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DETECTION OF INTERSTELLAR NH_3 IN THE FAR-INFRARED: WARM AND DENSE GAS IN ORION-KL

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ABSTRACT

We report the detection of the $(J, K) = a(4,3) \rightarrow s(3,3)$ rotation-inversion transition of ammonia at 124.6 μm toward the center of the Orion-KL region. The line is in emission and has a FWHM $\geq$ 30 km s^{-1}. The far-IR ammonia line emission probably comes mainly from the "hot core," a compact region of warm, very dense gas previously identified by the radio inversion lines of NH_3. The $a(4,3) \rightarrow s(3,3)$ line is very optically thick ($\tau \sim 10^3$), and since it is seen in emission, radiative excitation of the (4, 3) NH_3 level by far-IR emission from dust within the source can be ruled out. Radiative excitation via the 10 μm rovibrational transitions of NH_3 also seems unlikely. Hence, the (4, 3) level is probably collisionally excited and the gas in the hot core region is warmer than the dust. Since the far-IR line emission is highly trapped, densities of $\sim 10^7$ cm^{-3} are high enough to explain the observations. Shock heating by the mass outflow from IRc2 may account for the high gas temperatures in the hot core region.

Subject headings: infrared: sources — infrared: spectra — interstellar: molecules — nebulae: Orion Nebula

I. INTRODUCTION

Interstellar ammonia was first discovered by its inversion transitions at 1.3 cm (Cheung *et al.* 1968). Recent single-dish observations of several of the 1.2 cm inversion transitions toward the center of the Orion-KL region have shown the existence of a component of "hot" ammonia gas in addition to the "spike" NH_3 emission from the quiescent molecular cloud (cf. Barrett, Ho, and Myers 1977; Wilson, Downes, and Bieging 1979; Morris, Palmer, and Zuckerman 1980; Ziurys *et al.* 1981). High-resolution mapping with the VLA has shown that the hot core NH_3 emission comes from a region of angular diameter of about 10″ (7×10^{16} cm) and a temperature of 150–200 K. The opacities of many of the inversion lines are large ($\tau \geq 10$), and peak NH_3 column densities reach 5×10^{18} cm^{-2} (Genzel *et al.* 1982; Pauls *et al.* 1983; Palmer *et al.* 1983). The NH_3 hot core component may come from both streaming gas within and turbulent, swept-up gas in interaction with outflow from the luminous infrared source IRc2 (Downes *et al.* 1981). Morris *et al.* have shown that the radio data are consistent with pure collisional excitation of the NH_3 levels if densities are 10^9 cm^{-3}, or with a combination of far-IR radiative excitation and collisional excitation at densities $\sim 10^7$ cm^{-3} (assuming the NH_3 lines are optically thin). Emission lines from SO and SiO and from vibrationally excited HC_3N, CH_3CN and torsionally excited CH_3OH with characteristics similar to the NH_3 have also been found and probably also come from the hot core Orion-KL region (Plambeck *et al.* 1982; Wright *et al.* 1983; Clark *et al.* 1976; Loren *et al.* 1981; Goldsmith *et al.* 1982; Hollis *et al.* 1983). In the present *Letter*, we report the first detection of a far-IR emission line of NH_3. The new data are inconsistent with radiative excitation of the (4, 3) NH_3 level and support a collisional model. A hydrogen density of $\sim 10^7$ cm^{-3} is all that is required, as the far-IR NH_3 lines are very optically thick. Collisional excitation at that density may also explain the emission of some of the other molecular lines from the hot core region.

