6960.9084	6.696	6.229	5.901	5.950		6962.8626	5.232	4.661	4.307	4.369	6980.7984	5.263	4.683	4.326	4.387
6960.9141	6.697	6.230	5.898	5.951		6962.8809	5.233	4.665	4.309	4.374	6980.8307	5.263	4.685	4.329	4.391
6960.9203	6.697	6.230	5.901	5.951		6962.8960	5.237	4.667	4.312	4.376	6980.8520	5.258	4.683	4.327	4.388
6965.8142	6.673	6.213	5.882	5.936		6962.9082	5.237	4.670	4.313	4.380	6980.8701	5.255	4.678	4.324	4.385
6965.8165	6.679	6.217	5.886	5.938		6963.7719	5.264	4.692	4.328	4.373	6980.8936	5.253	4.677	4.322	4.382
6965.8397	6.676	6.216	5.886	5.939		6965.8206	5.269	4.691	4.334	4.384	6980.9044	5.251	4.674	4.320	4.381
6965.8571	6.675	6.215	5.884	5.936		6965.8447	5.268	4.691	4.334	4.385	6982.7232	5.293	4.698	4.339	4.403
6965.8799	6.671	6.216	5.884	5.936		6965.8626	5.264	4.687	4.331	4.380	6991.6139	4.877	4.191	4.161	4.424
6965.8939	6.674	6.217	5.883	5.938		6965.8841	5.263	4.686	4.330	4.381	6991.6622	4.864	4.402	4.367	4.429
6967.7317	6.680	6.216	5.881	5.937		6965.8989	5.267	4.687	4.333	4.382	6998.5703	4.857	4.369	4.329	4.397
6967.7711	6.682	6.222	5.887	5.942		6967.7744	5.261	4.682	4.330	4.387	6998.6287	4.854	4.372	4.338	4.393
6967.8108	6.684	6.228	5.892	5.946		6967.8146	5.259	4.683	4.326	4.387	6999.5802	4.844	4.357	4.322	4.380
6967.8441	6.679	6.223	5.888	5.941		6967.8466	5.261	4.680	4.327	4.387	6999.6321	4.842	4.365	4.331	4.391
6967.8674	6.676	6.219	5.886	5.935		6967.8707	5.260	4.681	4.327	4.387	7000.5960	4.832	4.371	4.324	4.380
6968.7034	6.679	6.215	5.882	5.939		6968.7853	5.239	4.673	4.318	4.378					
6968.7425	6.697	6.232	5.897	5.953		6968.8226	5.245	4.676	4.320	4.382					
6968.7800	6.692	6.223	5.890	5.946		6968.8511	5.257	4.689	4.332	4.393					
6968.8194	6.687	6.225	5.887	5.942		6968.8782	5.256	4.685	4.329	4.390					
6968.8485	6.697	6.230	5.897	5.949		6968.9003	5.251	4.679	4.323	4.383					
6968.8753	6.685	6.217	5.884	5.937		6970.7738	5.253	4.682	4.327	4.386					
6968.8973	6.686	6.220	5.887	5.940		6970.8097	5.252	4.679	4.324	4.384					
6969.7085	6.685	6.223	5.891	5.943		6970.8313	5.255	4.681	4.326	4.386					
6969.7479	6.690	6.221	5.891	5.942		6970.8595	5.248	4.676	4.319	4.380					
6970.7308	6.689	6.226	5.892	5.944		6970.8698	5.248	4.673	4.319	4.379					
6970.7703	6.696	6.218	5.892	5.944		6970.8863	5.250	4.675	4.321	4.380					
6970.8057	6.686	6.221	5.888	5.941		6970.8956	5.252	4.677	4.322	4.382					
6970.8280	6.685	6.222	5.888	5.942		6971.7635	5.276	4.699	4.341	4.400					
6970.8571	6.686	6.221	5.890	5.944		6971.8075	5.255	4.682	4.326	4.386					
6970.8669	6.690	6.227	5.896	5.950		6971.8378	5.250	4.677	4.321	4.384					
6970.8780	6.690	6.227	5.895	5.949		6971.8616	5.246	4.674	4.318	4.382					
6970.8933	6.689	6.227	5.895	5.948		6971.8842	5.248	4.672	4.317	4.379					
6971.7058	6.695	6.224	5.888	5.944		6971.8957	5.245	4.671	4.315	4.378					
6971.7605	6.691	6.216	5.884	5.942		6971.9073	5.248	4.671	4.317	4.377					
6971.8049	6.689	6.219	5.885	5.942		6972.7551	5.269	4.690	4.334	4.392					
6971.8347	6.696	6.220	5.889	5.945		6972.8013	5.263	4.684	4.330	4.391					
6971.8582	6.694	6.222	5.890	5.945		6972.8357	5.255	4.680	4.324	4.384					
6971.8813	6.698	6.226	5.893	5.947		6972.8763	5.248	4.676	4.322	4.383					
6971.8926	6.700	6.227	5.893	5.947		6972.8985	5.249	4.677	4.323	4.382					
6971.9044	6.698	6.225	5.892	5.949		6973.7722	5.279	4.693	4.344	4.402					
6972.6706	6.701	6.224	5.894	5.942		6977.7843	5.238	4.679	4.324	4.384					
6972.7073	6.699	6.224	5.896	5.948		6977.8139	5.235	4.679	4.323	4.385					
6972.7523	6.689	6.216	5.888	5.940		6977.8367	5.239	4.677	4.322	4.385					
6972.7981	6.686	6.214	5.886	5.937		6977.8524	5.235	4.673	4.320	4.382					
6972.8313	6.694	6.220	5.890	5.943		6977.8660	5.228	4.673	4.318	4.381					
6972.8698	6.695	6.224	5.895	5.943		6977.8859	5.233	4.674	4.320	4.381					
6972.8952	6.699	6.230	5.899	5.950											
6973.6677	6.673	6.209	5.886	5.934											

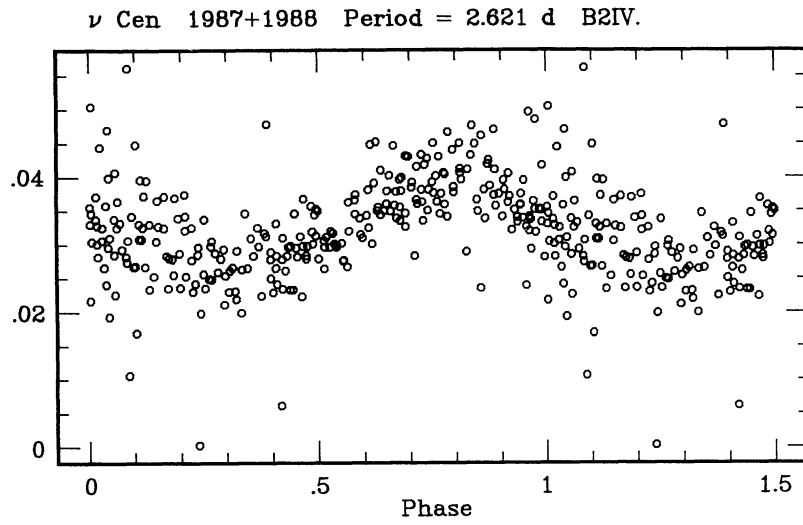


FIGURE 1. — Light curve of ν Cen ($P = 2.621$ d). In this and subsequent figures, the scale is in magnitudes and the epoch of phase zero is JD2446000.000.

