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SUPERGIANT AND GIANT M TYPE STARS IN THE LARGE MAGELLANIC CLOUD (*)

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Summary.-We present coordinates, finding charts, and near infrared magnitudes, I , for near 500 probable M supergiants and over 600 M giants in the field of the Large Magellanic Cloud. Photoelectrically observed colours, $R-I$, are given for 208 objects and $V-R$ for 144 objects. The catalogues give also the numbers of stars appearing in the Sanduleak-Philip list as well as various other designations.

Key words : supergiant and giant M stars - Large Magellanic Cloud - V , R , I photometry.

1. Introduction.- A survey of the Large Magellanic Cloud (LMC) was carried out in the late 50s by Westerlund with the 50/65/175 cm Schmidt telescope of Uppsala Southern Station at Mount Strömlo Observatory, using objective-prism techniques in the near infrared spectral region. Over 300 probable carbon stars were identified and several hundred M type stars were classified. Preliminary results were published (Westerlund 1960, 1961, 1964). Recently after a confirmation of the objective-prism identifications of some of the carbon stars by slit spectroscopy, a catalogue of the identified carbon stars was presented (Westerlund et al., 1978). In the mean time also a number of our identified M supergiants have been studied in higher dispersion (Humphreys 1974, 1979), and a large number of M giants in the LMC have been detected (Blanco et al., 1975, 1978). We have ourselves obtained slit spectrograms of several tens of M stars in the LMC and carried out photoelectric VRI photometry for about 200 of our objects. We find it, therefore, appropriate to present a catalogue of probable M supergiants in the LMC - this catalogue is believed to be rather complete. We present also a catalogue of probable M giants in the LMC. This catalogue should be looked upon as giving a good sample of the giant M stars in the region, brighter than $I = 13$, but it is far from complete even in this magnitude range. The detailed discussion of our material is postponed until the analysis of our slit spectrograms of the M stars has been completed.

2. Observations.- The photographic material used for the survey and for the determination of the I magnitudes for stars not observed photoelectrically is identical to that used by Westerlund et al. (1978) for cataloguing the carbon stars in the region; details may be found in that paper.

The photoelectric photometry in the VRI bands has been carried out by Westerlund, Olander and Lundgren in 1975-78 with the 1 m telescope at the European Southern Observatory

(*) Partly based on observations collected at the European Southern Observatory, La Silla, Chile.

(ESO) on La Silla, Chile, and to a limited extent, by Olander, with the 61 cm telescope of the Bochum University on La Silla. The photoelectric standards were taken from the Johnson et al. list (1966), and in 1978 also from the Kunkel and Rydgren RI standards (1978) among the Landolt equatorial UBV stars (1973). The observations have been carried out partly with an ITT FW 118 tube, partly with an RCA C 31034 tube. We have not encountered any difficulties in transferring our observations to the Johnson VRI system. However, the uncertainties discussed by Humphreys (1979) may exist in our investigations, too, for the reddest stars.

Many, may be most, of the M supergiants are variable, and we find it therefore difficult to give a meaningful value on the external accuracy of the catalogued data. The comparisons available in table III should suffice to give a feeling for the reliability of the magnitudes and colours for stars of these types.

We note also that our colour indices are in the range to be expected for M0-M4I stars with little or no reddening (Lee, 1970).

3. The Catalogues.- The classification of the M stars from the low-dispersion objective-prism spectra in the near-infrared, 2200 Å/mm at the A-band, follows closely the Case system (Nassau and Albada, 1949). All the probable M supergiants are of early subtypes and occupy a relatively limited interval of magnitude. The giant M stars are of later types and generally fainter; we believe that all the bright stars are galactic foreground stars.

Table I gives the catalogued data for 532 possible M supergiant stars in the LMC. The columns contain in order: the star number, the number of the identification chart, the coordinates (1975), the spectral sub-type, the photometric data I , $R-I$, $V-R$. The column S and P gives the numbers used by Sanduleak and Philip (1977) in their table I. The final column gives other designations of the stars.

Photoelectrically measured I magnitudes are given with two decimals, photographically measured ones are given with one. The error in a photoelectric magnitude is estimated to be ± 0.05 and the error in a photographic magnitude ± 0.2 .

During the observations with the Bochum telescope a colour system V, R, R_2 was used, where R_2 is located between R and I . For these stars $V-R$ is given with two decimals in the Catalogue, whereas I and $R-I$, as deduced quantities, are given with one.

Most of the stars in the catalogues having I and $R-I$ but no $V-R$ were observed photoelectrically by Mianes. We are grateful to him for having received these data in advance of publication.

Table II gives the catalogued data for 655 giant M stars in a similar way. We have, however, not included here spectral subtypes; the majority of the objects are later than M4.

We wish to emphasize that the separation of the stars into supergiants and giants follows primarily from their spectral types. Stars classified M4 or earlier are consequently generally included in table I. In some cases their apparent magnitudes contradict a supergiant character by being too faint, or, occasionally, too bright, and table I may thus contain a few foreground stars, easily recognisable (see Notes to tables I and II).

Figure 1 gives the division of the LMC into "finding chart" fields, and the finding charts are presented in figure 2. Here, an M supergiant is identified by one line towards the stellar image; an M giant has its stellar image in between two lines.

4. Comments.- 4.1. The distributions of the stars in magnitude.- Figure 3 gives the distributions of the stars from table I (full-drawn line) and from table II (dotted line) in magnitude. It appears certain that most of the early M stars with $8.5 - I - 12$ are supergiant stars in the LMC. Those outside these limits are probably galactic foreground stars or, among the faint ones, giants in the LMC. Obviously, the late M stars brighter than $I = 11$ must be considered as galactic foreground stars.

A comparison with the results of Blanco et al. (1978) for the giant stars appears to indicate that the cut-off in our material comes just before the peak in their distribution is reached (provided our magnitude scales are comparable). The very steep increase in numbers at $I = 12$ is probably real in spite of the fact that our material is incomplete. We have previously concluded (Westerlund, 1960) that the absolute magnitudes of the detected red giant stars in the LMC, of about $M_T \sim -6$, indicate that they belong to Population II; this is confirmed by ours as well as the Blanco et al. results.

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4.2. Individual data.- Table III gives the individual photoelectric values and the dates (month) of their observations for M supergiants (Table IIIa) and M giants (Table IIIb) observed more than once. Here we include the observations by Mianes (October 76) and by Humphreys (1979). The latter are identified in our table by dates having asterisks. Many, perhaps most, stars are seen to be variables. Nevertheless, a reasonable agreement between the various series of observations may be noted.

For information, we give in table IV the dates (months) of the photoelectric observations of stars observed in V, R, I during one period, only.

4.3. The spectral classification.- Sanduleak and Philip (1977) give 609 suspected late-type supergiant stars in their table I. Of these we have identified 209 among our supergiants and 12 among our giant M stars. We believe that the remaining 388 of their stars are earlier than M0 in our classification system. It is not quite clear why over 300 of our stars do not appear in the Sanduleak-Philip list. The area searched by us is somewhat larger than that searched by them; our survey extends from R.A. $\sim 4^{\text{h}}15^{\text{m}}$ to $6^{\text{h}}30^{\text{m}}$ and from Decl. $\sim -64^{\circ}$ to -74° , whereas their survey covers the region between R.A. $\sim 4^{\text{h}}30^{\text{m}}$ and $6^{\text{h}}10^{\text{m}}$ and Decl. $\sim -65^{\circ}5$ and -74° . However, only 28 of the stars in our table I fall outside the Sanduleak-Philip boundaries, and of those we have listed 6 as likely foreground stars. It is true that our survey extends to fainter magnitudes and also includes about 50 foreground stars (cf. notes to table I and II), but the majority of our supergiant stars are in the interval $9 - I - 11$ which should be covered by them. Possibly the identification of weak TiO bands in the near infrared is easier than the determination of a pronounced gradient in the continuum of these stars; the crowding of images on the deep IIIa-J plates may also contribute to the difficulties in identifying them.

We have also compared our classification from the objective-prism plates with that by Humphreys (1979) based on slit spectra. The result is summarized in table V. The agreement is reasonable when considering the different techniques used; our classification appears earlier than Humphreys' by about one subdivision.

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TABLE II.- (Continued)

NO	CHART	R.A. (1975)	DECL.	I	R-I	V-R	S&P	OTHER	NO	CHART	R.A. (1975)	DECL.	I	R-I	V-R	S&P	OTHER
601	8:14	5 51 53.3	-70 36 22.	13.1					631	7:5	5 55 52.9	-65 35 28.	14.06	0.20			
602	8:2	5 51 53.4	-67 35 59.	12.4					632	8:10	5 56 0.5	-69 44 22.	12.2				
603	7:9	5 52 14.1	-66 29 37.	11.1			59:13		633	7:8	5 50 1.4	-66 14 23.	13.1				
604	9:3	5 52 15.9	-71 5 0.	12.0					634	8:1	5 50 4.8	-67 41 22.	10.0				61:39
605	7:3	5 52 24.5	-64 53 2.	11.6					635	8:7	5 56 12.5	-69 11 14.	12.9				
606	7:6	5 52 28.4	-65 40 14.	13.0					636	8:7	5 56 12.9	-69 18 22.	13.3				
607	9:1	5 52 31.0	-70 57 20.	12.7					637	8:7	5 56 44.9	-69 6 4.	12.0				
608	7:3	5 52 38.7	-64 57 20.	10.1					638	5:6	5 56 45.9	-73 8 53.	13.0				
609	7:3	5 52 57.9	-65 2 38.	10.9					639	8:4	5 56 54.2	-68 27 34.	11.5				
610	7:11	5 52 59.1	-66 54 38.	12.4					640	9:11	5 57 29.4	-73 31 6.	8.96	1.70			
611	7:11	5 53 12.4	-67 4 22.	9.7					641	7:2	5 57 33.2	-65 25 25.	12.4				
612	7:6	5 53 18.0	-65 38 22.	11.0					642	7:8	5 59 21.0	-66 31 59.	10.4				
613	7:11	5 53 18.4	-66 53 35.	12.8					643	7:5	5 59 33.4	-65 56 58.	13.1				
614	8:14	5 53 37.7	-70 34 9.	13.48	0.61	0.30			644	9:6	5 59 36.2	-72 7 57.	12.7				
615	8:11	5 53 38.1	-69 35 23.	12.6					645	9:11	6 0 15.7	-73 17 47.	13.1				
616	7:2	5 53 42.0	-65 24 40.	12.7					646	8:1	6 1 23.1	-67 25 29.	11.7				
617	7:2	5 53 46.0	-65 13 36.	13.1					647	8:1	6 3 28.4	-67 47 8.	11.7				
618	9:3	5 53 49.9	-71 34 20.	12.2					648	5:6	6 3 39.5	-72 24 48.	13.0				
619	7:2	5 53 51.0	-65 17 0.	12.5					649	9:8	6 5 19.2	-73 0 53.	13.0				
620	7:9	5 53 59.3	-66 10 0.	12.8					650	9:8	6 5 53.5	-72 39 39.	13.1				
621	7:6	5 54 9.9	-66 8 53.	12.8					651	9:8	6 6 4.0	-73 0 13.					
622	7:2	5 54 33.2	-64 56 32.	7.7					652	9:11	6 6 26.8	-73 22 26.	9.2				
623	4:1	5 54 46.0	-71 2 29.	9.45	1.03				653	9:10	6 8 10.2	-73 30 49.	9.7				
624	7:5	5 54 51.3	-65 38 57.	13.1					654	9:5	6 9 30.8	-72 11 9.	13.1				
625	7:2	5 55 12.3	-65 4 52.	12.7					655	9:10	6 10 4.1	-73 41 9.	12.7				
626	7:2	5 55 19.9	-65 4 18.	12.0					656	9:5	6 11 41.1	-72 10 56.	13.1				
627	8:1	5 55 24.0	-67 30 58.	12.8					657	9:5	6 11 57.3	-72 13 16.	5.8				
628	8:7	5 55 36.1	-69 16 14.	13.0					658	9:5	6 13 59.7	-72 14 41.	7.0				
629	8:10	5 55 36.9	-69 35 25.	11.5					659	9:5	6 14 1.1	-72 14 30.	13.3				
630	7:8	5 55 47.7	-66 42 11.	9.5					660	24:3	6 20 18.6	-72 26 52.	12.5				

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TABLE III a.- Photoelectric observations of supergiant M stars in the LMC.