II. OBSERVATIONS AND RESULTS

The data were taken in 1982 February with the 91.4 cm telescope on board the NASA Kuiper Airborne Observatory. The spectrometer was a liquid-helium–cooled, tandem Fabry-Perot described by Storey, Watson, and Townes (1980), with a photoconductive detector. The angular resolution was 44″ (FWHM, or 55″ equivalent disk), and the chopper throw was 4′ at a chopping frequency of 29 Hz. The total system NEP (noise equivalent power) was 3×10^{-14} W $Hz^{-1/2}$. The

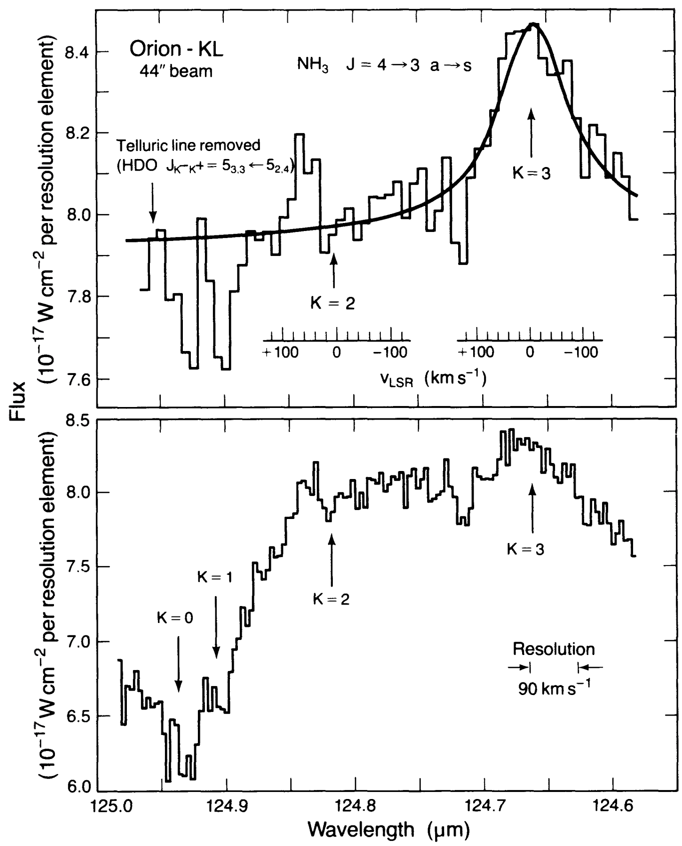

FIG. 1.—125 μm spectrum toward the core of the Orion-KL region, with a 44″ beam (FWHM). The spectral resolution is a Lorentzian of FWHM 90 ± 10 km s^{-1}. (*lower*) Observed spectrum, with arrows marking the positions of the $a(4, K) \rightarrow s(3, K)$ rotational inversion transitions of NH_3 toward Orion ($v_{LSR} = +5$ km s^{-1}). Also marked is a telluric HDO line on the left side of the spectrum. The $K = 0$ and $K = 1$ NH_3 lines are close to the bottom of this absorption feature. There is a second HDO absorption line at 124.3008 μm which may cause the downward slope on the right side of the spectrum. (*upper*) The same spectrum, but smoothed and with Lorentzians fitted to the HDO lines and then divided out. The positions of the $K = 2$ and $K = 3$ NH_3 lines are indicated, together with the appropriate velocity scales. A Lorentian fit to the $(4,3) \rightarrow (3,3)$ transition (*thin line*) gives an LSR velocity centroid of 0 ± 10 km s^{-1} and a FWHM of 140 ± 20 km s^{-1}, that is, significantly wider than the resolution. The $K = 2$ line is at least 3 times weaker than the $K = 3$ line. The positive bump at LSR + 70 km s^{-1} is probably narrower than our resolution and therefore is not interpreted as real.

$J = 4 \rightarrow 3$, $a \rightarrow s$[1] NH_3 rotation inversion transitions at 125 μm in Orion-KL were observed at a resolving power $\lambda/\Delta\lambda = 3400$, resulting in a Lorentzian instrumental profile of FWHM 90 km s^{-1}. Wavelength and velocity calibration was provided by the NH_3 $J = 4 \rightarrow 3$ lines in a gas cell and an HDO line at 124.9547 μm (McClatchey *et al.* 1973), with a precision of about ±10 km s^{-1}. To calibrate the line intensities near 125 μm, we used the line-to-continuum ratios and assumed a flux density of 6.0×10^4 Jy for the central 60″ of the Orion-KL region at 125 μm (Werner *et al.* 1976). The $J = 5 \rightarrow 4$, $a \rightarrow s$ transition region at 100 μm was observed at a resolving power of $\lambda/\Delta\lambda = 4000$, resulting in a velocity resolution (FWHM) of 75 km s^{-1}. Absolute wavelength calibration was provided by the NH_3 lines in a gas cell, and a telluric $H_2{}^{18}O$ line at 100.2601 μm. The observed 125 μm spectrum of Orion-KL is shown in Figure 1, and the main results are as follows:

1. The $(J, K) = a(4,3) \rightarrow s(3,3)$ line (rest wavelength, 124.6474 μm, Urban *et al.* 1981) is present in emission with a local standard of rest (LSR) velocity centroid of 0 ± 10 km s^{-1}, in agreement with the velocity centroids of the radio inversion lines (v_{LSR} = 5–8 km s^{-1}). The observed line width (FWHM) is 140 ± 20 km s^{-1}, which is significantly wider than the instrumental resolution (90 ± 10 km s^{-1}). Hence, the intrinsic velocity width of the line in Orion-KL has to be ≥ 30 km s^{-1} if the line profile is Lorentzian. For a Gaussian shape, the width would be substantially larger than this and may be due to high-velocity gas with too small a column density to have been detected in the inversion

[1]The symbols s and a refer to the symmetry of the rotation inversion wave function with respect to reflection about the plane of the hydrogens. An $a \rightarrow s$ transition is between the upper inversion level in the upper rotational state to the lower inversion level in the lower rotational state. Each inversion level is further split by hyperfine structure, which is not resolved in the present measurements.

spectrum. The peak line flux ($\sim 5 \times 10^{-18}$ W cm^{-2} within the passband of the spectrometer) is $\sim 6\%$ the intensity of the continuum at the resolution used. The effective line brightness temperature is 25–30 K, assuming that the NH_3 source fills the beam.

2. The neighboring $(J, K) = (4,2) \rightarrow (3,2)$ line at 124.7957 μm (Urban *et al.* 1981) is not evident. At $v_{LSR} = 5$ km s^{-1}, the 3 σ upper limit to the line intensity is 2×10^{-18} W cm^{-2}, that is, about 3 times weaker than the (4, 3) line. There is a feature at $v_{LSR} = +70$ km s^{-1} which, however, is not interpreted as significant, as its width is less than the instrumental resolution. The $(4,1) \rightarrow (3,1)$ and $(4,0) \rightarrow (3,0)$ transitions (rest wavelengths 124.8835 and 124.9125 μm, Urban *et al.*) are buried in the wings of a strong H_2O line.

3. None of the observed $J = 5 \rightarrow 4$ transitions falling within observed wavelengths ($K = 3$, 100.1046 μm; $K = 2$, 100.2129 μm; $K = 1$, 100.2772 μm, Urban *et al.* 1981) are present to a 3 σ limit greater than 5×10^{-18} W cm^{-2} (in absorption or emission). Only the $(5,3) \rightarrow (4,3)$ transition is not confused by telluric absorption. The $(5,4) \rightarrow (4,4)$ line wavelength was not observed since this line was known to be almost coincident with a strong telluric H_2O line.

It may not be surprising that the $(4,3) \rightarrow (3,3)$ line is stronger than the other transitions, since it is the only observed transition whose upper level has $K = J - 1$. Nonmetastable levels with $K \leq J - 2$ are expected to require substantially higher densities or a more intense radiation field to be populated equivalently to the $K = J - 1$ levels.

