

# Four-colour Photometry of Eclipsing Binaries. II

## RZ Cha, Light Curves, Photometric Elements and Determination of Helium Content

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Received June 24, 1975

**Summary.** Photometric elements are derived for the eclipsing binary RZ Cha. Combined with spectroscopic elements the absolute dimensions are derived. A comparison with new stellar model calculations was carried out and the helium content of RZ Cha was found to be  $Y=0.21 \pm 0.05$ .

**Key words:** eclipsing binary – light curves – photometric elements and helium content

### 1. Introduction

RZ Cha (HD 93486 = CoD – 81° 391 = BV 473,  $m_V = 8.1-8.5$ , Sp F5 IV – V) was mentioned as double-lined spectroscopic binary by Popper (1966) and recommended by Batten (1971) to be observed photometrically. Spectroscopic observations have been obtained by Andersen *et al.* (1975) who derive  $M_1 \sin^3 i$  and  $M_2 \sin^3 i$  to an accuracy of 1.7%. The masses are equally large to this accuracy. The components seem to be somewhat evolved according to the spectral classes, and the masses also seem a bit high for the spectral type. The period is 2<sup>d</sup>.8 and the light curve seems uncomplicated, see Strohmeier *et al.* (1964). One should therefore expect that fundamental parameters for this binary could be derived with good accuracy and that a comparison with stellar evolution calculation could be made and helium content derived.

### 2. The Observations

The observations were carried out at the European Southern Observatory, La Silla, Chile on 23 nights during the period March 1972–June 1973. The observations on the last two nights were made by Grønbech. All observations were done using the 50 cm reflecting telescope belonging to the Copenhagen Observatory. The Strömgren four channel photometer was used to give observations simultaneously in the four colour *uvby* system. Two comparison stars were selected in order to check these stars for variability. Altogether 775 observations of RZ Cha were obtained in each of the four colours. The scatter in the magnitude difference of the comparison stars is found to be  $\sigma(\Delta u) = 0.007$ ,  $\sigma(\Delta v) = 0.005$ ,  $\sigma(\Delta b) = 0.005$  and  $\sigma(\Delta y) = 0.007$  which is very nearly the same as for the comparison stars of

HS Hya discussed by Gyldenkerne *et al.* (1975). The magnitude difference between the two comparison stars was plotted night by night and no variation was found. The light curves of RZ Cha in the instrumental system are given in Table 1 and Fig. 1, where HR 4312 has been used as comparison star. For further details of methods we refer to Gyldenkerne *et al.* (1975).

The following times of minimum (Table 2) are derived using the method described by Kwee and van Woerden (1956). The period is not well determined from the times of minimum and instead we have developed a computer program which derives the scatter around the mean light curve for an assumed period. A scan through periods were made and the period giving the smallest scatter around the mean curve was accepted. This method includes branches of minima with high weight even if only one branch is observed on a particular night. We find

$$\text{Pr. Min} = 2441401.7711 + E \times 2^d.832084 \quad (1)$$

from which the  $O-C$  values are computed for the primary minima. The  $O-C$  values for the secondary minima are determined assuming phase 0.5 exactly. Obviously the secondary minima fall very precisely at phase 0.5 and we assume that the system has circular orbits, in agreement with the spectroscopic evidence.

### 3. Photometric Element Determination

First we used the rectification procedure described more in detail for HS Hya by Gyldenkerne *et al.* (1975). The rectification constants are given in Table 3.

From the further discussion and the results by Andersen (1975) we derive the mean dynamical ellipticity  $z = 0.019$

TABLE 1. RZ CHA - HR 4312 IN THE INSTRUMENTAL SYSTEM.

Table with 12 columns: IJD-2440000, AU, AV, AB, AY, IJD-2440000, AU, AV, AB, AY, IJD-2440000, AU, AV, AB, AY, IJD-2440000, AU, AV, AB, AY, IJD-2440000, AU, AV, AB, AY. The table contains a dense grid of numerical data points for various astronomical observations.