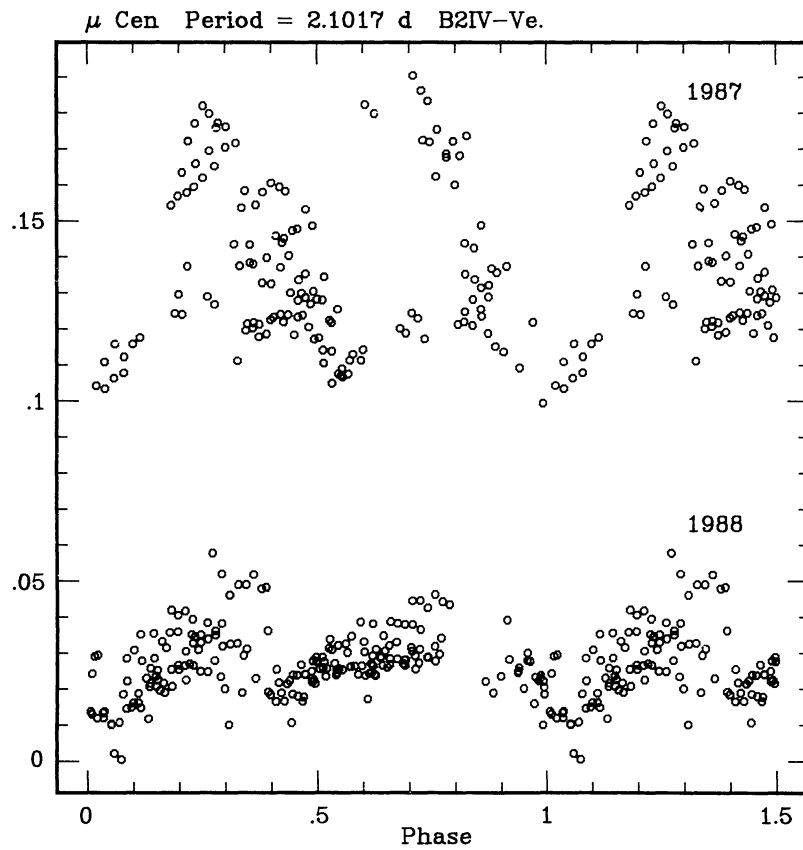
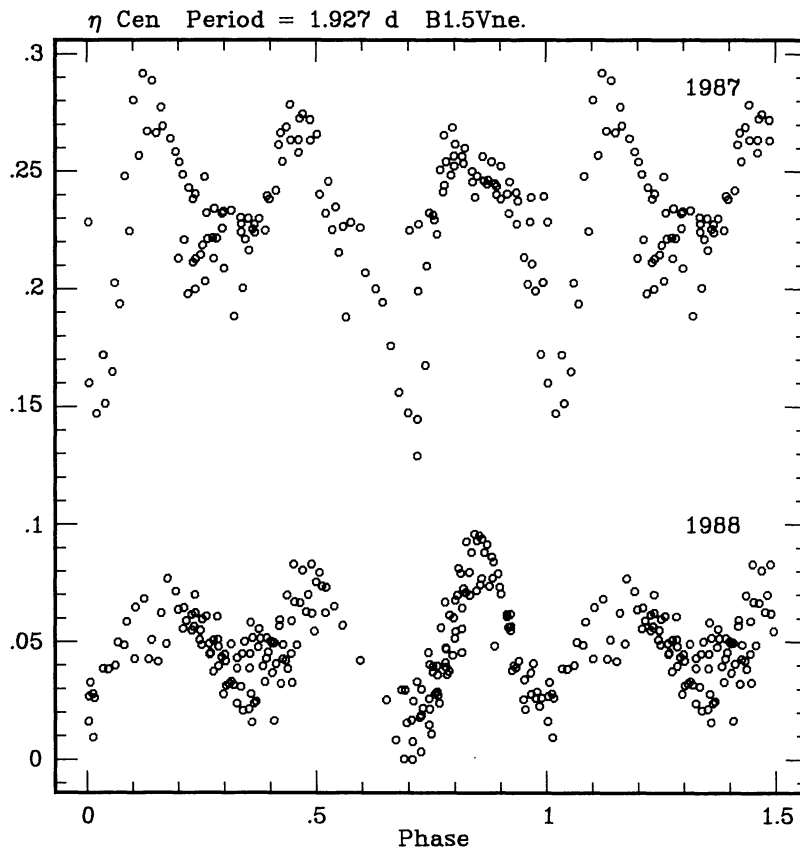
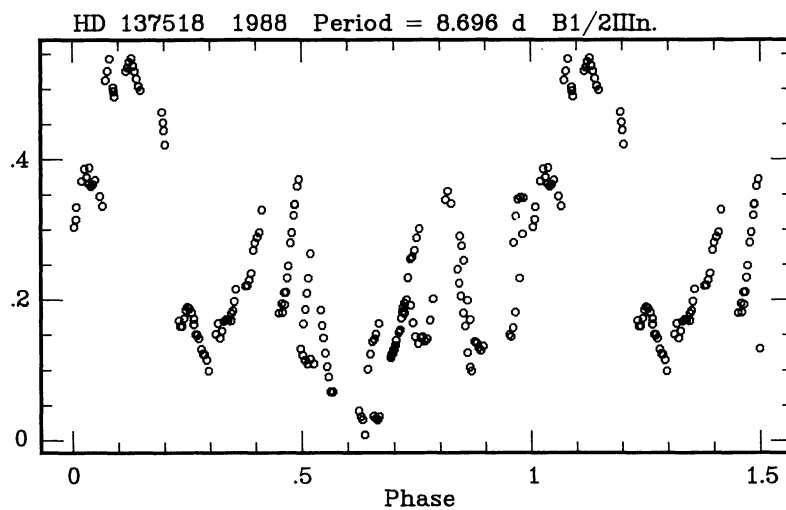
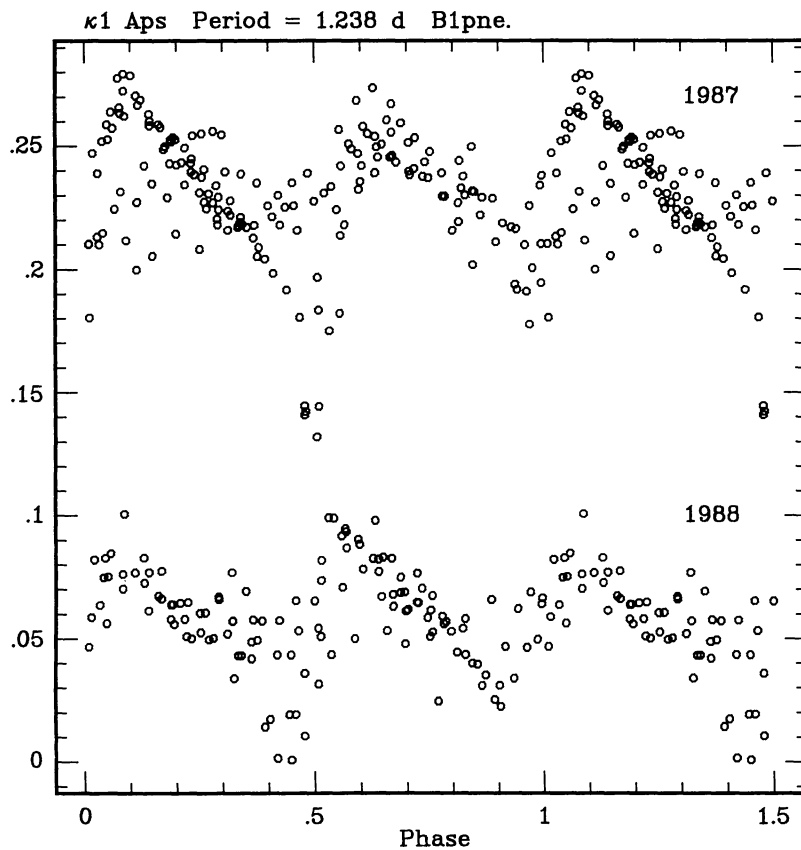
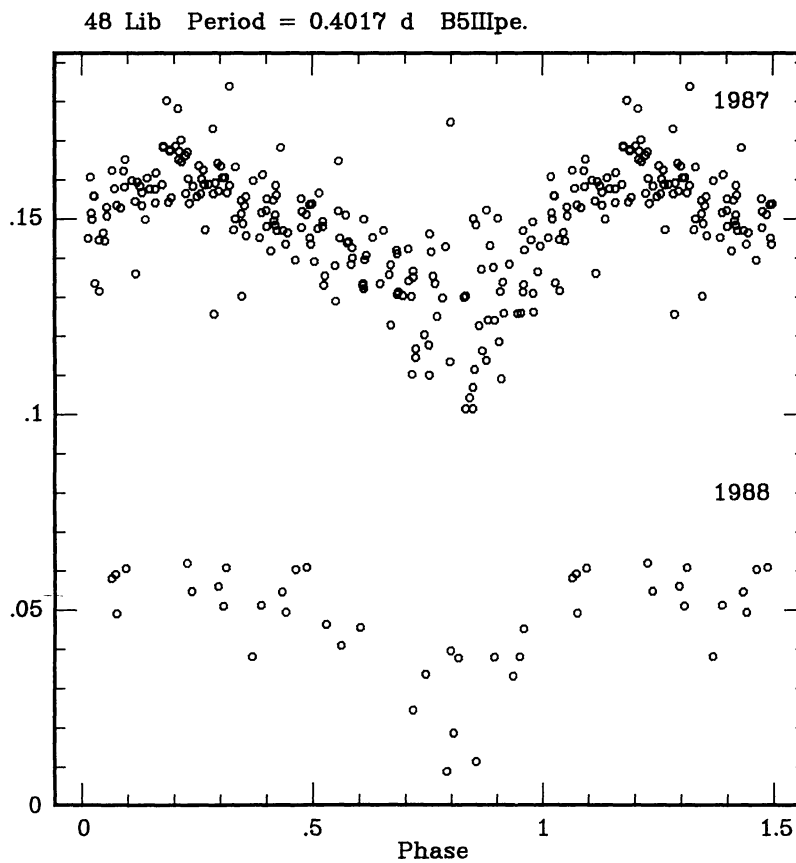