III. DISCUSSION

In the following, we discuss possible excitation mechanisms for the NH_3 gas and investigate which of the known gas components at the center of Orion-KL may account for the far-IR emission. Table 1 lists the characteristics of the different components. In addition to the hot core and spike seen in the NH_3 inversion lines, there are also the "plateau" and "shocked gas" components in other molecular lines which probably come from gas within and in interaction with the "high-velocity" outflow (e.g., Scoville 1981; Beckwith 1981).

a) *Radiative Excitation*

The detection of an emission line immediately rules out that far-IR continuum radiation by dust mixed with the gas or by a far-IR source embedded in the line-emitting region can *alone* account for the excitation of the (4, 3) level. Far-IR radiation by an *internal* source can only result in an absorption line or in a redistribution of the continuum radiation, but not in a *net* emission line. Resonant scattering by NH_3 molecules behind or beside a continuum source could conceivably produce the emission line, but there is no observational evidence for a strong far-IR continuum source *external* to the center of the region (Wynn-Williams *et al.* 1983). Radiative excitation is also possible via the rovibrational bands of NH_3, particularly the ν_2 bands at 10 μm. The luminous, compact source IRc2, for example, could efficiently pump the 10 μm transitions, since its spectrum peaks at about 8 μm. However, detailed considerations, based on present knowledge of the structure of this region, rule out such excitation. To pump the far-IR transitions, the total number of 10 μm photons in a given rovibrational line should be greater than or equal to the observed number of photons in the detected far-IR line ($1.5 \pm 0.5 \times 10^{47}$ s^{-1}). The number of 10 μm photons emitted by IRc2 per rovibrational transition is about $N_{IRc2} \approx 5 \times 10^{46} \Delta v_{10} (s^{-1})$, where Δv_{10} is the FWHM of a rovibrational line in units of 10 km s^{-1}. This value is an upper limit, for which one must assume that the luminosity of IRc2 is 10^5 $L_\odot$, that is, equal to the luminosity of the whole Orion-KL region. Furthermore, the number of 10

TABLE 1

FAR-INFRARED NH_3 LINE EMISSION AND ABSORPTION AND THE KINEMATIC COMPONENTS IN ORION-KL

Parameter	Spike	Hot Core	Plateau	Shocked Gas
Size of region	$\geq 1'$	$10''$	$40''$	$1'$
Velocity width (km s^{-1})	2.5 (FWHM)	10 (FWHM)	60 (FWHM)	60 (FWHM)
		30–40[a]	150[a]	120[a]
Temperature (K)	70	200 ± 50	100	1000
Hydrogen density (cm^{-3})	10^6	$\geq 10^7$	10^5–10^6	1–3×10^6
Hydrogen column density (cm^{-2})	5×10^{23}	5–20×10^{23}	1–5×10^{22}	3–10×10^{21}
Total mass ($M_\odot$)	100–200	1–10	1–10	0.5–3
$\tau_{NH_3}^{RADIO}$ (4, 3)	0.1	8 ± 2	...	...
$\tau_{NH_3}^{FIR}$ (4, 3)	2	600	$\sim 1 \times \left(\frac{N_{H_2}}{3 \times 10^{22}}\right)\left(\frac{\chi NH_3}{10^{-7}}\right)$	$\sim 0.1 \times \left(\frac{N_{H_2}}{5 \times 10^{21}}\right)\left(\frac{\chi NH_3}{10^{-7}}\right)$
Max. FIR photons s^{-1} [b]	$10^{46.5}$	$10^{48.3}$	$10^{44.7}$	10^{46}

[a] Full width at zero power.

[b] The maximum far-IR photons s^{-1} which can be emitted or absorbed is taken to be $N_{NH_3} n_H \sigma v \exp(-h\nu/kT)$, where N_{NH_3} is the number of NH_3 molecules in the $J = 3$, $K = 3$ state if $\chi NH_3 = 10^{-7}$; σv is assumed to be 10^{-10} cm^3 s^{-1}, and $h\nu$ corresponds to the rotational energy difference. This is an upper limit to the number of photons emitted if the line has substantial opacity.