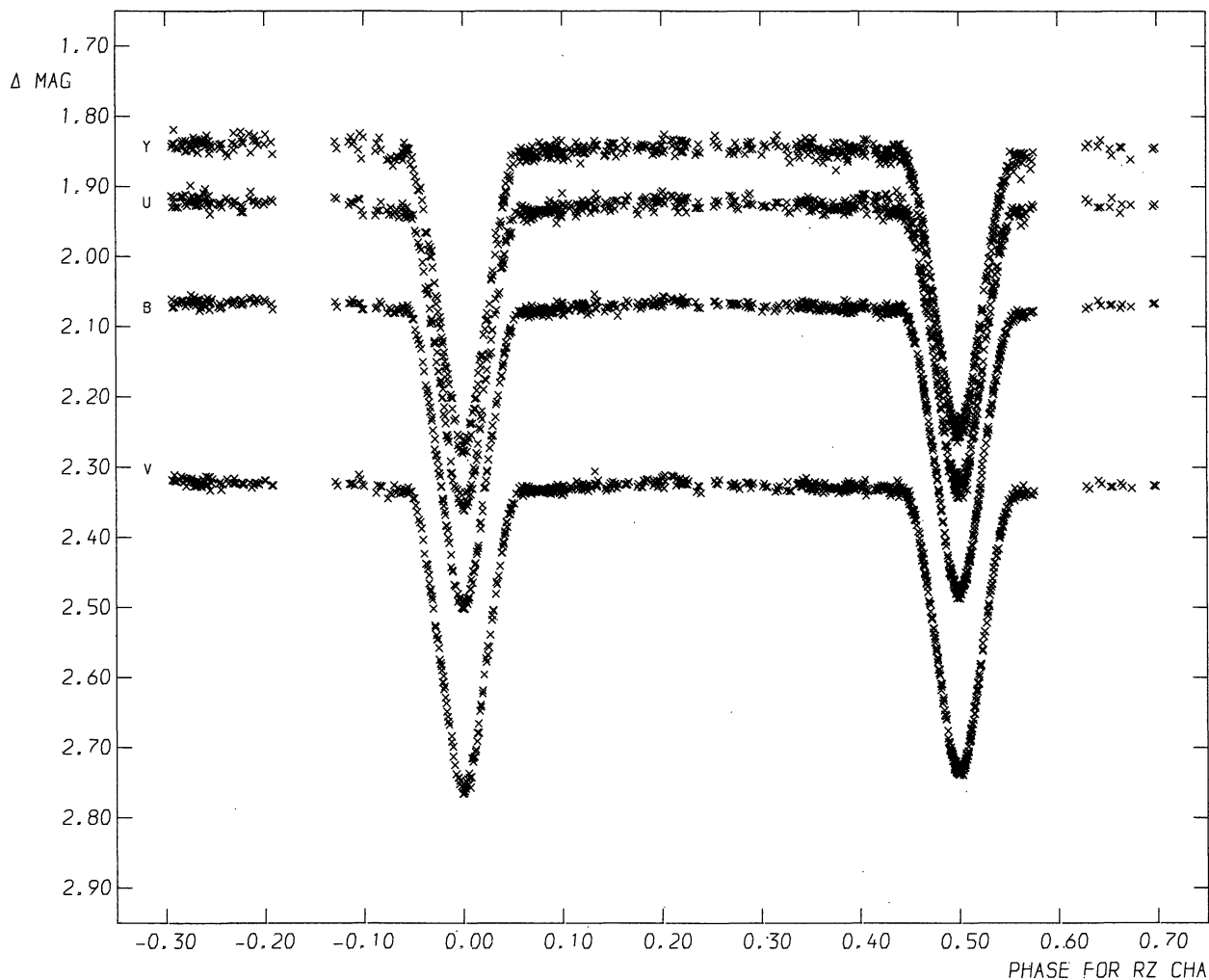


Fig. 1. *uvby* light curves in the instrumental system for RZ Cha. HR 4312 is used as comparison star