FIGURE 2. — Light curve of μ Cen ($P = 2.1017$ d).

FIGURE 3. — Light curve of η Cen ($P = 1.927$ d).FIGURE 4. — Light curve of HD137518 ($P = 8.696$ d).

FIGURE 5. — Light curve of κ^1 Aps ($P = 1.238$ d).FIGURE 6. — Light curve of 48 Lib ($P = 0.4017$ d).

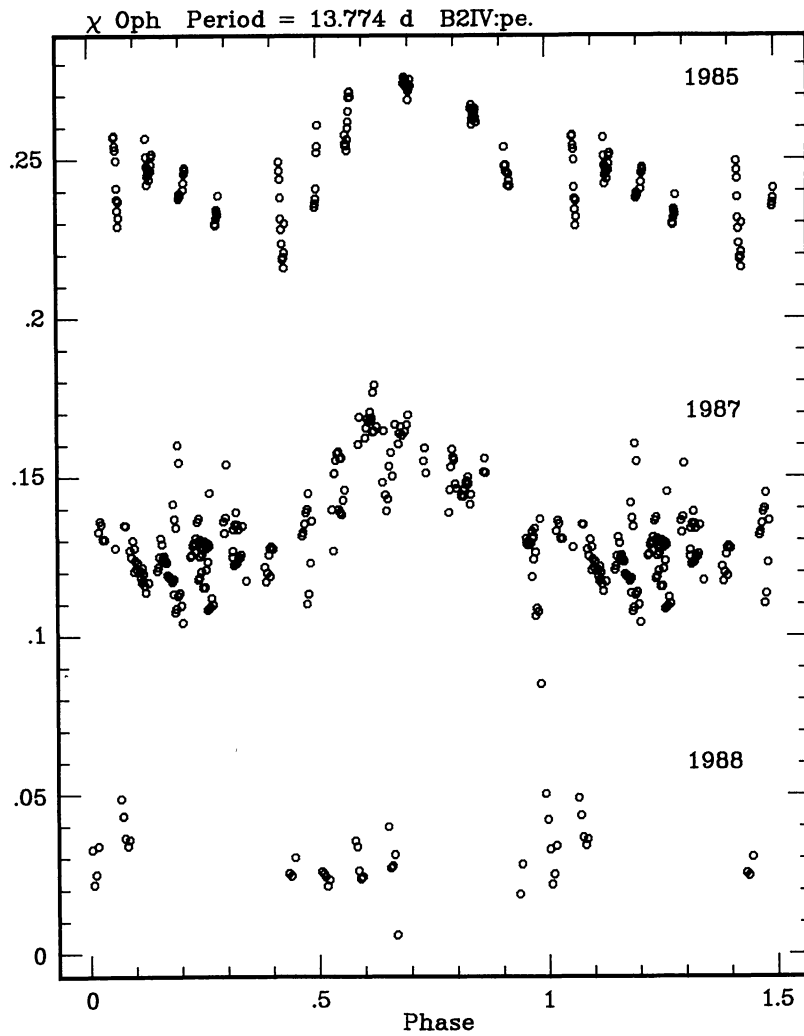


FIGURE 7. — Light curve of χ Oph ($P = 13.774$ d).