μm photons coming from IRc2 has to be corrected for emission and absorption of dust along the path to the far-IR emitting region. Recent infrared observations suggest that there is a cavity of low dust density out to a radius of $\sim 4 \times 10^{16}$ cm from IRc2, which is surrounded by the dense, clumpy hot core region and a region which contains most of the quiescent and the high-velocity gas (Werner, Dinerstein, and Capps 1983; Wynn-Williams *et al.* 1983). The 10 μm dust opacities through these regions are substantial ($\tau \sim 3$ to > 10). Because of the presence of the large cavity around IRc2, the dust grains in the hot core, spike, plateau, and shocked regions absorb near-IR radiation from IRc2, but are too far from that source to reach temperatures so that they significantly emit at 10 μm. The dust in the far-IR emitting region, therefore, cannot contribute significantly to 10 μm pumping. Since the far-IR NH_3 lines are optically thick, the emission must come from the outer surface of the region, where the 10 μm radiation of IRc2 is attenuated by one to four orders of magnitude. Therefore, the number of 10 μm photons available in the far-IR emitting region, $N_{\mathrm{IRc2}} \times e^{-\tau}$, is probably much smaller than N_{FIR}. This probably rules out 10 μm radiative pumping.

b) Collisional Excitation: Far-Infrared Emission Comes from the "Hot Core"

We have used the values of temperature, hydrogen density, and total mass within the field of view listed in Table 1 to derive the maximum number of photons from $(4,3) \rightarrow (3,3)$ transitions for the different components of molecular gas in Orion-KL. For the hot core and spike, the column density of NH_3 can be directly estimated from the radio opacities. In other regions, the total number of NH_3 molecules is assumed to be approximately 10^{-7} that of H_2. Table 1 shows that only the hot core gas has sufficient density and numbers of NH_3-molecules to produce the observed $(4,3) \rightarrow (3,3)$ emission intensity. The spike region does not produce quite enough photons, and furthermore, its Doppler velocity width is much too small to provide the width found in the $(4,3) \rightarrow (3,3)$ transition. Other regions would produce far too few collisional excitations. We propose, therefore, that the NH_3 far-IR emission comes from the hot core region and that the (4,3) level is collisionally excited. As a consequence of the high opacity expected for the $(4,3) \rightarrow (3,3)$ transition from the hot core ($\tau_{\mathrm{FIR}} \sim 600$), the far-IR line also has to be broader than the 1.3 cm inversion lines, since amounts of gas at high velocity which are almost transparent at the inversion frequency can be observed in the rotational transition. This is consistent with the observed broadening of the 125 μm line. If it is assumed that all of the 125 μm line emission comes from a source 10″ in diameter, the radiation temperature would be about 200 K. This value is in good agreement with the brightness and rotational temperatures derived from the radio data. Estimates based on the total luminosity and the 8–30 μm continuum observations indicate that the dust temperature in this 10″ region is $\lesssim 150$ K. Hence, the gas temperature in the hot core is a few tens of degrees higher than the dust temperature.

The apparent absence of the $(4,2) \rightarrow (3,2)$ and $(5,3) \rightarrow (4,3)$ lines suggests that densities are not high enough to populate the (4,2) and (5,3) levels to the equivalent excitation temperature of the (4,3) transition. Microwave observations indicate that, indeed, the population of the (4,2) level is less than its thermal equilibrium value (Table 2). While this is an indication that the

TABLE 2

COLLISIONAL EXCITATION OF THE NH_3 LEVELS

Upper Level	τ^{RAD} [a]	b_u [b]	τ^{FIR}	$\beta(\tau^{\mathrm{FIR}})$ [c]	$n_{0.5}^{\mathrm{crit}}$ [d]	$\beta x n_{0.5}^{\mathrm{crit}}$ [e]	$\beta x n_{0.9}^{\mathrm{crit}}$ [f]
$(J, K) =$							
(4,3)	8 ± 2	0.9 ± 0.2	600	6×10^{-4}	10^9	6×10^5	10^7
(4,2)	0.7 ± 0.2	0.4 ± 0.3	100	3×10^{-3}	2×10^9	6×10^6	6×10^7
(7,6)	0.7 ± 0.2	0.3 ± 0.2	30	10^{-2}	5×10^9	5×10^7	5×10^8
(10,9)	0.15 ± 0.08	0.4 ± 0.1	3	10^{-1}	10^{10}	10^9	10^{10}

[a] The opacities of the 1.3 cm inversion lines are estimated from the single-dish observations by Morris, Palmer, and Zuckerman 1980 and Ziurys *et al.* 1981 and represent an average value across the source. For the (4,2), (7,6), and (10,9) lines, the opacities are obtained by comparing to the line temperature of the optically thick (4,3) line.