No	I	R-I	V-R	Date	No	I	R-I	V-R	Date
6	10.72	1.37	1.90	Dec. 76	168	9.99	1.28		Oct. 76
	10.48	1.47	1.66	Jan. 78		10.12	1.25	1.71	Jan. 78
						10.11	1.61	1.83	Oct. 78
8	10.21	0.95	1.30	Dec. 76	193	9.98	1.63	1.98	Dec. 76*
	10.14	0.96	1.37	Jan. 78		9.38	1.37	1.64	Dec. 77*
						9.25	1.85	1.88	Oct. 78
46	10.32	1.19		Oct. 76	194	9.82	1.71		Oct. 76
	10.26	1.02	1.47	Jan. 78		9.89	1.69	1.87	Dec. 76*
49	10.09	1.22	1.75	Dec. 76		9.66	1.85	1.84	Dec. 77*
	9.74	1.35	1.75	Jan. 78					
51	10.39	1.49		Oct. 76	200	10.40	1.20	1.35	Dec. 76*
	10.38	1.48	1.81	Jan. 78		10.14	1.36	1.61	Dec. 77*
						10.18	1.25	1.33	Oct. 78
52	9.94	1.52		Oct. 76	234	9.08	1.36	1.66	Dec. 76
	10.20	1.54		Dec. 76		9.26	1.26	1.73	Dec. 76*
	10.19	1.57	1.79	Jan. 78		8.98	1.27	1.67	Dec. 77*
						9.16	1.11	1.69	Jan. 78
69	9.97	1.69	2.33	Dec. 76	239	9.66	0.63		Oct. 76
	9.49	1.61	1.89	Jan. 78		9.62	0.64	0.95	Dec. 76
						9.66	0.63	0.91	Oct. 78
74	10.16	2.06		Oct. 76	245	10.39	1.24	1.94	Nov. 75
	10.53	1.85		Dec. 76		10.18	1.17	1.55	Dec. 76*
	9.81	1.76	2.15	Jan. 78		9.98	1.36	1.63	Dec. 77*
80	9.92	1.14		Oct. 76	258	7.60	1.36		Oct. 76
	9.87	1.22	1.54	Dec. 76		7.51	1.40	1.54	Dec. 76
86	9.71	1.29		Oct. 76	289	9.54	1.34	1.54	Dec. 76*
	9.64	1.35	1.63	Dec. 76		9.23	1.37	1.56	Dec. 77*
89	10.23	1.11		Oct. 76		9.34	1.34	1.60	Jan. 78
	10.41	0.96	1.64	Dec. 76	296	9.06	1.06		Oct. 76
						9.18	1.09	1.32	Dec. 76
91	9.57	1.19		Oct. 76		9.25	1.13	1.29	Oct. 78
	9.78	1.10	2.12	Dec. 76	297	9.92	1.62	1.69	Dec. 76
						10.16	1.51	1.84	Jan. 78
95	9.86	1.64		Oct. 76	299	10.26	1.30	1.60	Dec. 76*
	9.64	1.35	1.63	Dec. 76		9.93	1.19	1.52	Dec. 77*
	9.95	1.47	1.65	Dec. 76*		10.06	1.13	1.50	Jan. 78
	9.88	1.71	1.92	Dec. 77*					
161	10.03	1.69	2.01	Dec. 76*					
	9.78	1.56	1.78	Dec. 77*					
	10.10	1.92	1.84	Oct. 78					
166	8.97	1.19		Oct. 76					
	8.85	1.11	1.54	Dec. 76					
	8.91	1.14	1.54	Jan. 78					

TABLE IIIa .-(continued)

No	I	R-I	V-R	Date
319	8.89	1.17	1.61	Dec. 76
	8.99	1.14	1.68	Dec. 76*
	9.15	1.42	1.72	Dec. 77*
	9.31	1.20	1.78	Jan. 78
321	9.95	1.43	1.74	Dec. 77*
	10.09	1.29	1.69	Jan. 78
322	9.75	1.34	1.70	Dec. 76
	9.85	1.30	1.65	Dec. 76*
	9.76	1.54	1.80	Dec. 77*
	9.86	1.33	1.75	Jan. 78
323	9.72	1.39	1.66	Dec. 76
	9.96	1.33	1.71	Dec. 76*
	9.74	1.49	1.75	Dec. 77*
	9.92	1.40	1.70	Jan. 78
329	10.42	1.03	1.42	Dec. 76*
	10.31	1.10	1.42	Dec. 77*
	10.45	1.05	1.21	Oct. 78
330	10.25	1.40	1.64	Dec. 77*
	10.35	1.27	1.62	Jan. 78
331	9.70	1.66	1.94	Dec. 76
	10.16	1.83	2.09	Dec. 77*
	10.27	1.69	2.09	Jan. 78
338	9.06	1.18	1.62	Dec. 76
	9.24	1.17	1.65	Dec. 76*
	8.99	1.25	1.62	Dec. 77*
	9.16	1.13	1.58	Jan. 78
341	8.60	1.19	1.59	Dec. 76
	8.82	1.18	1.64	Dec. 76*
	8.61	1.24	1.58	Dec. 77*
354	9.44	1.49	1.71	Dec. 76
	9.66	1.38	1.80	Dec. 76*
	9.48	1.58	1.80	Dec. 77*
356	10.40	1.20	1.58	Dec. 76*
	10.22	1.33	1.58	Dec. 77*
	10.38	1.43	1.54	Oct. 78
381	10.28	1.32	1.46	Nov. 75
	10.20	1.26	1.54	Dec. 76*
	10.18	1.22	1.51	Dec. 77*
397	10.96	1.04	1.42	Nov. 75
	10.79	1.12	1.54	Dec. 76
	10.81	1.17	1.39	Oct. 78
401	9.99	1.45	1.68	Nov. 75
	11.99	0.73	1.12	Dec. 76*
	11.79	0.84	1.14	Dec. 77*
413	10.18	1.44		Oct. 76
	10.21	1.62	2.11	Jan. 78
	9.84	1.60	1.84	Oct. 78
422	10.23	1.47	1.78	Nov. 75
	9.85	1.26	1.53	Dec. 76
	9.97	1.25	1.62	Dec. 76*
	10.04	1.32	1.58	Dec. 77*
439	10.30	1.33		Oct. 76
	10.34	1.23	1.66	Jan. 78
	10.36	1.51	1.66	Oct. 78
467	9.76	1.52	1.98	Dec. 76*
	9.32	1.64	1.98	Dec. 77*
	9.56	1.58	2.01	Jan. 78
470	9.57	1.38	1.81	Dec. 76*
	9.62	1.35	1.73	Dec. 77*
	9.77	1.27	1.75	Jan. 78
471	9.44	1.52		Oct. 76
	9.67	1.41	1.76	Dec. 76
	9.65	1.50	1.88	Dec. 76*
	9.19	1.35	1.70	Dec. 77*
477	10.08	1.13	1.63	Dec. 76
	10.15	1.33	1.65	Dec. 76*
	9.86	1.23	1.59	Dec. 77*
486	9.71	1.33		Oct. 76
	9.64	1.38	1.65	Dec. 76
	9.78	1.25	1.67	Dec. 76*
	9.58	1.32	1.65	Dec. 77*
	9.78	1.29	1.53	Oct. 78
508	8.88	1.43		Oct. 76
	8.78	1.46	1.55	Dec. 76

TABLE IIIb .-Photoelectric observations of giant M stars in the LMC.

No	I	R-I	V-R	Date
22	8.76	2.22		Oct. 76
	9.32	2.24	2.43	Jan. 78
333	10.20	0.38		Oct. 76
	10.20	0.37	0.56	Dec. 76
	10.29	0.35	0.55	Oct. 78
409	8.73	0.55	0.84	Nov. 75
	9.60	1.37	1.59	Oct. 78
412	10.45	1.28		Oct. 76
	10.50	1.21	1.58	Jan. 78
	10.48	1.37	1.51	Oct. 78
552	9.80	1.29		Oct. 76
	9.73	1.37	1.77	Dec. 76

TABLE IV .- Dates of photoelectric VRI observations of supergiant and giant (g) M stars in the LMC.

	1975		1976		1978		1978	
	Nov.	Dec.	Dec.	Jan.	Jan.	Oct.		
252	5		459	4	453	121		
265	83		502	9	461	148		
267	92		506	225		160		
268	98		507	235		190		
274	112		510	243		204		
279	151		511	247		298		
282	292		514	254		320		
382	301		517	269		411		
387	313			287				
403	335			305				
404	337			318				
409	339			332	g	6		
418	344			424	g	63		
425	440	g	176	435	g	114		
427	450	g	614	447	g	428		

TABLE V .- Comparison of the present (WOH) classification of M supergiants in the LMC with that by Humphreys (RH) from slit spectra.

RH \ WOH	WOH			
	0	0.5	1	2
0	3	1		
0.5	1			
1	1	4	6	
2	2	3	3	
3		1		
4				1

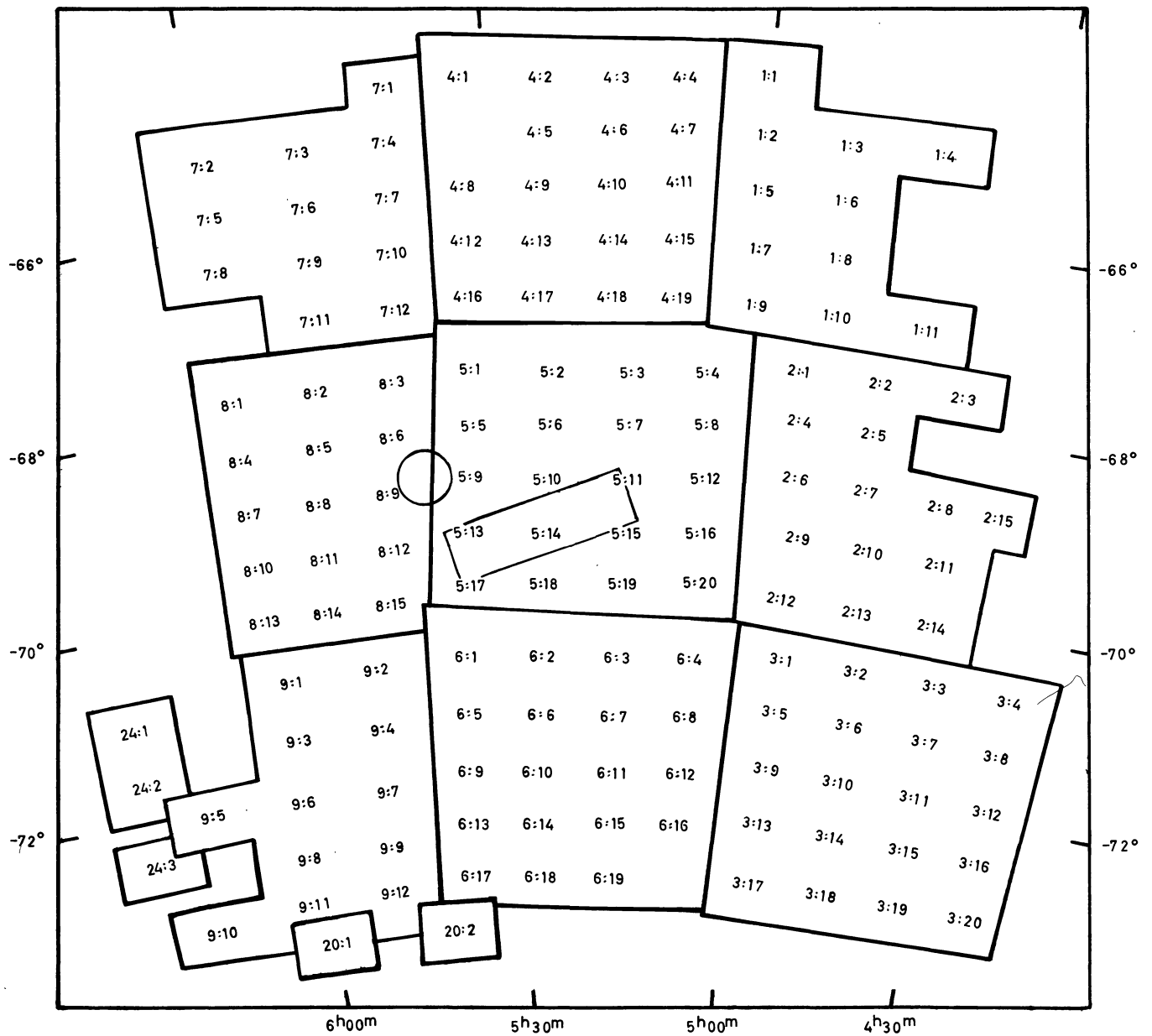
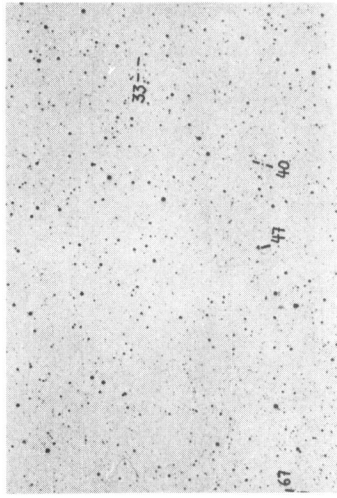
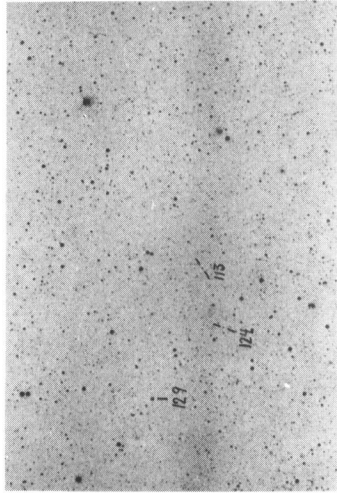


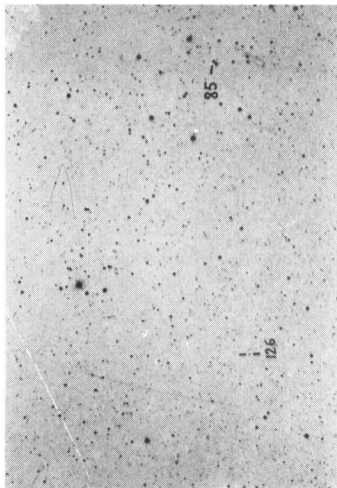
FIGURE 1.- The division of the region into finding-chart fields. For orientation the positions of the Bar (rectangle) of the LMC and of 30 Doradus (circle) are marked.