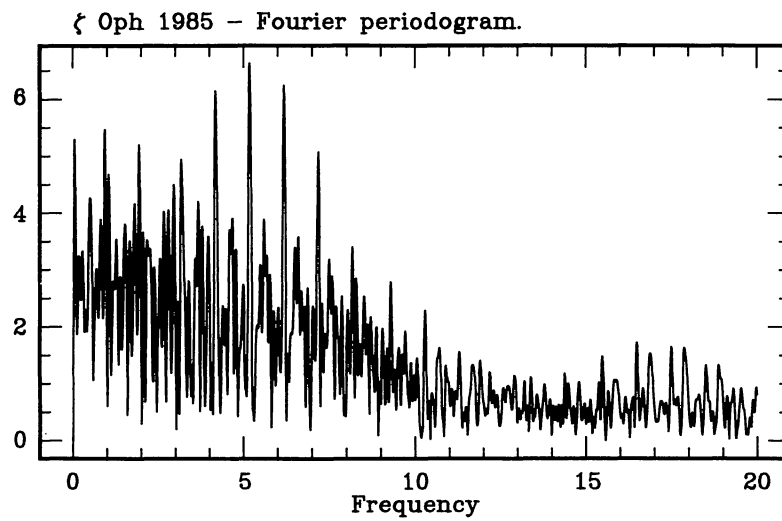
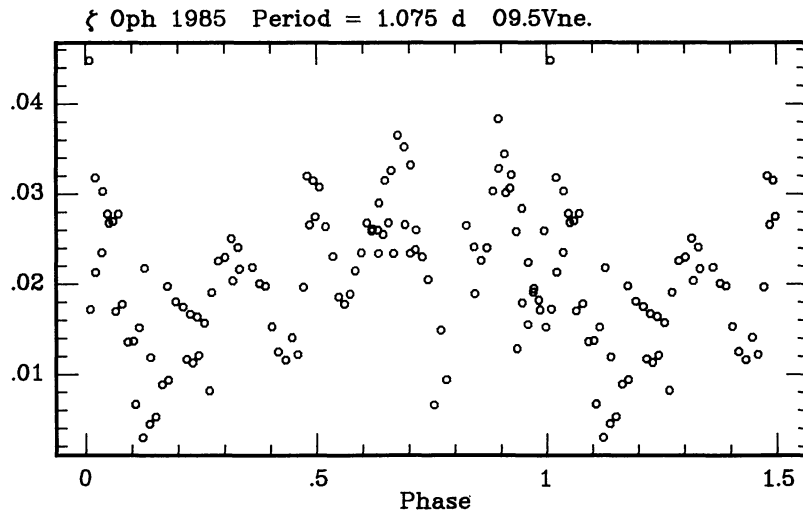
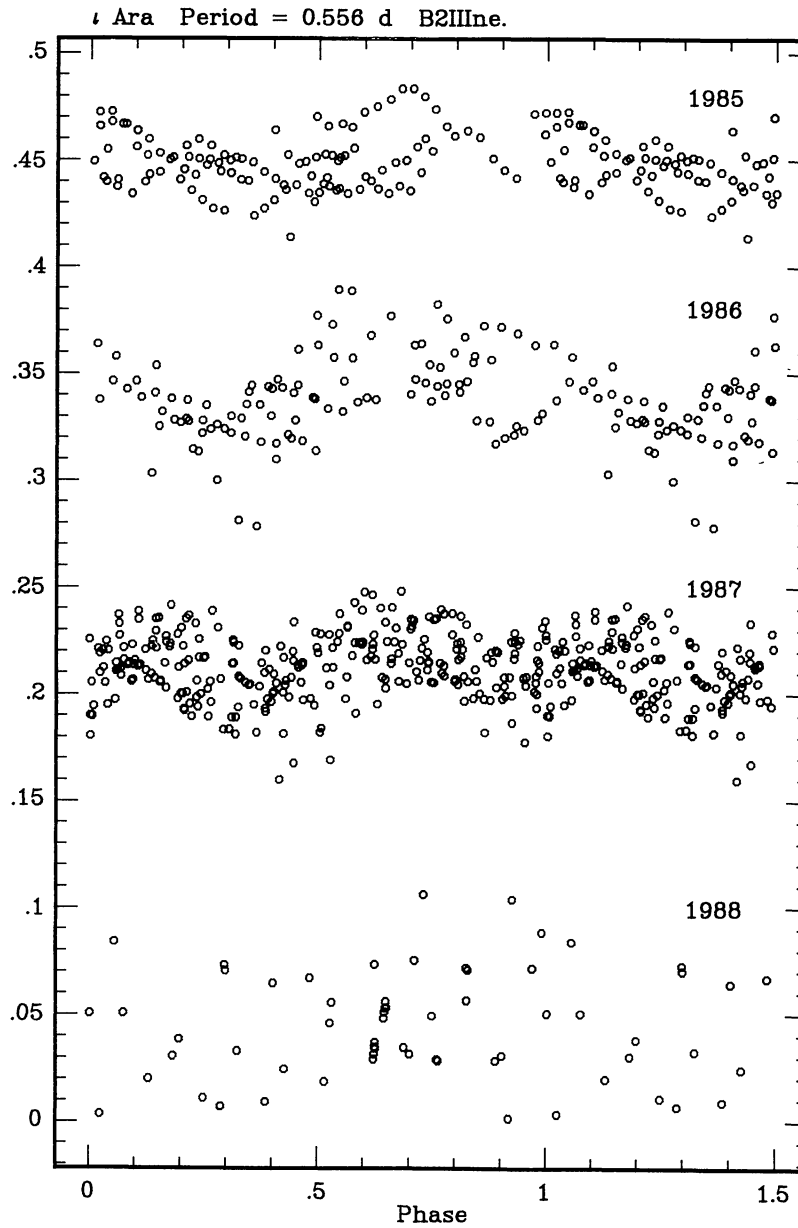


FIGURE 8a. — Fourier periodogram for 1985 data of ζ Oph. The frequency is in cycles d^{-1} and the semi-amplitude in millimagnitudes.

FIGURE 8b. — Light curve of ζ Oph for 1985 ($P = 1.075$ d).FIGURE 9. — Light curve of ι Ara ($P = 0.556$ d).