[b] The departure coefficients $b = n_{\mathrm{upper}}/n_{\mathrm{upper(LTE)}}$ are estimated from the observed opacities and the opacities extrapolated from $\tau(3,3) = 20$ at $T = 200$ K.

[c] The escape probability for far-IR line radiation: $\beta(\tau) = 1 - e^{-3\tau}/3\tau$. The value used is a compromise between an expanding source and Gaussian turbulence (de Jong, Chu, and Dalgarno 1975; de Jong, Dalgarno, and Boland 1980).

[d] The variable $n_{0.5}^{\mathrm{crit}}$ is the density where $b = \frac{1}{2}$ without trapping of far-IR radiation $[b = n_{\mathrm{upper}}/n_{\mathrm{upper(LTE)}} = (1 + A_{ul}\beta(\tau^{\mathrm{FIR}})/C_{ul})^{-1}$, where A_{ul} is the Einstein coefficient and C_{ul} the collisional rate between upper and lower levels].

[e] $\beta x n_{0.5}^{\mathrm{crit}}$ is the density where $b = \frac{1}{2}$ if trapping is taken into account.

[f] $\beta x n_{0.9}^{\mathrm{crit}}$ is the density where $b = 0.9$ if trapping is taken into account.

(4, 2) level is not in local thermodynamic equilibrium (LTE), it is the population ratio between (4, 2) and (3, 2) states to which the far-IR is sensitive, rather than the population itself, since the gas is optically thick at 125 μm. For the (4, 2) → (3, 2) transition to show no emission at all, or perhaps an absorption of the dust continuum, its effective temperature must be reduced to about 150 K. This implies that relative populations are $(n_{4,2}-n_{3,2})/(n_{4,3}-n_{3,3}) \lesssim 0.77$. Thus, assuming that excitation and radiative rates are comparable, a 23% change in relative populations between upper and lower states due to a smaller ratio between the collisional and radiative rates would explain the nondetection of emission from the (4, 2) → (3, 2) transition. Table 2 gives parameters for several nonmetastable NH_3 levels. The values of τ_{FIR} and $\beta(\tau_{FIR})$ are derived from the observed opacities of the 1.3 cm inversion lines. Values for collisional cross sections between He and NH_3 were taken from Green (1982) and were multiplied by 2.5 to account for the faster thermal speed of H_2 over He and for enhanced collisional cross sections of H_2 molecules which are in rotational states $J > 0$. The quantities $n_{0.5}^{crit}$ and $\beta \times n_{0.5}^{crit}$ listed in Table 2 are the densities at which the population is half of that at LTE without and with trapping corrections. $\beta \times n_{0.9}^{crit}$ is the density at which the population is 0.9 that at LTE, with trapping included. Table 2 shows that, because of the large trapping corrections, densities of $\sim 10^7$ cm^{-3} are sufficient to account for the population of the (4, 3) and (4, 2) levels. Higher densities may be necessary to account also for the observed radio brightness temperatures of higher excitation lines, such as (7, 6) and (10, 9). Such high densities, however, may be inconsistent with the weakness of the (4, 2) radio and infrared lines. A possible solution might be a mixture of radiative and collisional excitation.

At hydrogen densities of $\sim 10^7$ cm^{-3}, the NH_3 abundance in the hot core is $\sim 10^{-6}$ to 10^{-5}. A possible mechanism to heat the 1–10 $M_\odot$ of gas in the hot core region above the dust temperature may be the mechanical (shock) heating by the mass outflow from IRc2. The mechanical luminosity of this source is estimated to lie between 10 and 1000 $L_\odot$, which is sufficient to account for the total line cooling from ammonia and other molecules.

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