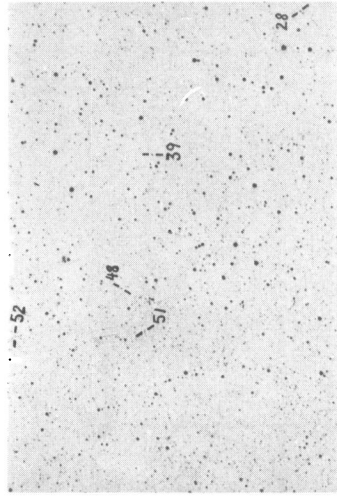
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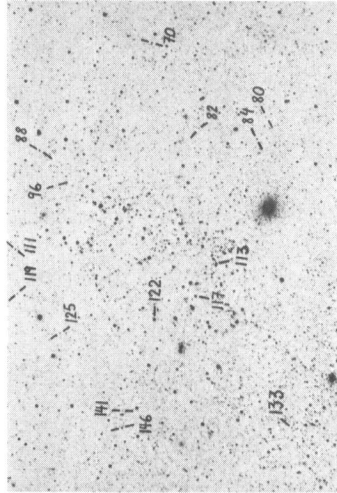
1:2



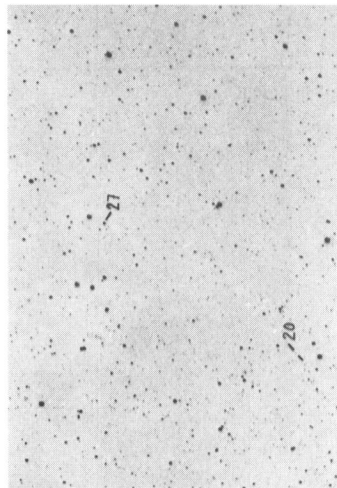
1:3



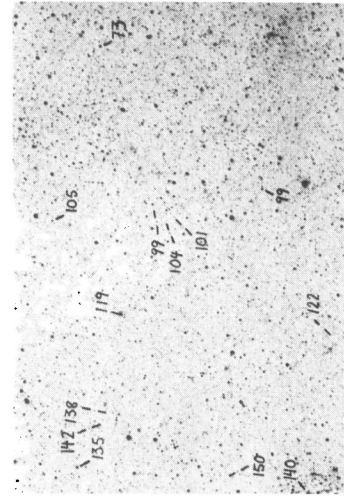
1:4



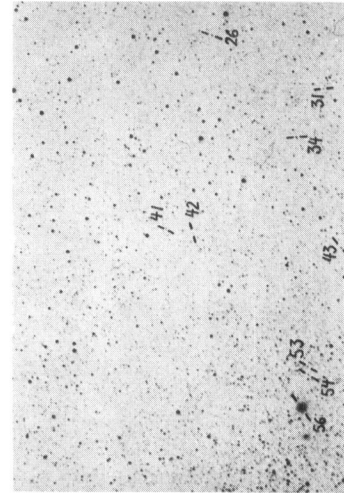
1:5



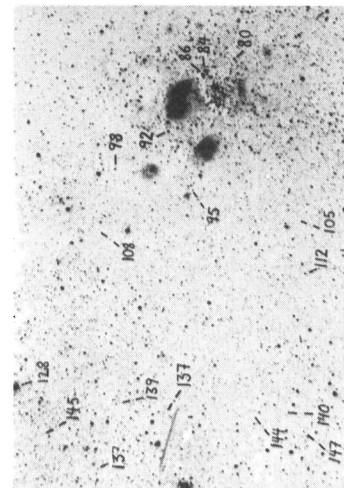
1:6



1:7

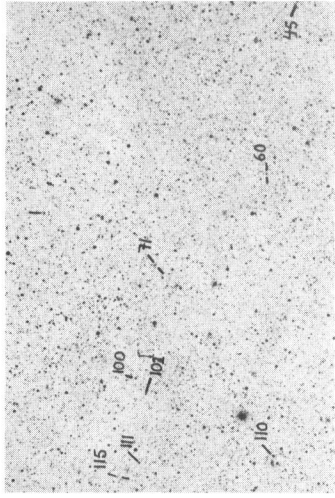


1:8

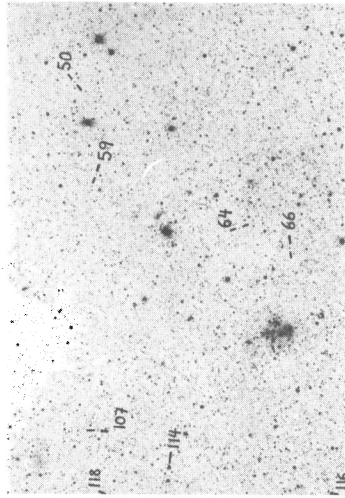


1:9

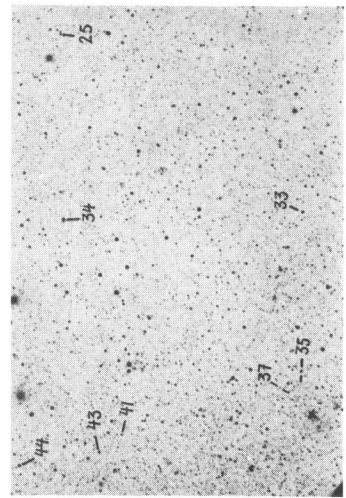
FIGURE 2.- Finding charts (in visual light) for the identified supergiant and giant M stars in the LMC. The former are identified by one line from the number to the star, the latter have an additional line on the other side of the star. - North is up. East is to the left. Each field is 37.8 x 54.7 minutes of arc.



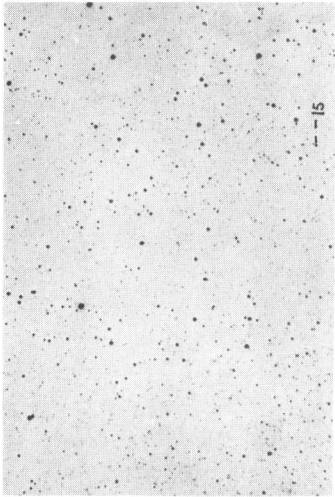
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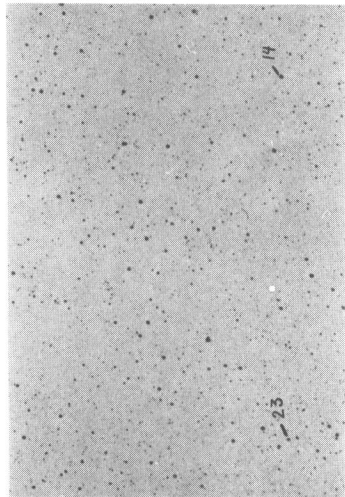
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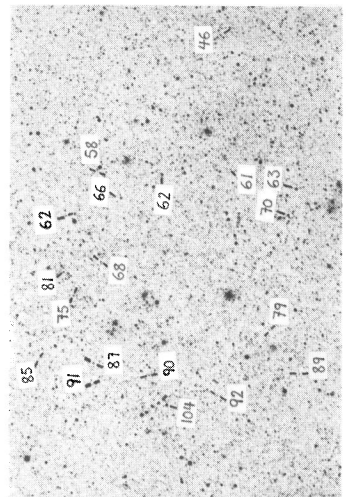
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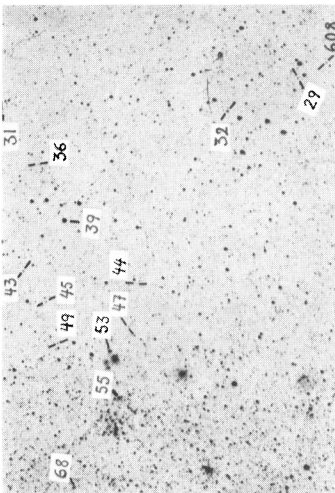
1:11