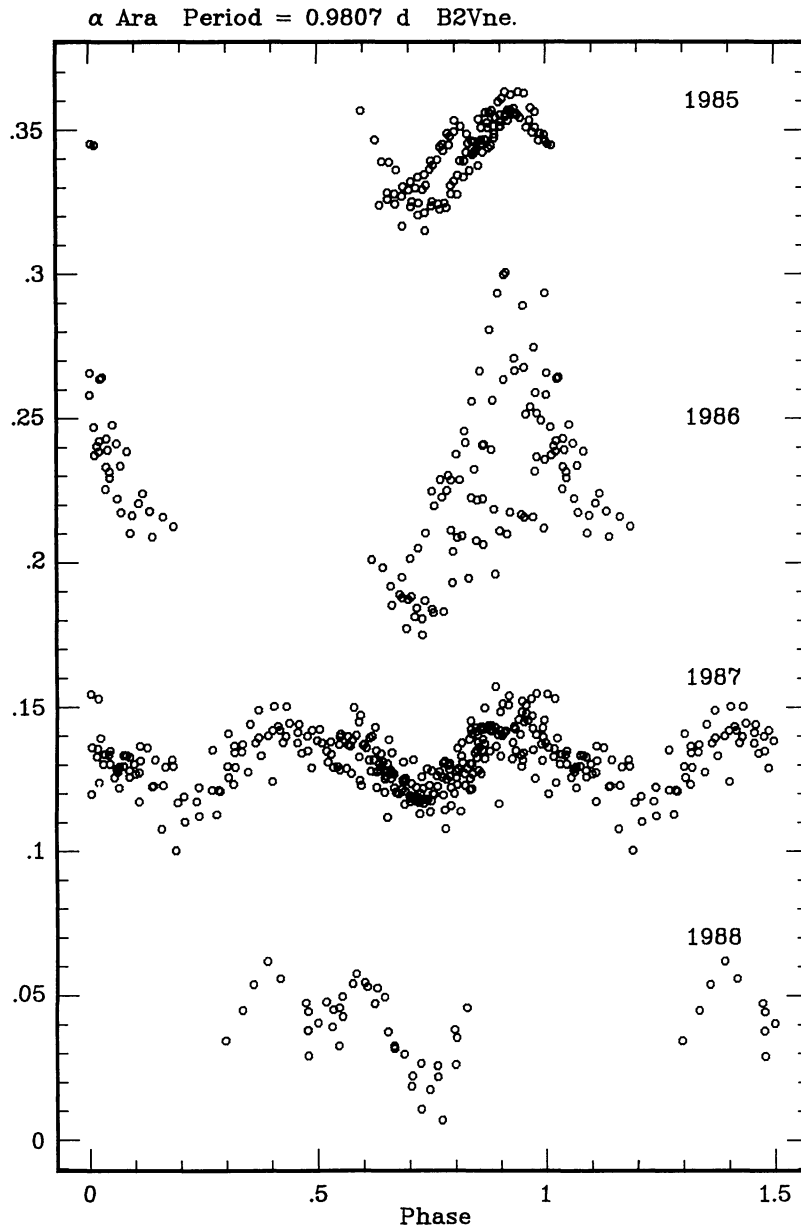


FIGURE 10. — Light curve of α Ara ($P = 0.9807$ d).

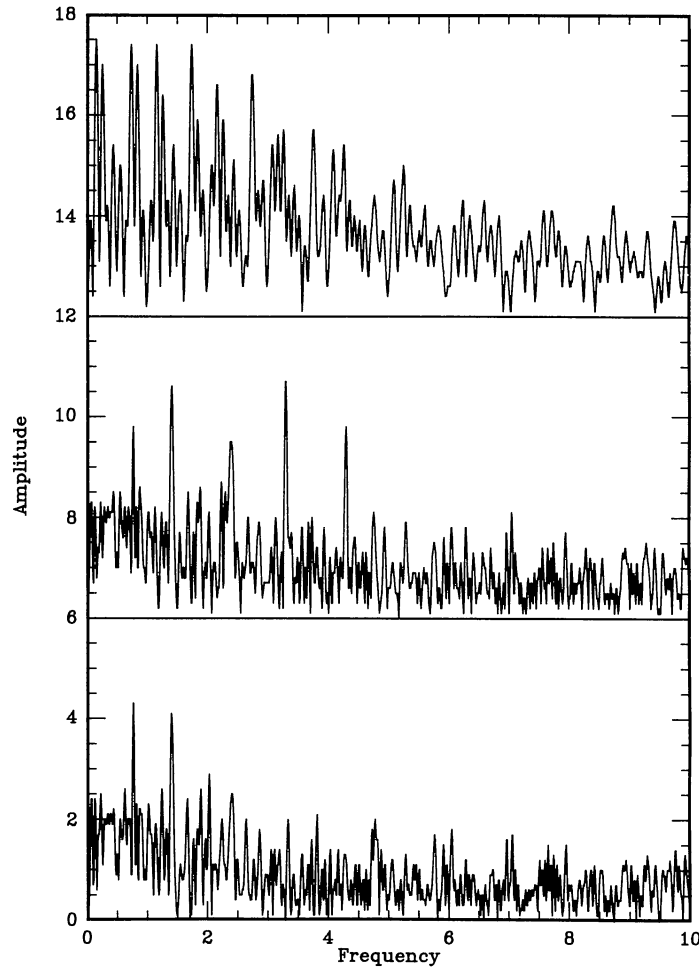


FIGURE 11a. — Fourier periodograms for V986 Oph. Top panel - 1985 data ; middle panel - 1987 data ; bottom panel - 1987 data prewhitened by 3.30 d^{-1} . The frequency is in cycles d^{-1} and the semi-amplitude in millimagnitudes.