Table 2. Times of minimum

	O-C	E	
Pr.min			
2441401.7709	-.0002	0	
1404.6033	+.0001	+ 1	
Sec. min			
2441394.6911	+.0002	- 2.5	
1411.6834	.0000	+ 3.5	
1836.4935	-.0025	+153.5	insufficiently observed

Table 3. Fourier coefficients with mean errors and rectification constants.  $\sigma_r$  is the r.m.s. scatter between the minima

	<i>u</i>	<i>v</i>	<i>b</i>	<i>y</i>
$A_0$	.9738 ± .0004	.9742 ± .0002	.9791 ± .0002	.9744 ± .0004
$A_1$	-.0016 ± .0006	.0003 ± .0004	.0004 ± .0003	.0010 ± .0006
$A_2$	-.0066 ± .0006	-.0061 ± .0004	-.0064 ± .0004	-.0067 ± .0006
$C_0$	.0137	.0137	.0137	.0137
$C_1$	.0016	-.0003	-.0004	-.0010
$C_2$	.0046	.0046	.0046	.0046
<i>z</i>	.0144	.0138	.0140	.0179
$\sigma_r$	.007	.004	.004	.007

of the two components, in fair agreement with the small geometrical ellipticity found from the light curves. We notice that the constants are small and that the residual scatter between the minima are the same as for HS Hya, which is of similar apparent magnitude. In order to derive the elements we tried to use the method of Kopal (1959), that is an iteration in *k* where

*k* is the ratio of the radii of the components. However, we ran into precisely the same troubles as for HS Hya and no value of *k* could be determined in this way. For some colours the iteration did converge but if the intensities in the minima ( $\lambda_{pr}$ ,  $\lambda_{sec}$ ) were changed by only 0.001 a completely different value of *k* was obtained. Other colours gave no solution in the iterative

process since the difference  $\Delta k$  between any starting value of  $k$  and an improved value was either positive or negative.

Instead we have used the method described for HS Hya by Gyldenkerne *et al.* (1975). For adopted values of  $(k, \lambda_{pr}, \lambda_{sec})$  we compute  $p_0 = (\cos i - r_b)/r_a$  and find the fractional light losses  $\alpha_0$  in the two minima. The constants  $(C_1, C_2)$  in the formula

$$\sin^2 \Theta = (p^2 - p_0^2)C_1 + 2(p - p_0)C_2 \quad (2)$$

are found by the usual method and a least squares fit. The radii  $r_a$  and  $r_b$  are then found from

$$r_a = \left( p_0^2 + \frac{1}{k^2} + \frac{2p_0}{k} + \frac{1}{k \cdot C_2} \right)^{-1/2}, \quad (3)$$

$$r_b = r_a/k, \quad (4)$$

and finally the inclination from

$$\cos i = r_a \cdot p_0 + r_b. \quad (5)$$

By a "best fit" method we have thus found a set of elements adopting  $(k, \lambda_{pr}, \lambda_{sec})$ . Finally the light curve is computed for these elements and the residual scatter is found. By varying  $k$  we should in principle be able to find the minimum of the scatter and therefore a final set of elements. However, this procedure does not work in practice since there is no well defined minimum of the scatter and because the  $k$  giving the minimum is extremely sensitive to the adopted  $(\lambda_{pr}, \lambda_{sec})$ . Also it was impossible to decide if the primary was a transit or an occultation. There was no need for a final adjustment of  $(\lambda_{pr}, \lambda_{sec})$  since the residual scatter around the computed light curves was sufficiently small. In Table 4 we give the intensities of the minima and the adopted limb-darkening coefficients from Grygar *et al.* (1972). Microphotometer tracings of spectra of RZ Cha by Andersen *et al.* (1975) show that the spectral types and