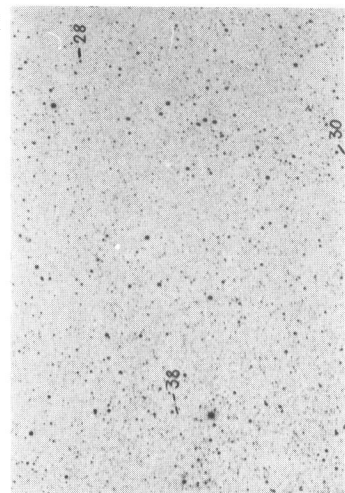
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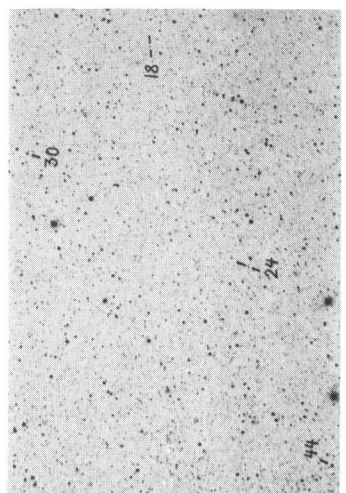
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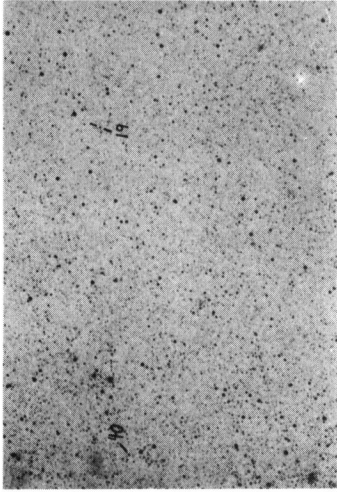
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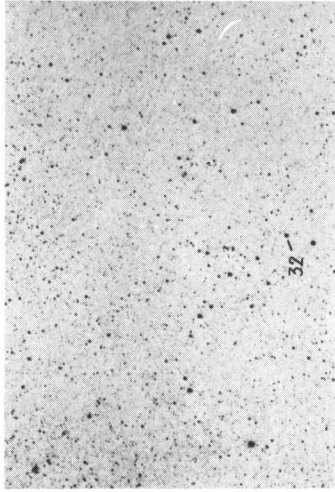
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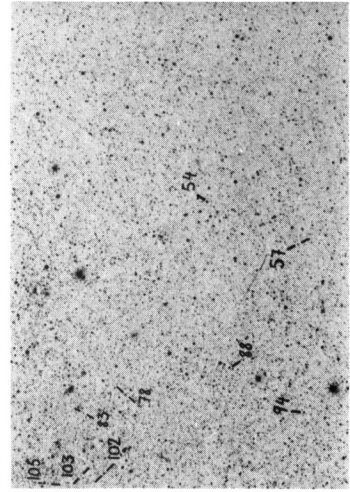
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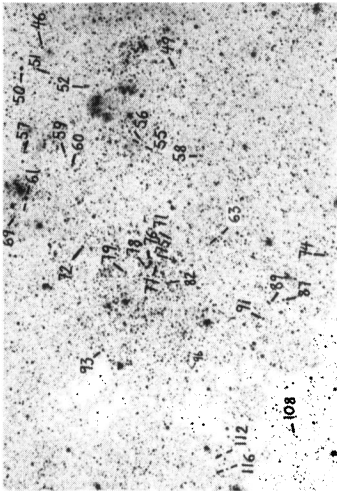
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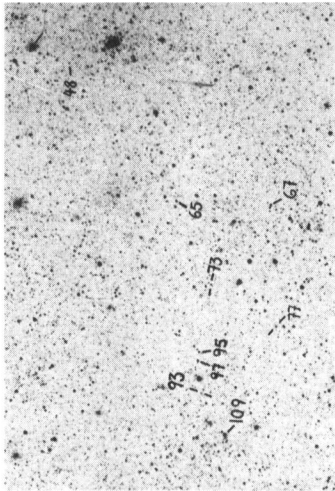
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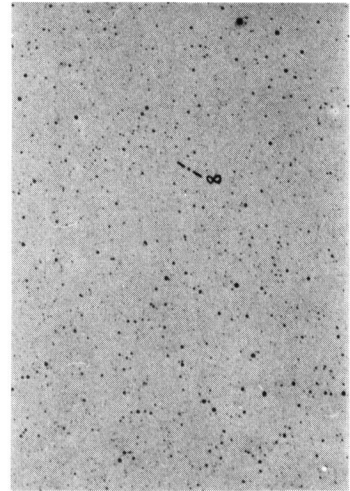
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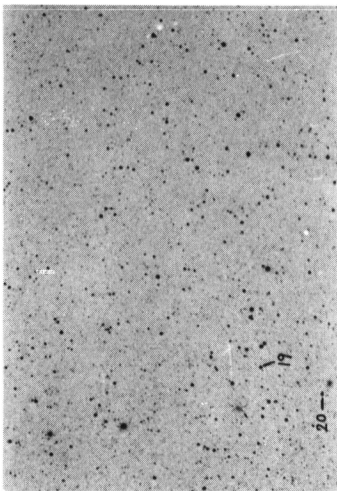
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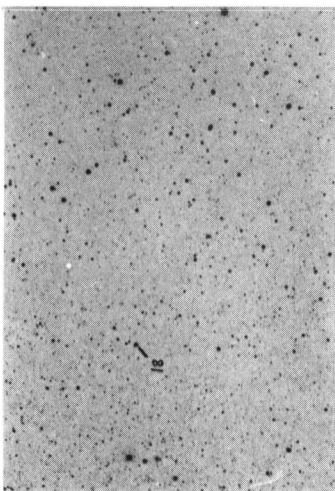
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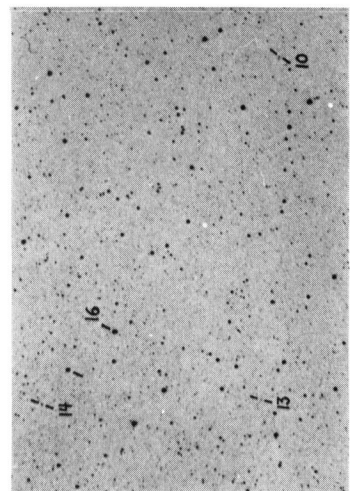
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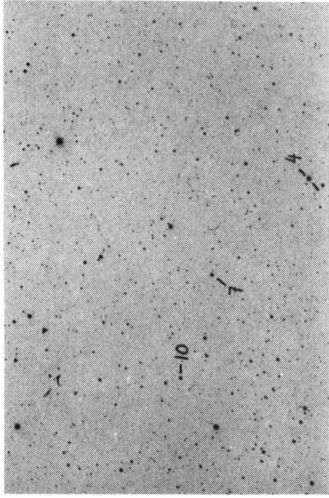
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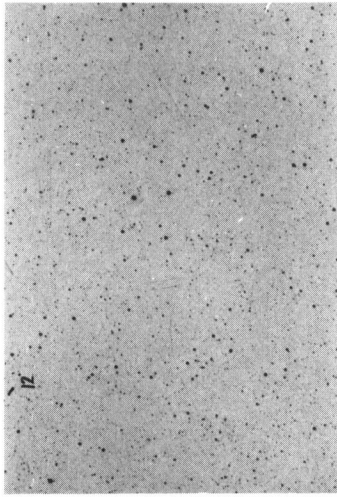
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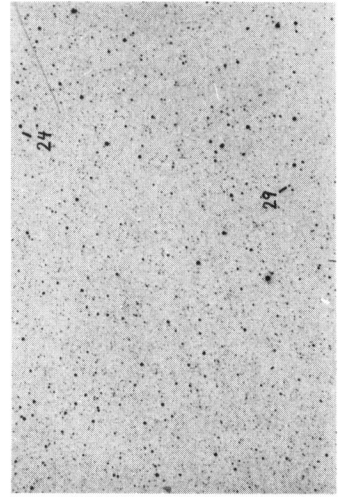
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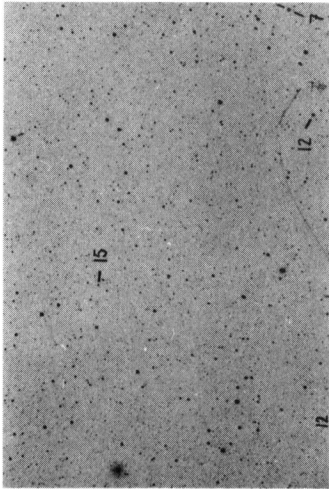
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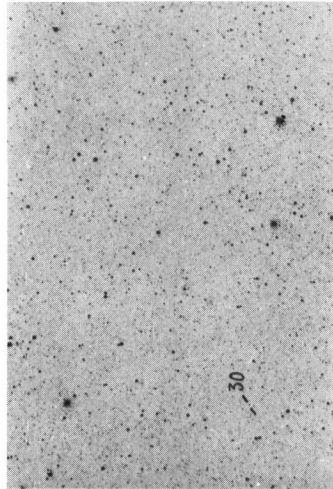
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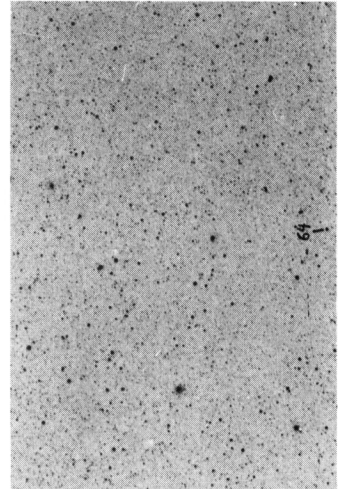
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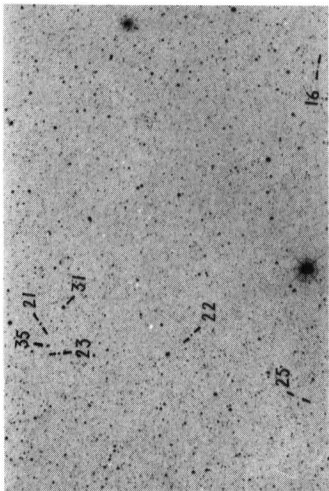
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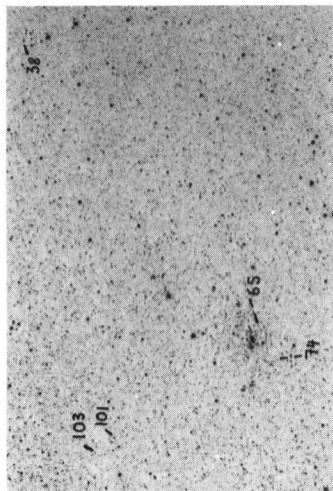
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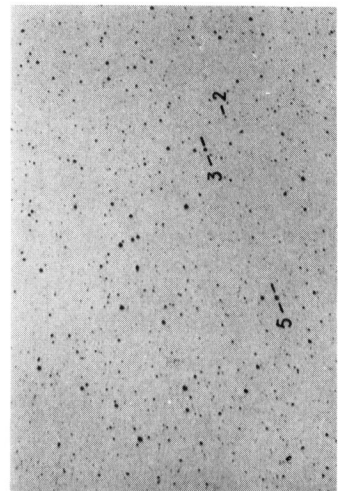
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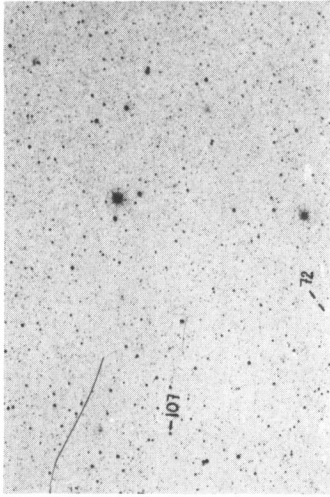
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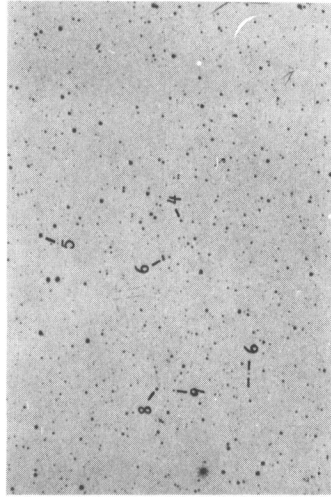
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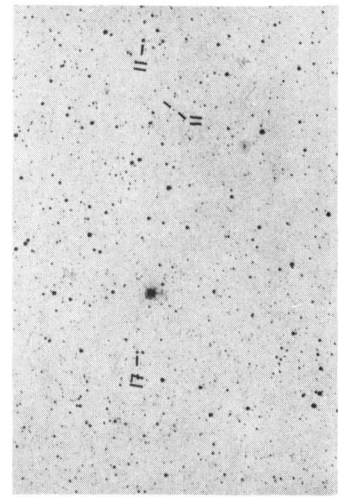
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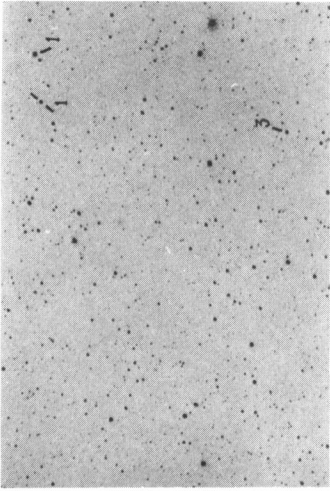
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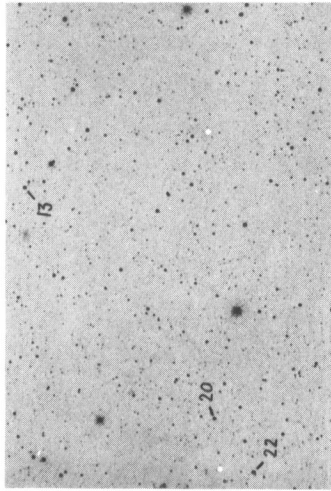
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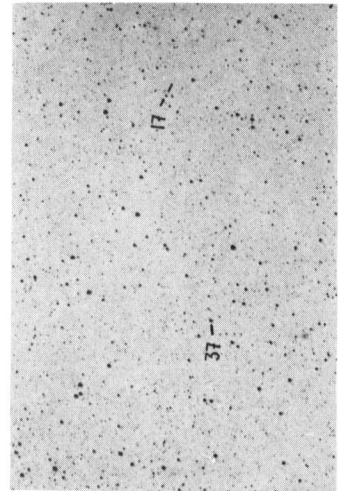
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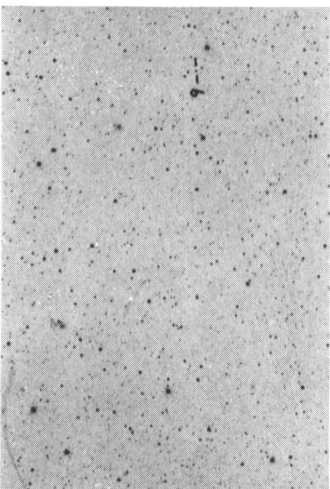
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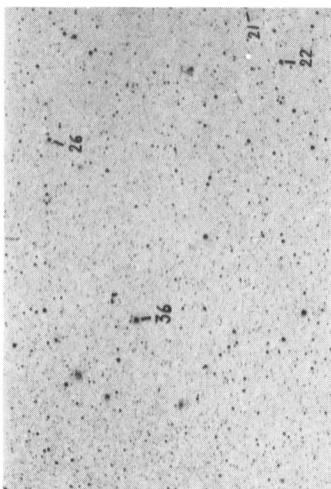
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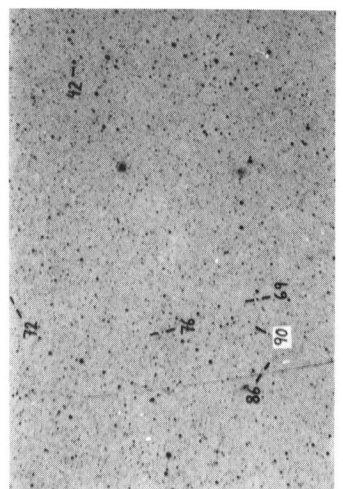