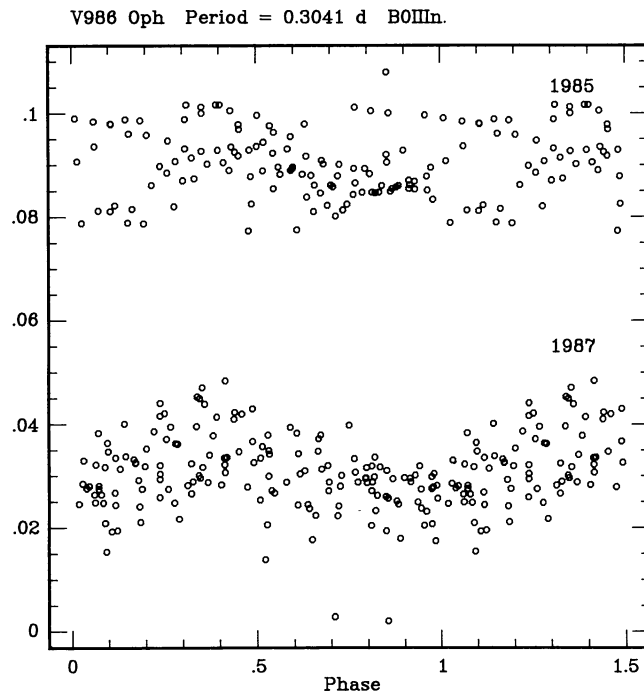
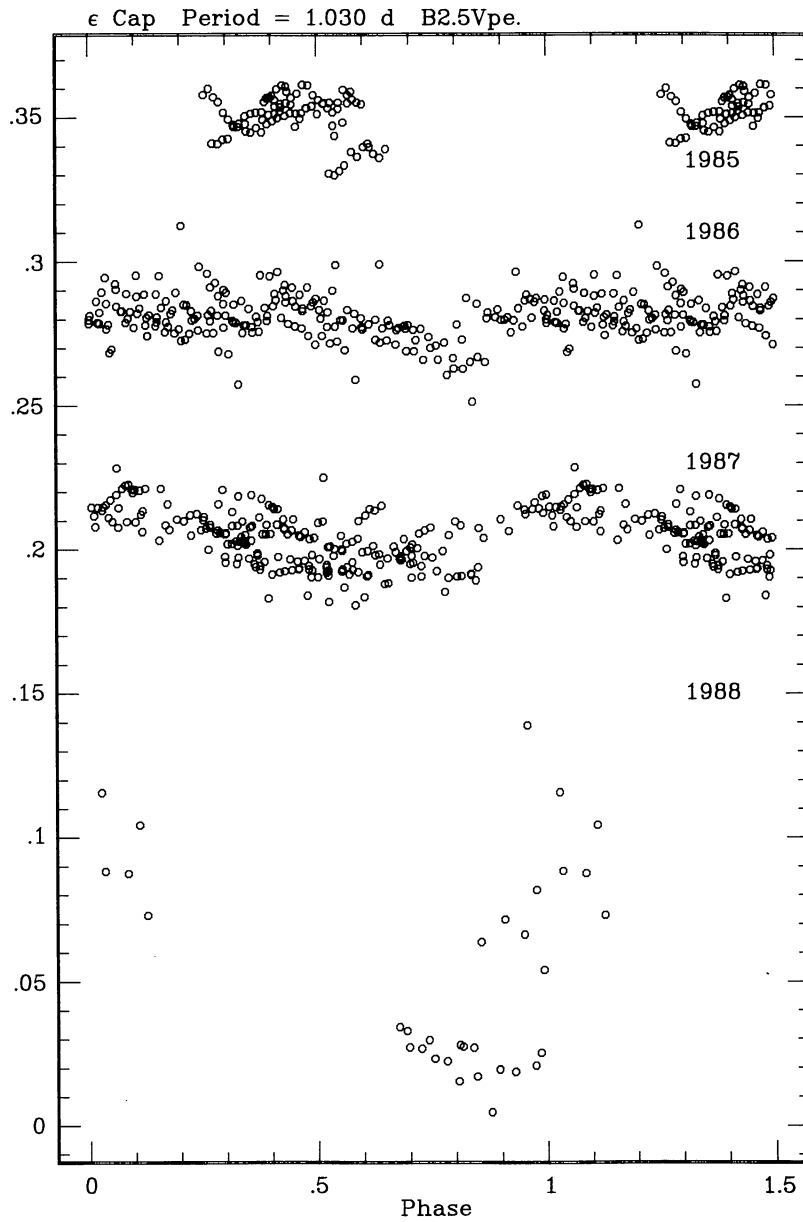


FIGURE 11b. — Light curve of V986 Oph ($P = 0.3041 \text{ d}$).

FIGURE 12. — Light curve of ϵ Cap ($P = 1.030$ d).