Table 4. Limb-darkening coefficients and the minimum intensities in the rectified light curves

	Limb-darkening $u$	Minimum intensities	
		$\lambda_{pr}$	$\lambda_{sec}$
$u$	0.73	.687	.700
$v$	0.82	.683	.698
$b$	0.73	.686	.699
$y$	0.61	.685	.701

luminosities of the two components are very similar. This is confirmed by our photometry since there is no detectable variation in the colour index  $b-y$  and the  $c_1$  index. If there is any variation it is less than approximately 0.002 in these indices.

In Table 5 we give solutions for  $k=0.90$  and  $k=1.00$  and this interval seems to give luminosity ratios acceptable from the spectroscopic point of view. From these solutions we find that the flux in all four colours is approximately 5% higher for the hotter star which corresponds to a temperature difference of  $\Delta \log T_e \approx 0.005$  or a spectral type difference of only 0.6 subclasses in agreement with no detectable difference. Comparing the  $b$  and  $y$  solutions we find that the difference in colour index between the two components is  $\Delta(b-y) = -0.010$  in the sense that the bluest star is the cooler one. This is of course not acceptable but we need only a change  $(\lambda_{pr}, \lambda_{sec})$  by  $\sim 0.001$  in order to make the cooler star the reddest one. The disagreement between colour differences and flux ratios is therefore no problem.

#### 4. Discussion

We have seen that the ratio of the luminosities and the radii are very close to unity and we shall therefore dis-

Table 5. Solutions for  $k=0.90$  and  $k=1.00$  assuming partial occultation and partial transit.  $\Delta \log F = \log F_b - \log F_a$ ;  $\sigma_L$  = r.m.s. scatter

Pr. min.	$k$	Colour	$r_a$	$r_b$	$i_r$	$\sigma_L$	$L_a$	$L_b$	$L_b/L_a$	$\Delta \log F$
Occ. par.	0.90	$u$	.1763	.1959	82.84	.0064	.4590	.5410	1.1786	-.0201
		$v$	.1766	.1962	82.83	.0040	.4606	.5394	1.1711	-.0229
		$b$	.1765	.1961	82.87	.0037	.4590	.5410	1.1786	-.0201
		$y$	.1755	.1949	83.01	.0063	.4613	.5387	1.1678	-.0241
	1.00	$u$	.1864	.1864	82.78	.0064	.5106	.4894	.9585	-.0184
		$v$	.1864	.1864	82.78	.0042	.5121	.4879	.9527	-.0210
		$b$	.1861	.1861	82.82	.0037	.5106	.4894	.9585	-.0184
		$y$	.1849	.1849	82.97	.0063	.5130	.4870	.9493	-.0226
Tr. par.	0.90	$u$	.1759	.1955	82.90	.0064	.4380	.5620	1.2831	+0.167
		$v$	.1761	.1957	82.90	.0047	.4367	.5633	1.2899	+0.190
		$b$	.1761	.1956	82.93	.0041	.4381	.5619	1.2826	+0.166
		$y$	.1751	.1946	83.07	.0064	.4356	.5644	1.2957	+0.210
	1.00	$u$	.1864	.1864	82.78	.0064	.4894	.5106	1.0433	+0.184
		$v$	.1864	.1864	82.78	.0042	.4879	.5121	1.0496	+0.210
		$b$	.1861	.1861	82.82	.0037	.4894	.5106	1.0433	+0.184
		$y$	.1849	.1849	82.97	.0063	.4870	.5130	1.0534	+0.226

cuss the mean component. We have

$$R = (r_a + r_b)/2 = (0.1860 \pm 0.0007) \cdot a = 2.26 \pm .02 R_\odot$$

$$a \sin i = 12.08 \pm .06 R_\odot \quad \text{and} \quad a = 12.17 \pm .06 R_\odot$$