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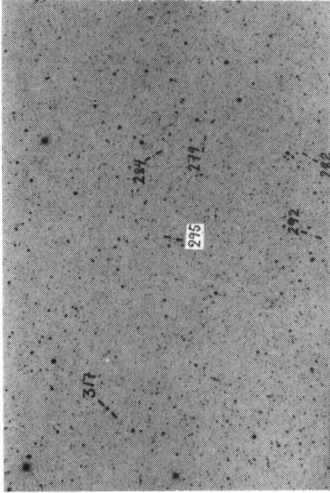
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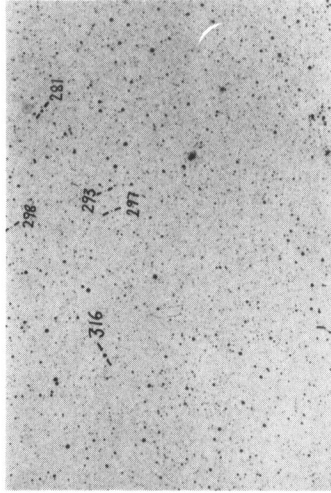
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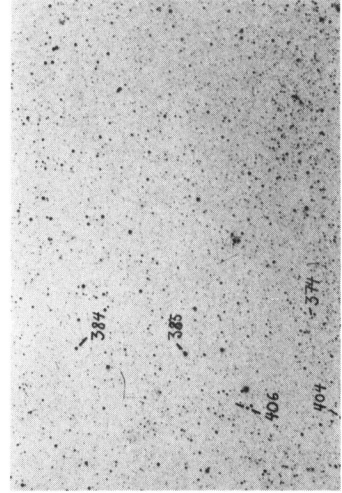
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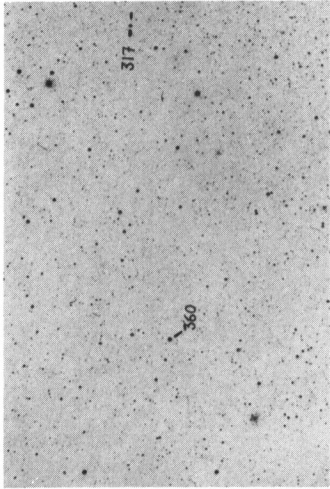
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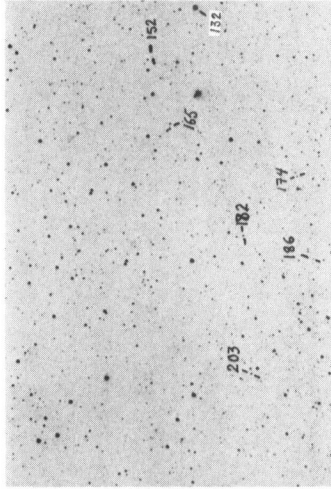
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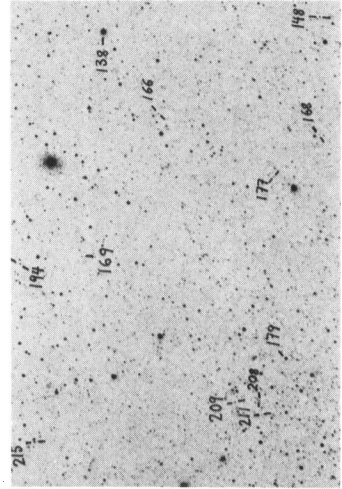
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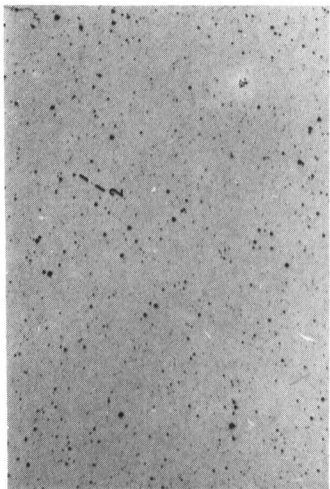
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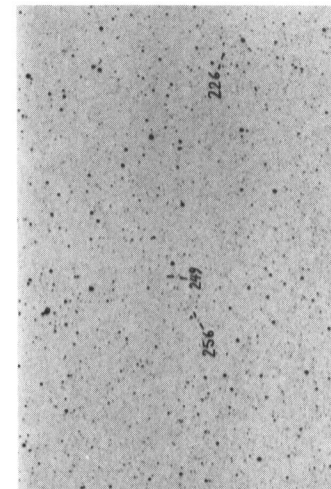
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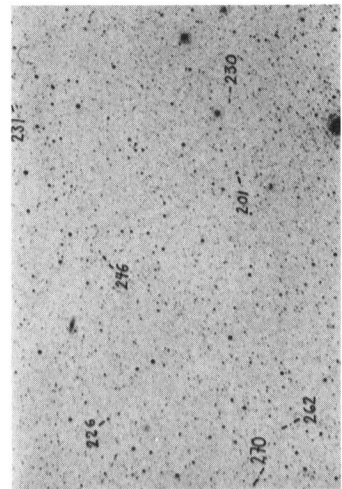
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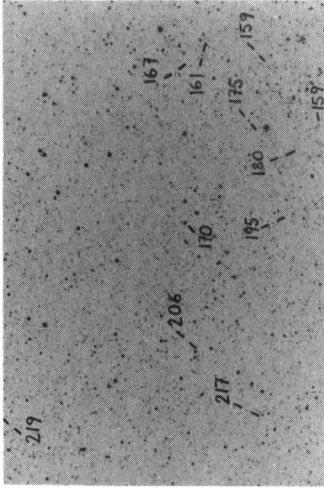
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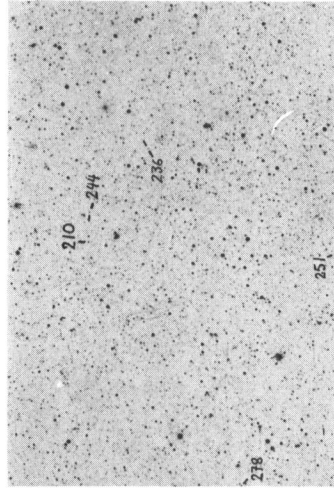
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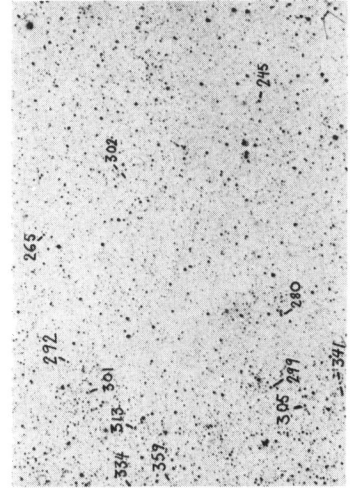
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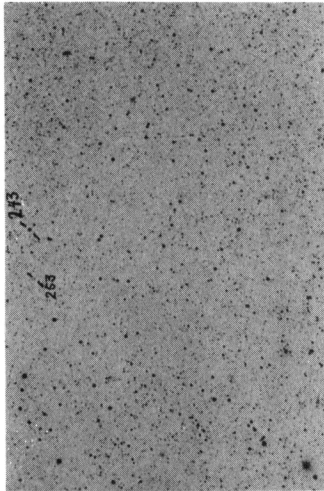
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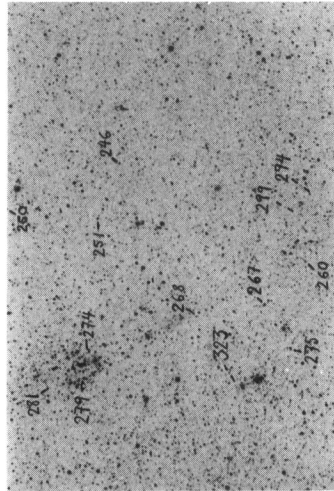
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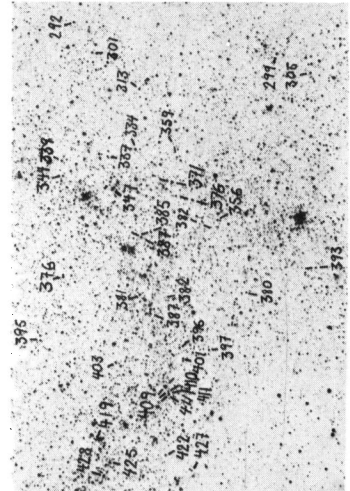
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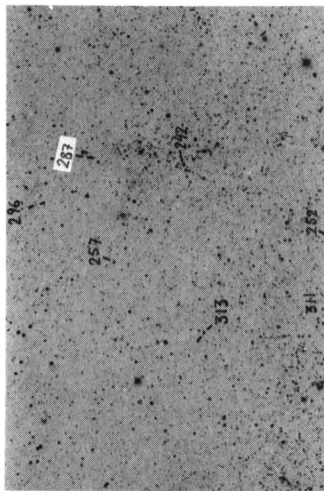
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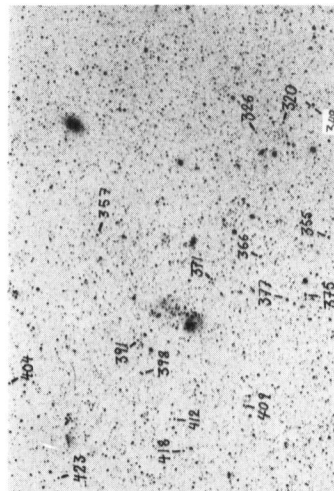
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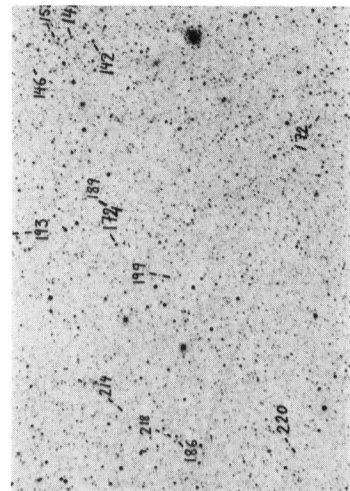
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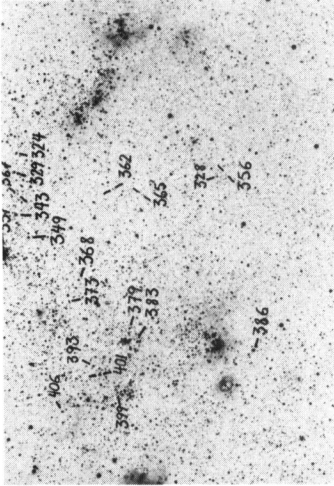
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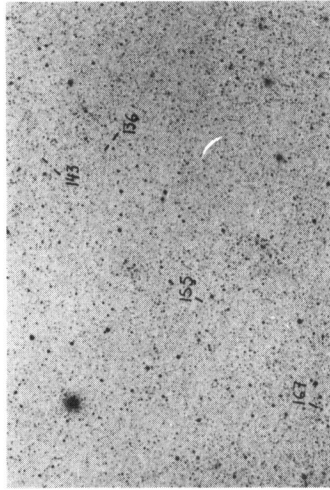
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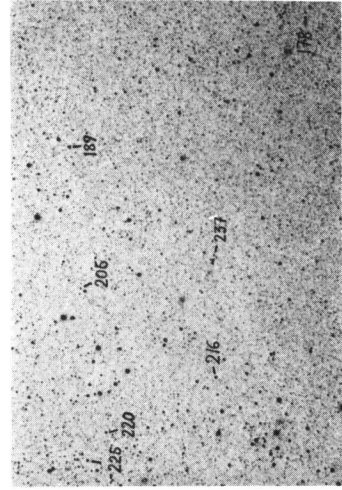
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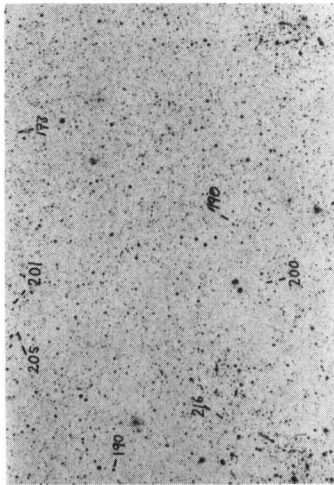
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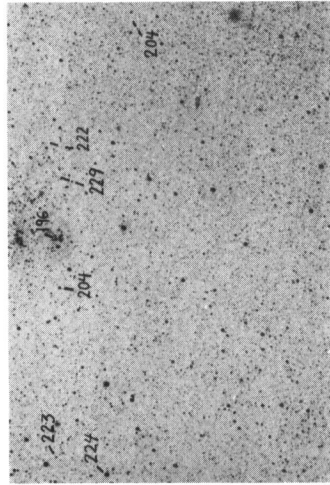
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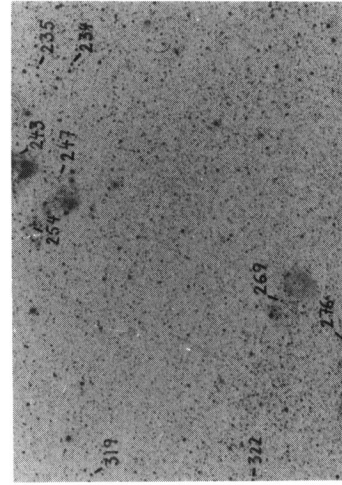
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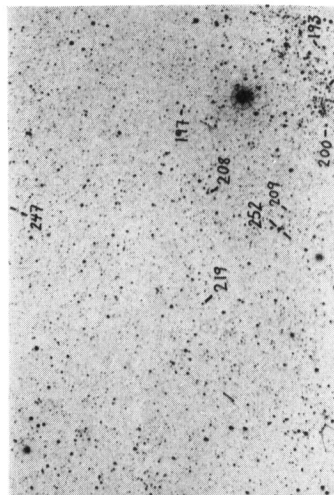
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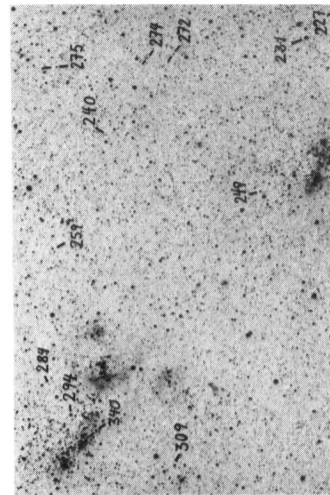
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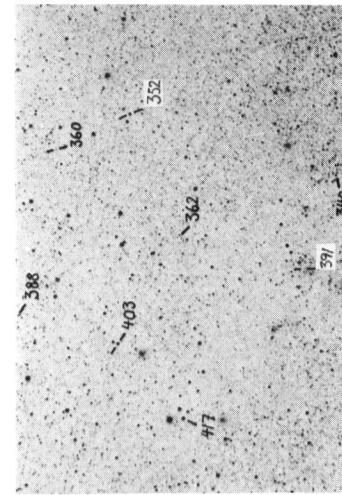
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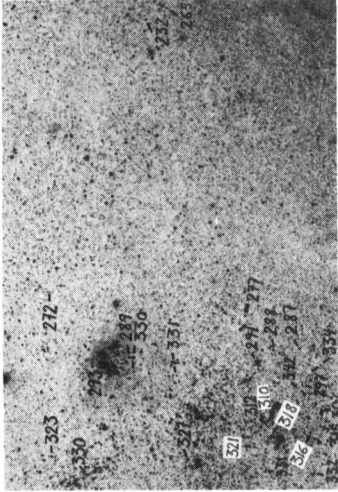
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5:2