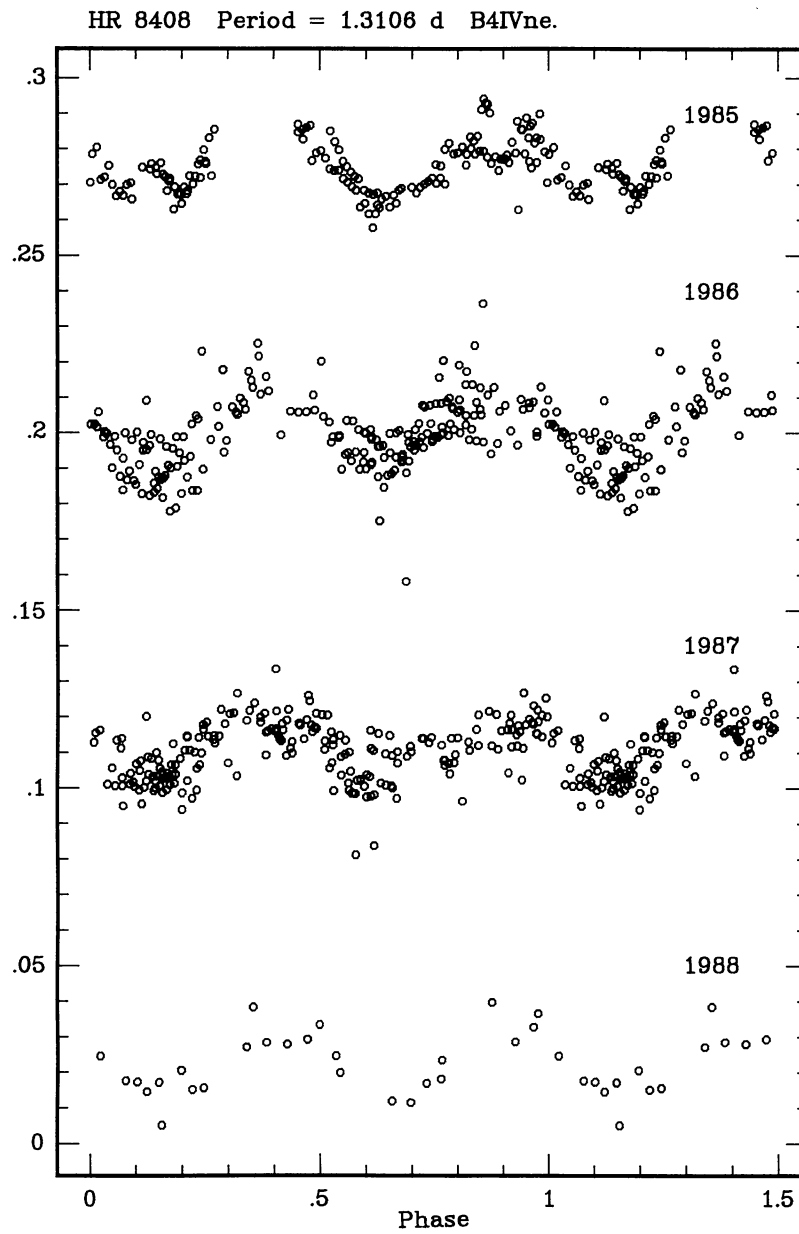
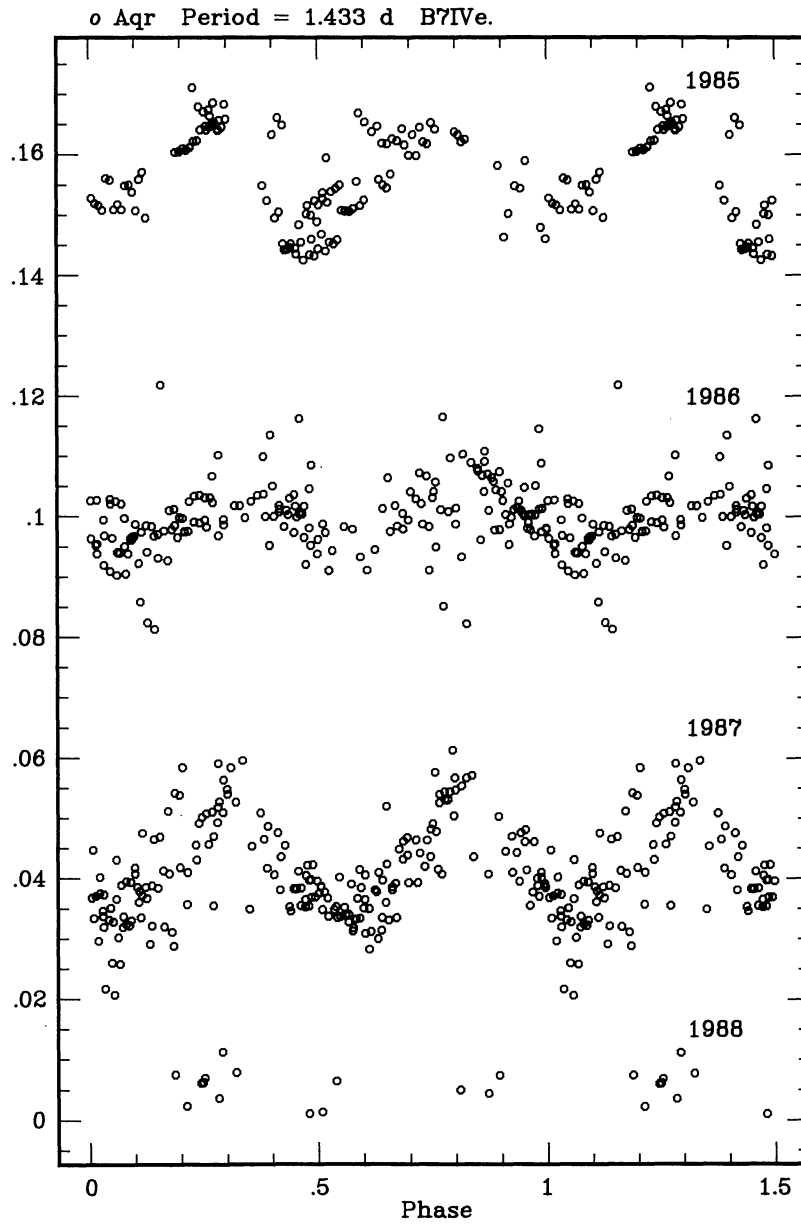
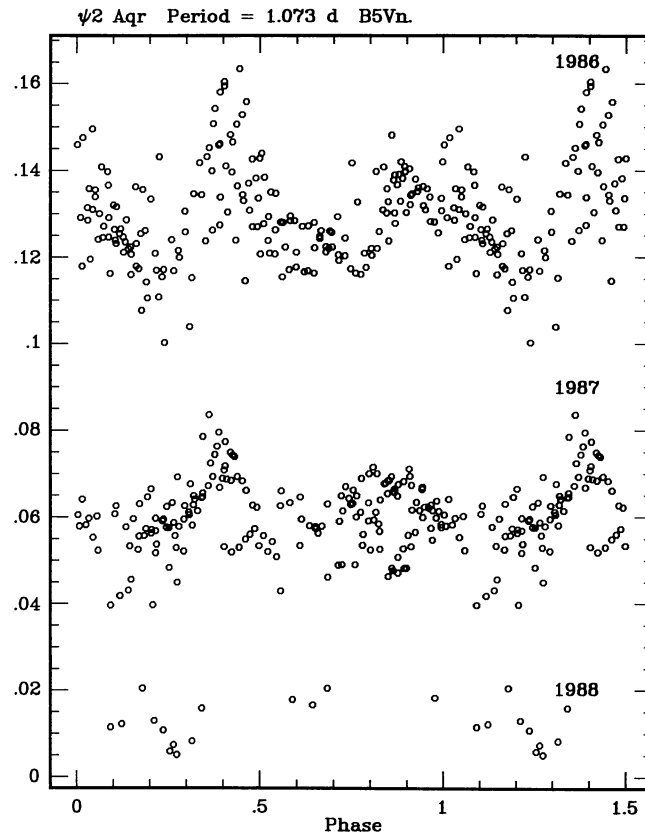
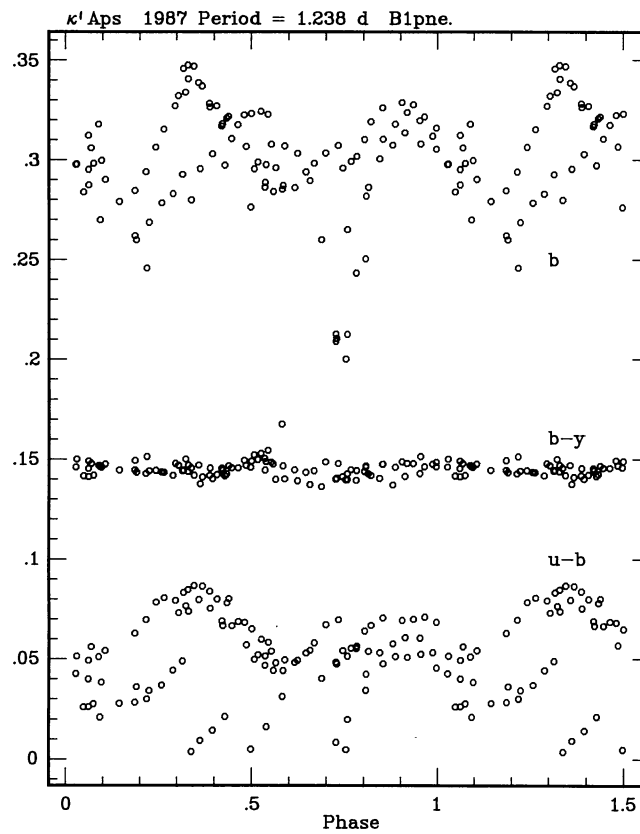
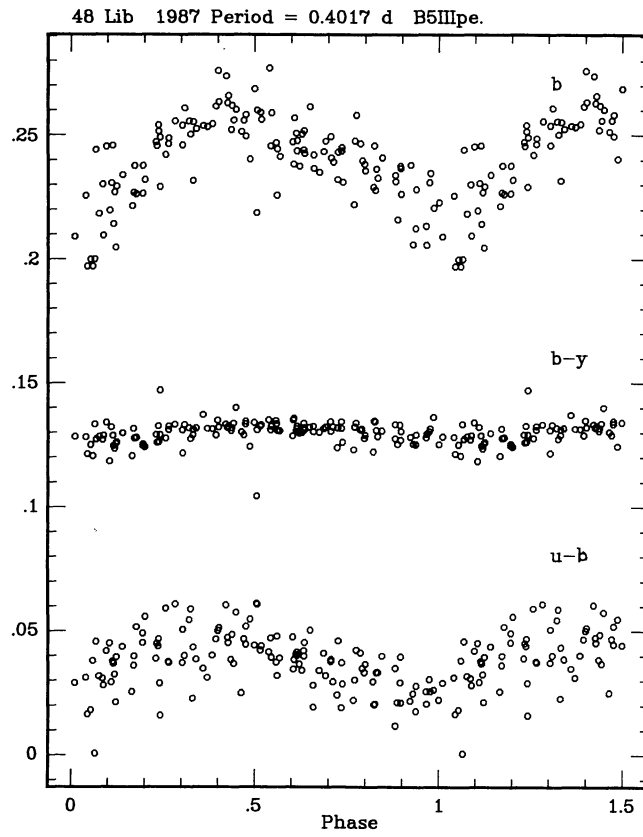
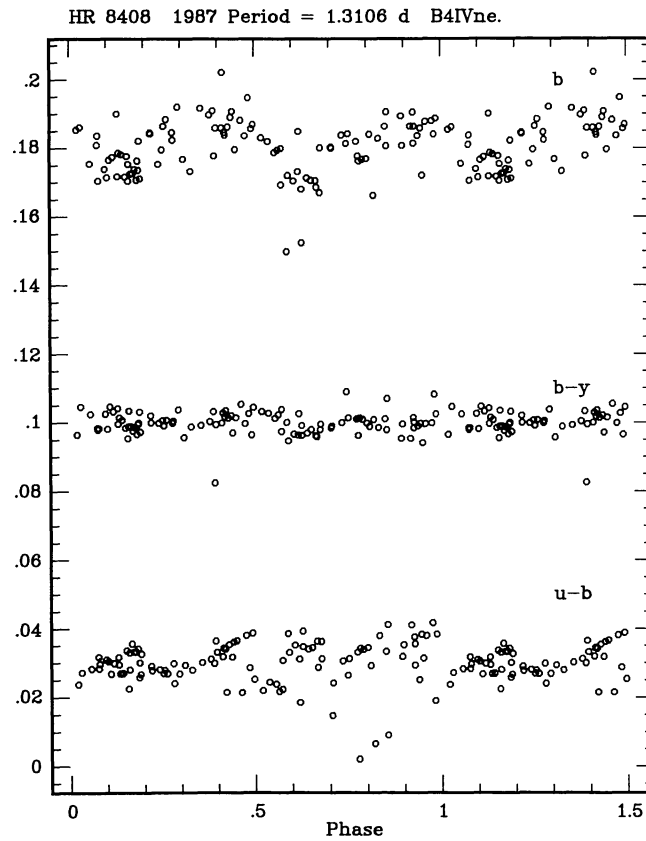
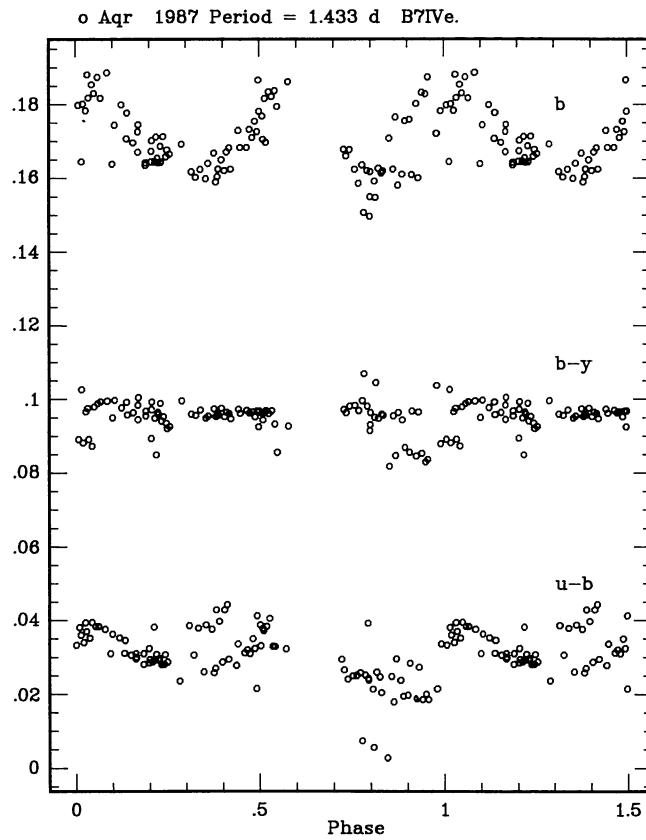
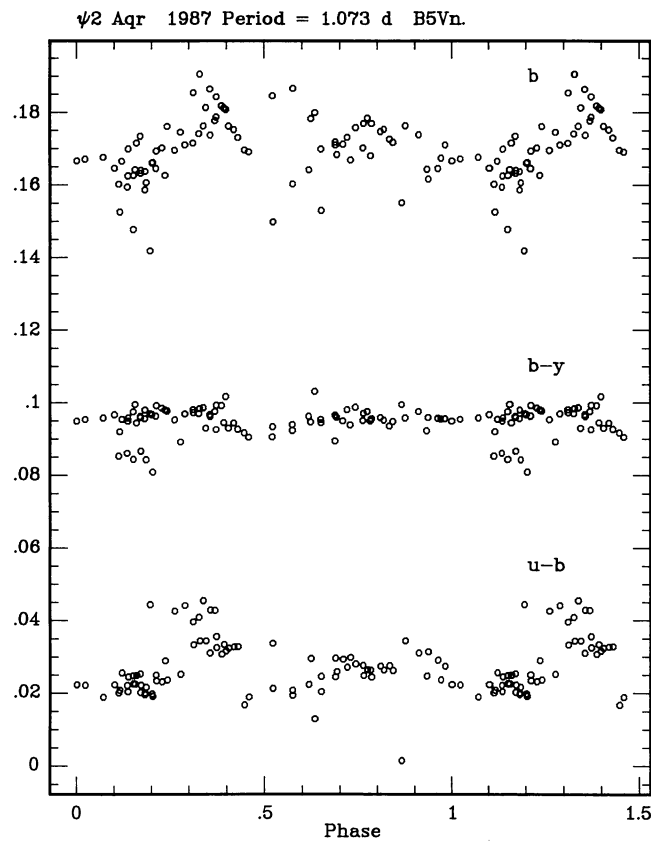


FIGURE 13. — Light curve of HR 8408 ($P = 1.3106$ d).

FIGURE 14. — Light curve of o Aqr ($P = 1.433$ d).

FIGURE 15. — Light curve of ψ^2 Aqr ($P = 1.073$ d).FIGURE 16. — Light and colour variations in κ^1 Aps ($P = 1.238$ d). In this and subsequent figures the colour and light variations are plotted according to the astronomical convention, i.e. with more negative values (bluer colour) upwards.

FIGURE 17. — Light and colour variations in 48 Lib ($P = 0.4017$ d).FIGURE 18. — Light and colour variations in HR 8408 ($P = 1.3106$ d).

FIGURE 19. — Light and colour variations in o Aqr ($P = 1.433$ d).FIGURE 20. — Light and colour variations in ψ^2 Aqr ($P = 1.073$ d).