$$M = (M_1 + M_2)/2 = 1.51 \pm .02 M_\odot$$

where  $a \sin i$  and  $M_1 \sin^3 i$  and  $M_2 \sin^3 i$  are given by Andersen *et al.* (1975). For the surface gravity we derive

$$\log g = 3.908 \pm .010$$

For the indices in the standard *uvby* system we get the following values to an accuracy better than 0.002 since three standard stars were observed

$$b - y = 0.309 \quad (\text{corresponding to } B - V = 0.47 \text{ and spectral type F6})$$

$$m_1 = 0.157, \quad \Delta m_1 = 0.017$$

$$c_1 = 0.474, \quad \Delta c_1 = 0.057.$$

The deviations  $\Delta m_1 = 0.017$  and  $\Delta c_1 = 0.057$  are found relative to the Hyades values assuming no reddening. For stars in the spectral range F5–G2 Nissen (1970) found a factor of 13 to convert  $\Delta m_1$  to metal content  $\left[\frac{\text{Fe}}{\text{H}}\right]$ . Adopting  $\left[\frac{\text{Fe}}{\text{H}}\right]_{\text{Hyades}} = 0.20 \pm .15$  for the metallicity of the Hyades compared to the Sun we have

$$\left[\frac{\text{Fe}}{\text{H}}\right]_{\text{RZ Cha}} = -0.02 \pm .15 \quad \text{or} \quad Z_{\text{RZ Cha}} \approx Z_\odot^{+0.008}_{-0.006}$$

$$= 0.020^{+0.008}_{-0.006}$$

where  $Z_{\text{RZ Cha}}$  is the fraction by mass of elements heavier than hydrogen and helium. RZ Cha is then a normal population I star with solar composition. The indicated errors in  $Z_{\text{RZ Cha}}$  are those found assuming  $\pm .15$  in  $\left[\frac{\text{Fe}}{\text{H}}\right]$ .

The value  $\Delta c_1 = +.057$  indicates that the star is evolved in agreement with the small value of the gravity found above.

Adopting the temperature scale by Morton and Adams (1968) we find the effective temperature which again gives the luminosity as we know the radius.

$$T_e = 6580^\circ, \quad \log T_e = 3.818,$$

$$\log \frac{L}{L_\odot} = 0.937, \quad M_{\text{bol}} = 2.41$$

(for  $M_{\text{bol}\odot} = 4.75$ ).

Comparing these data for RZ Cha to a new set of stellar models by Hejlesen (unpublished) we find that the value of  $\log g$  places the star above the gap at the end of the main-sequence phase and that the evolutionary phase is characterized by a thick hydrogen burning shell

Table 6. Changes in the helium content corresponding to changes in mass, content of heavy elements, effective temperature and radius

$\Delta M = +.03 M_\odot$	$\Delta Y = -.010$
$\Delta Z = +.008$	$\Delta Y = +.040$
$\Delta T_e = +150^\circ$	$\Delta Y = +.020$
$\Delta R = +.02 R_\odot$	$\Delta Y = \sim .000$

corresponding to subgiant stars. For the initial helium content we derive

$$Y_{\text{RZ Cha}} = 0.21 \pm .05$$

where the uncertainty is estimated from Table 6, giving the changes in helium content for the indicated changes in mass, content of heavier elements, effective temperature and radius. Particularly we notice the strong influence on the accuracy of  $Y$  from the uncertainty in the content of heavier elements.

The value of the helium content is very nearly the same as the value found for HS Hya, see Gyldenkerne *et al.* (1975). The average value of the two stars is somewhat lower than the mean value for stars of spectral type early B to F found previously by Popper *et al.* (1970).

*Acknowledgements.* The authors wish to thank Dr. B. Grønbech for his observing assistance on two nights, Drs. J. Andersen, Gjerløff and Imbert for supplying us with the spectroscopic results, and Mr. H. U. Nørgaard-Nielsen for his assistance at the computer.

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