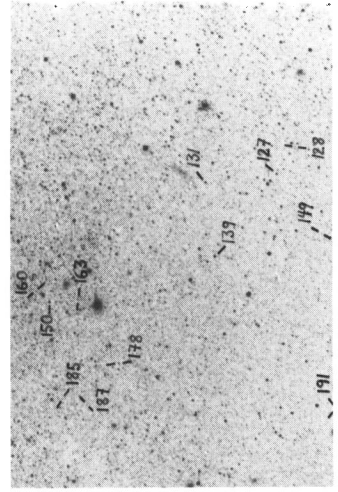
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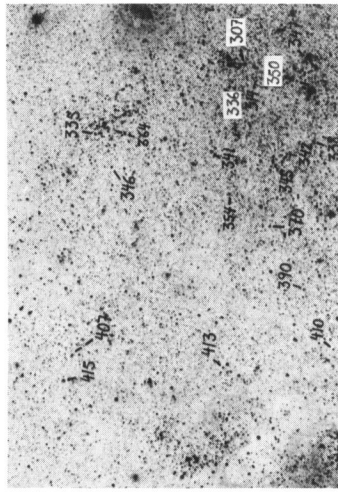
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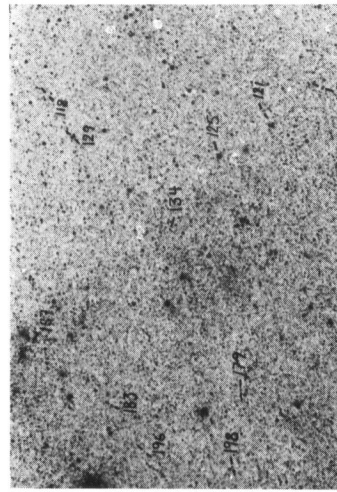
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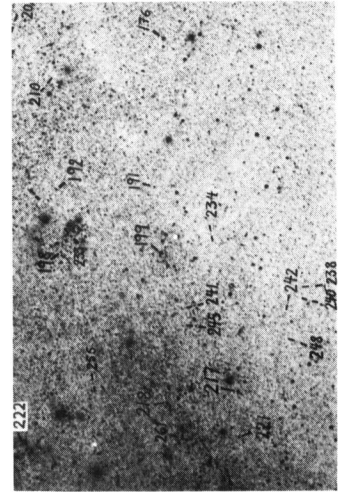
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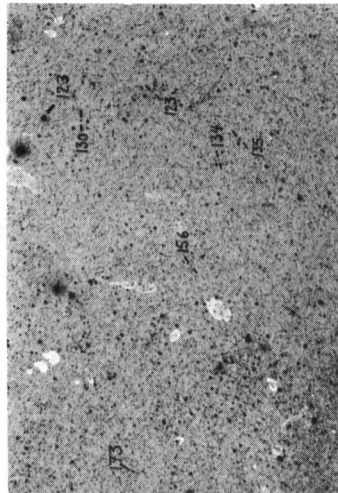
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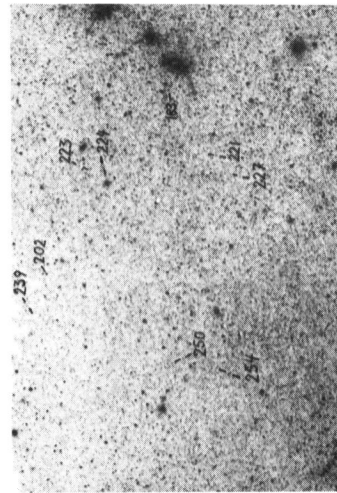
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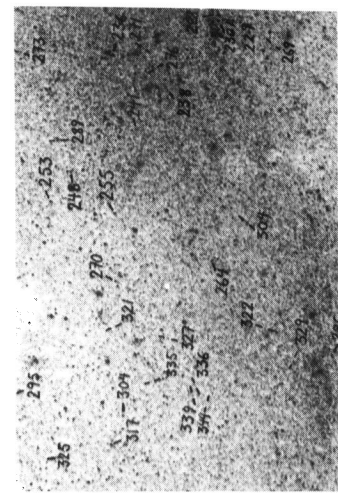
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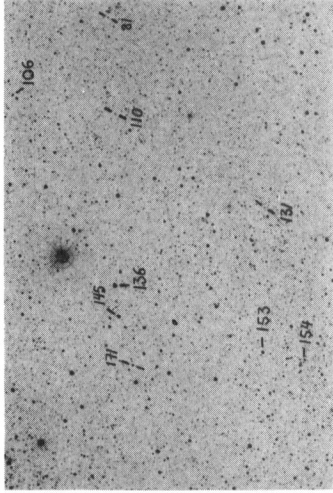
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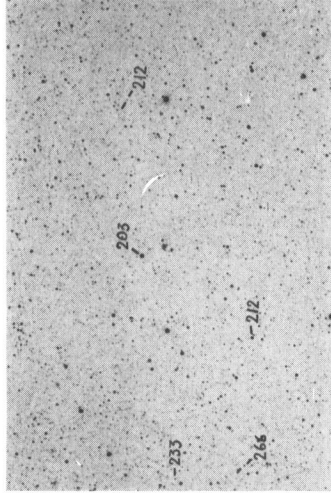
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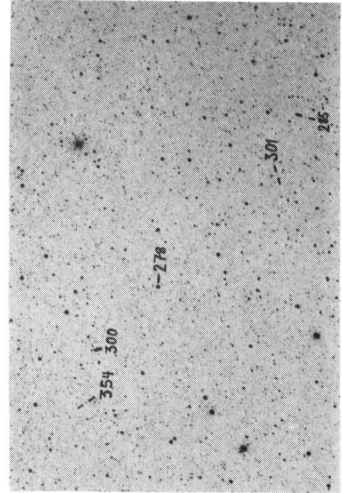
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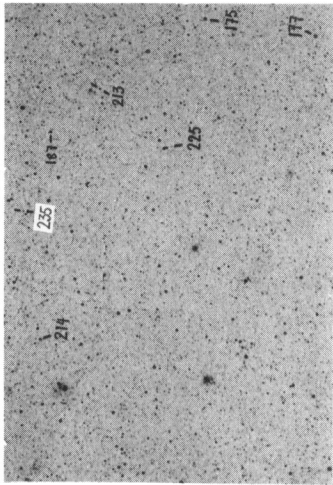
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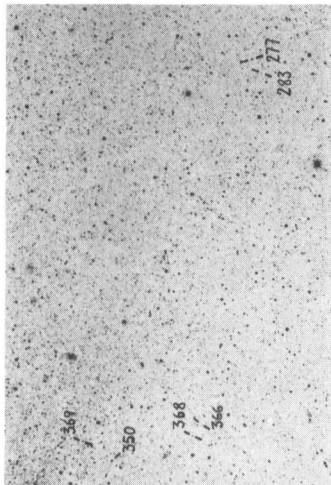
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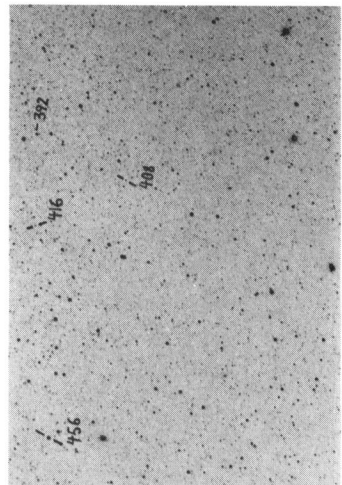
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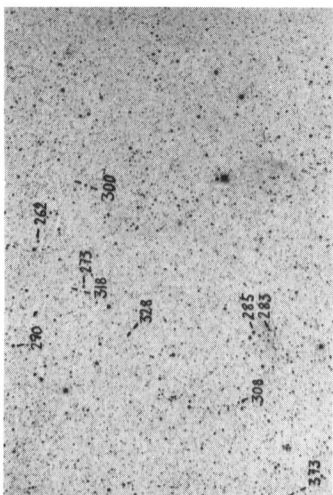
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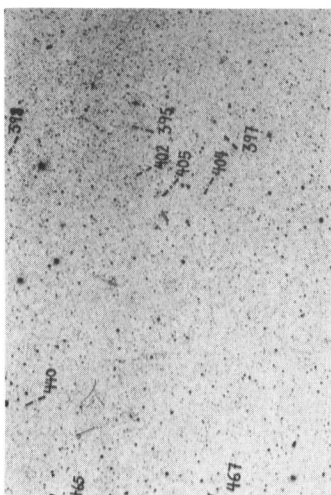
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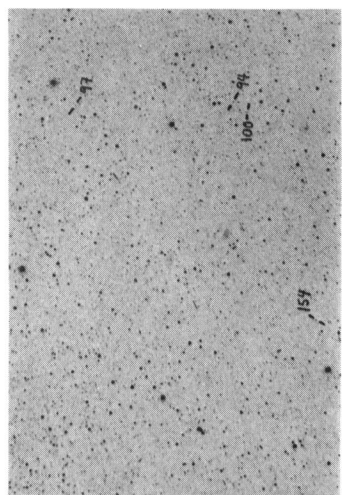
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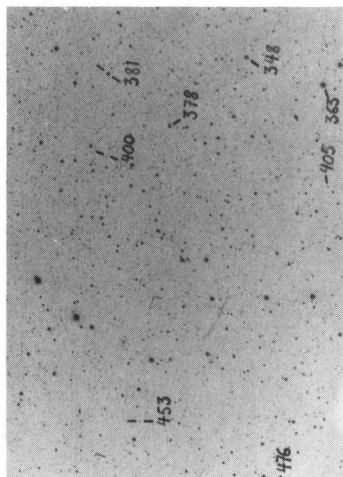
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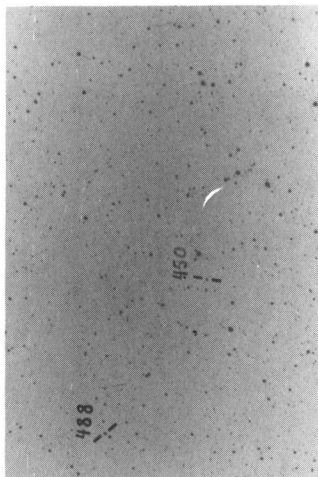
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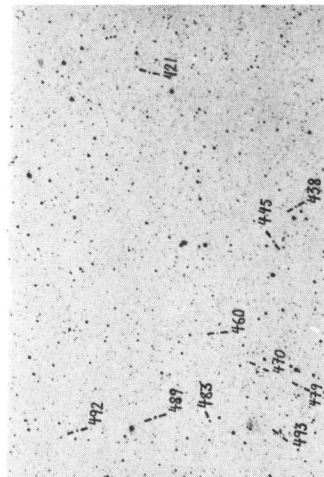
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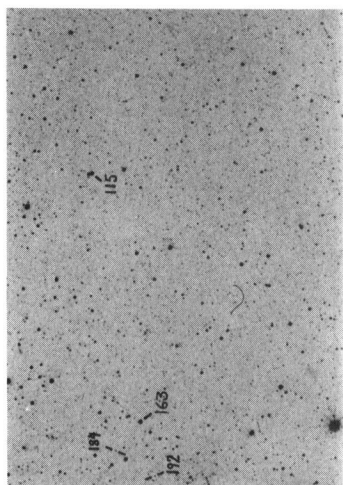
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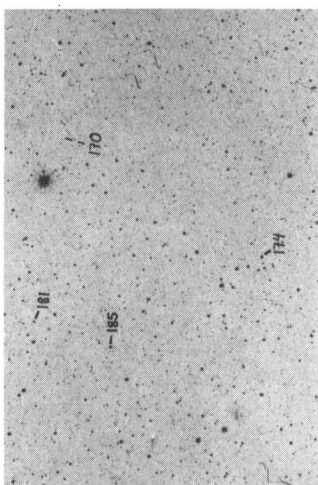
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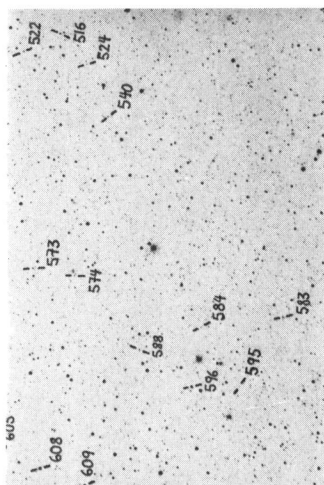
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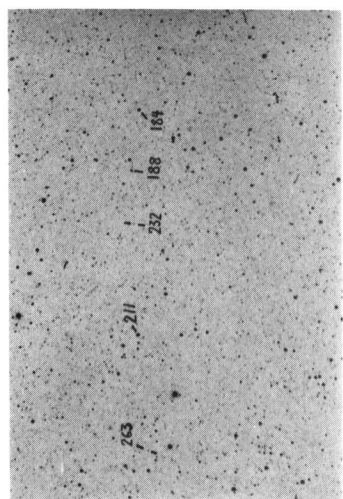
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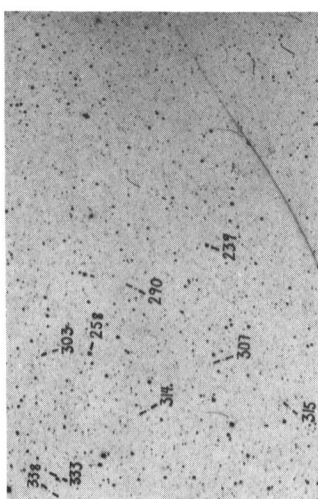
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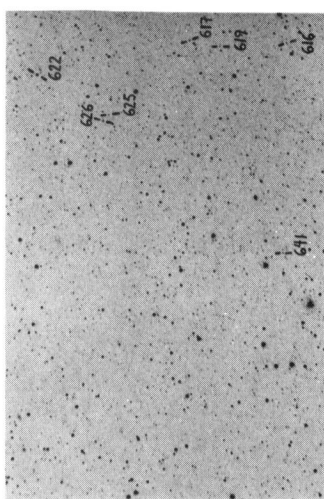
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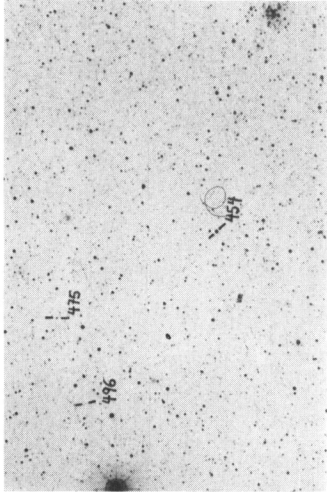
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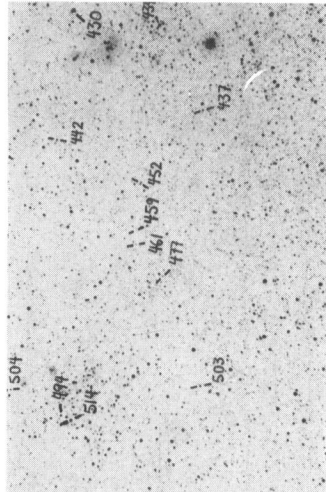
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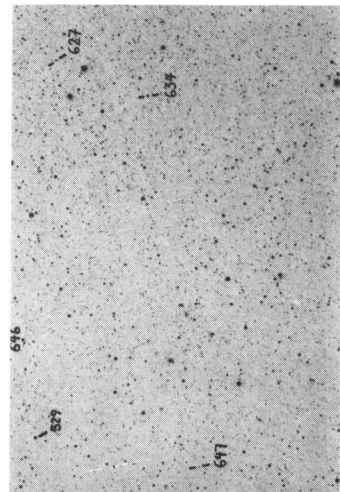
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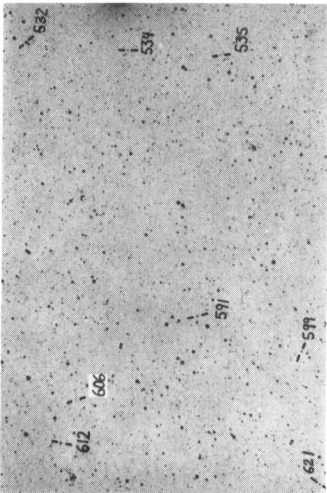
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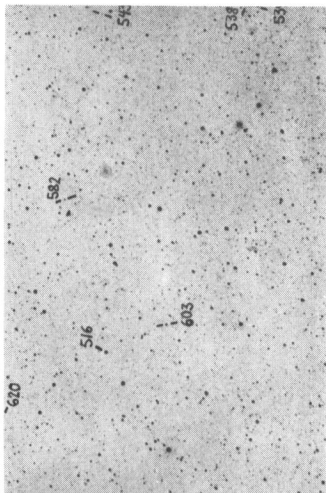
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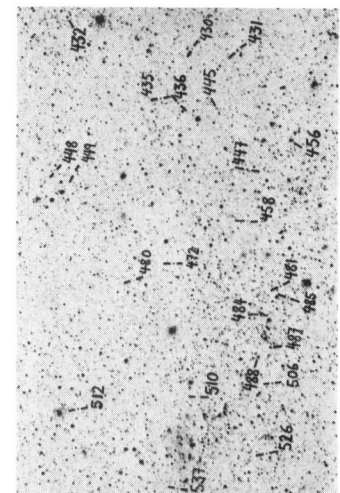
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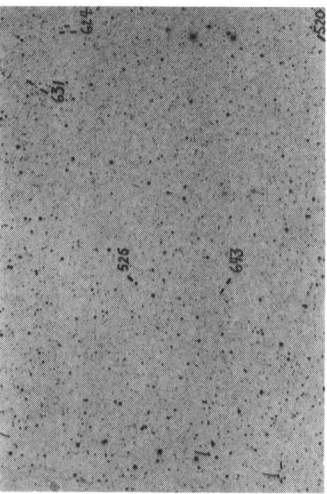
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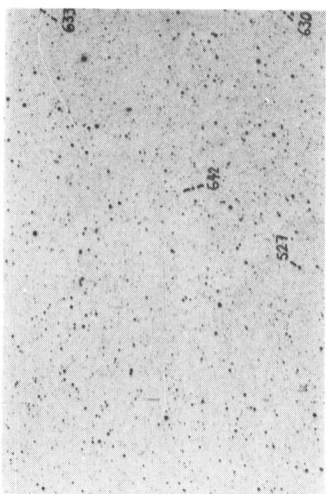
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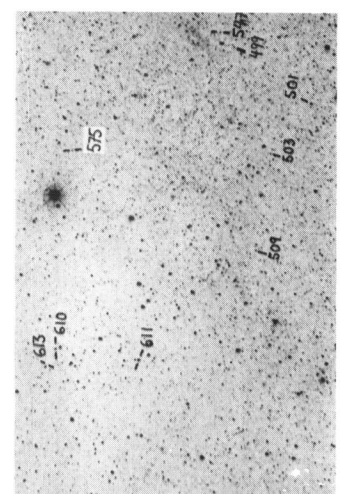
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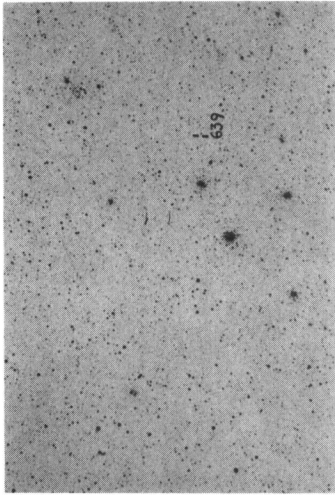
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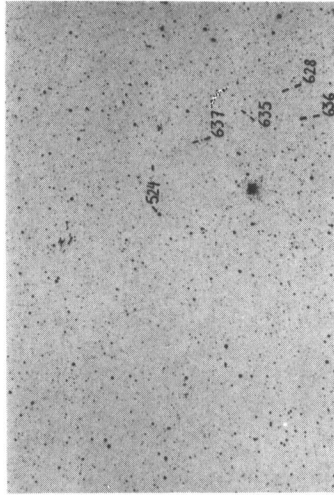
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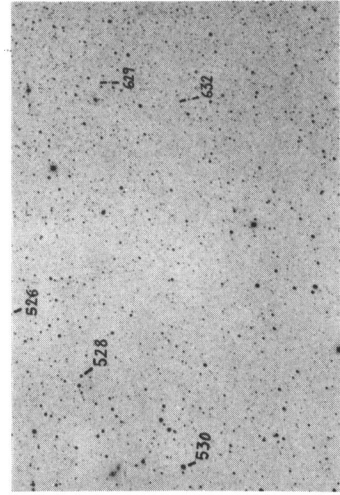
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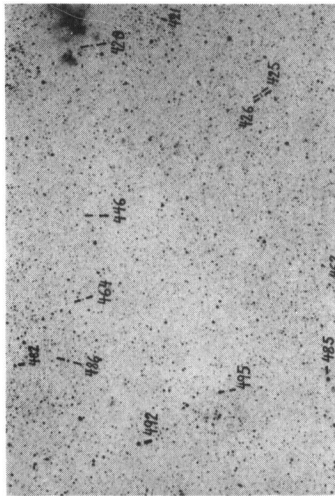
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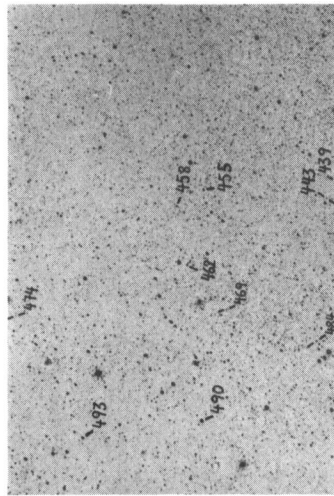
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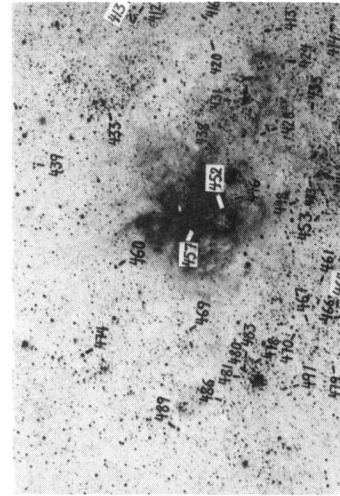
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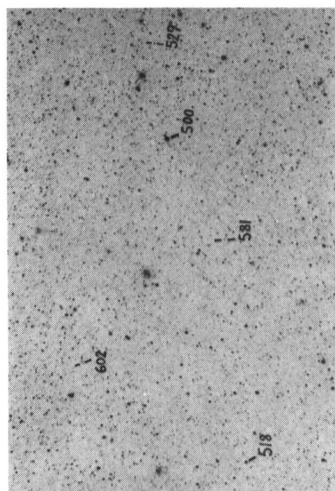
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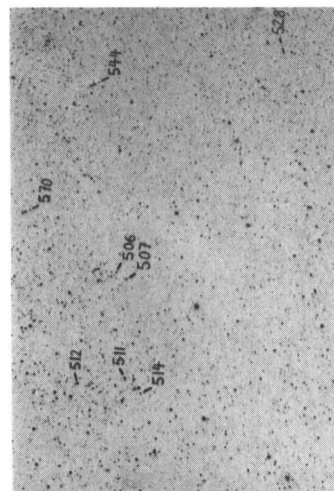
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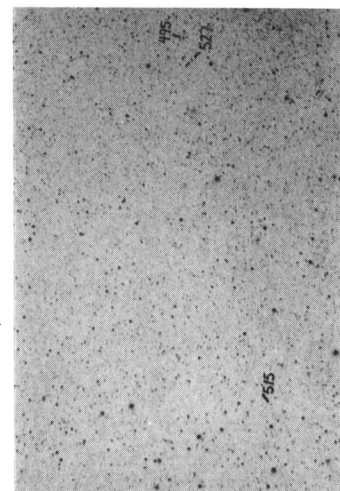
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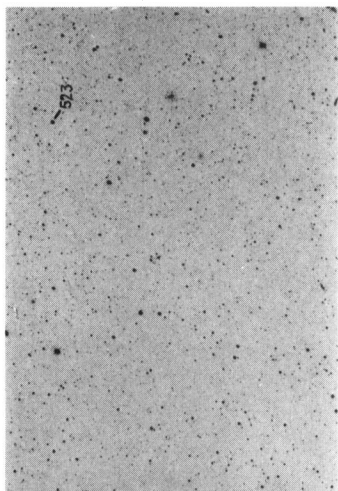
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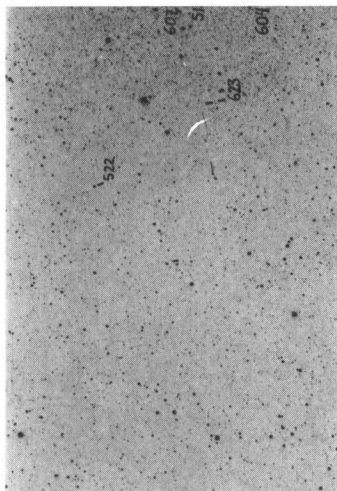
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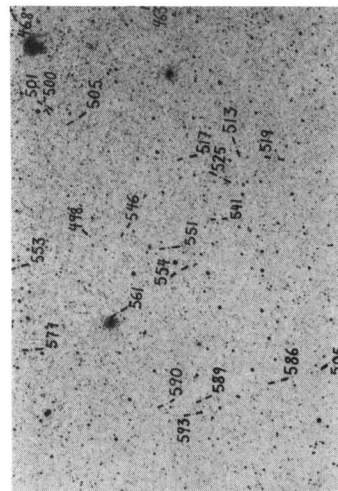
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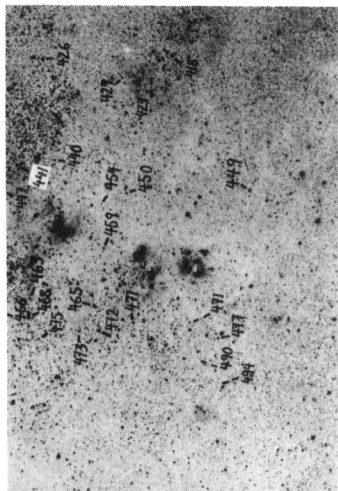
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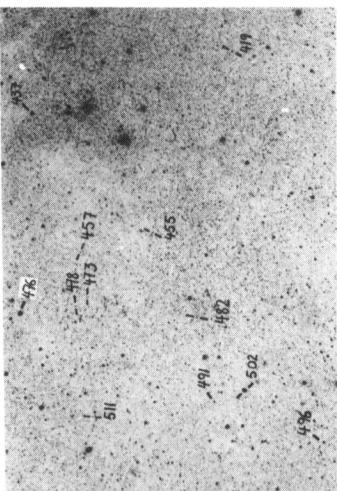
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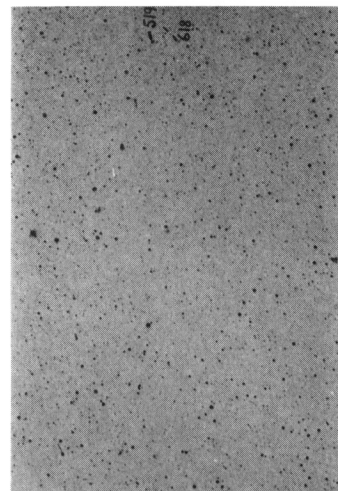
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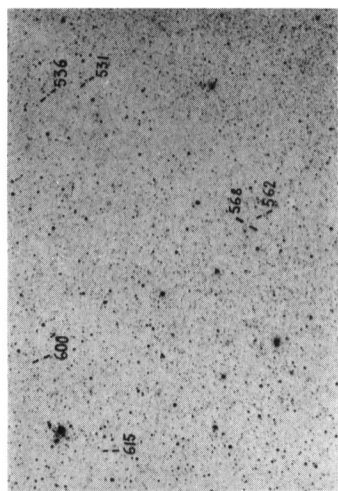
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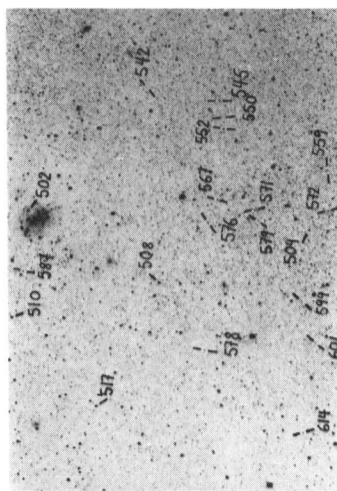
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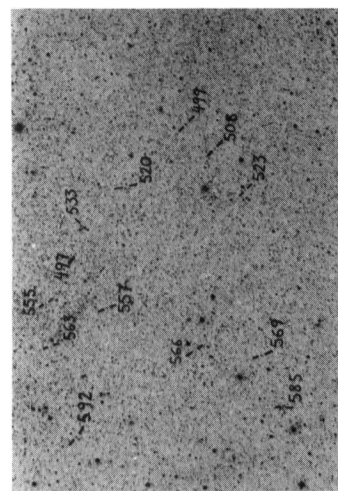
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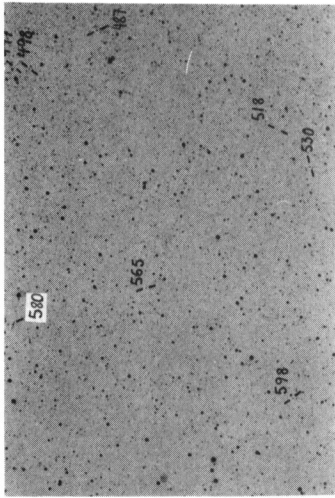
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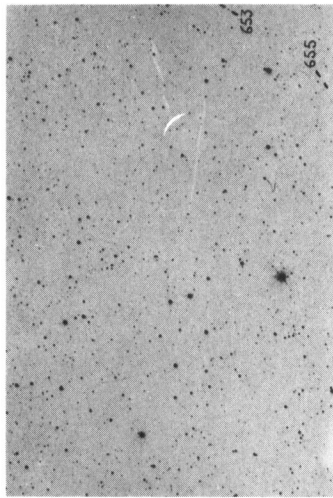
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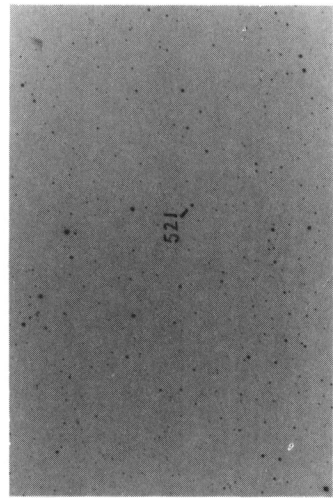
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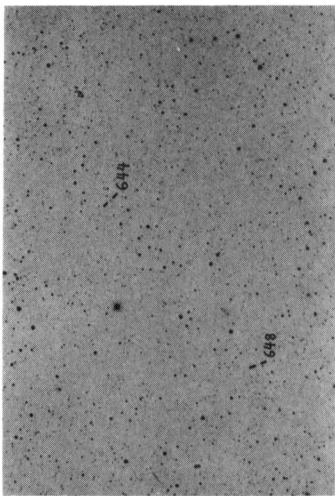
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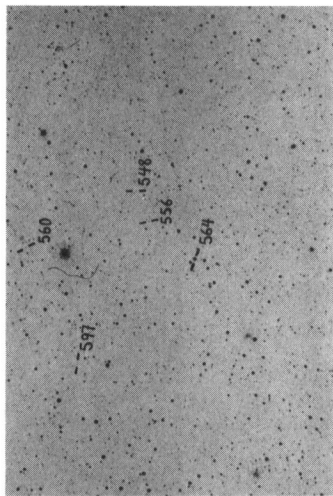
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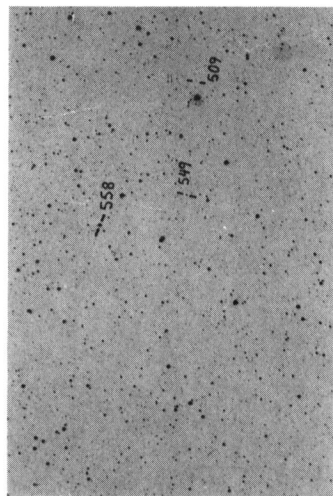
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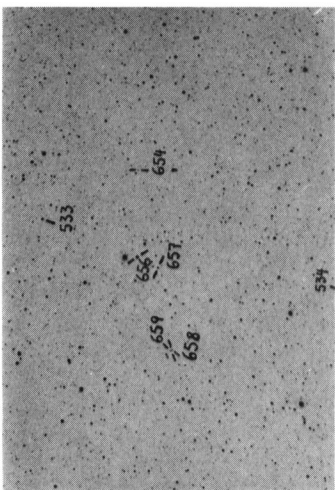
9:6



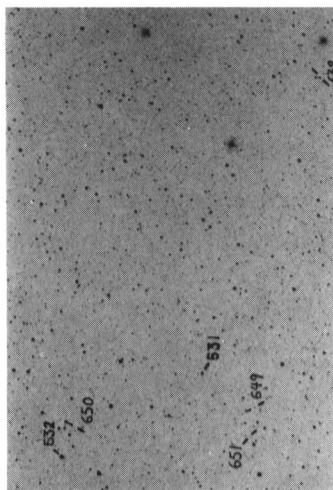
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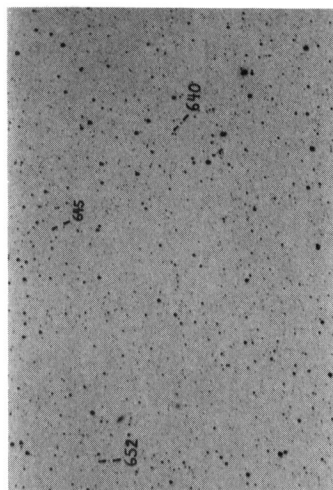
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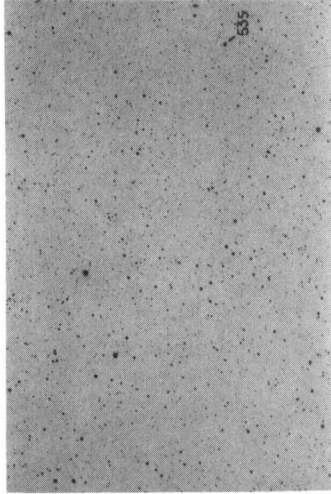
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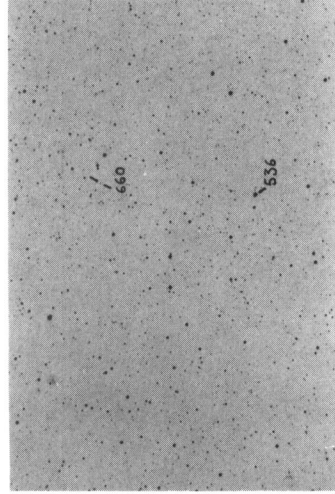
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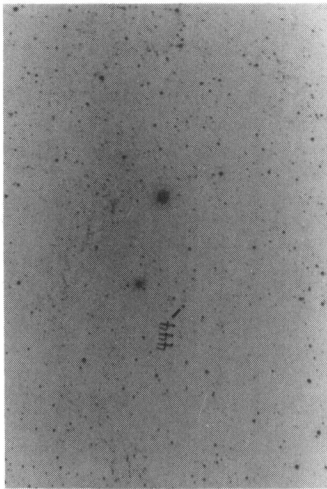
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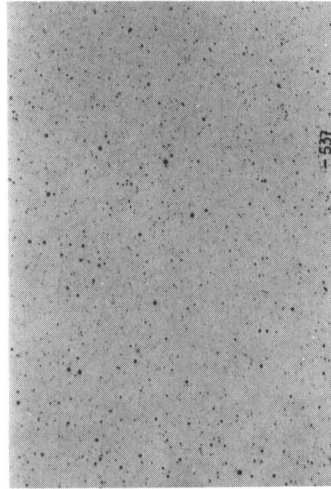
24:1



24:3



20:2



24:2

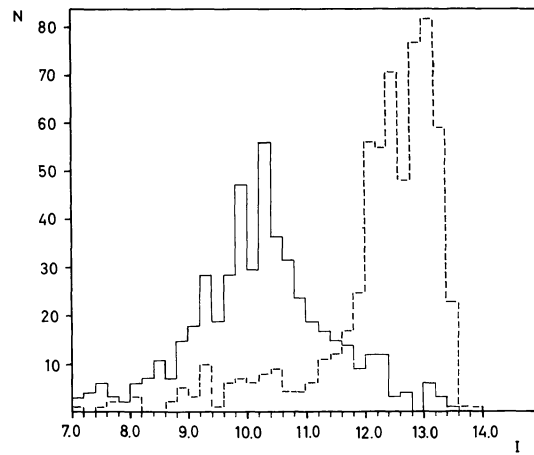


FIGURE 3.- The distributions of the supergiant and giant M stars in magnitude. The full-drawn line gives the histogram for the supergiant M stars and the dashed line that for the giant M stars.